

Gravitational waves from cosmic bubble collisions

Dong-Hoon Kim^{1,2,a}, Bum-Hoon Lee^{4,5,b}, Wonwoo Lee^{4,c}, Jongmann Yang^{1,2,3,d}, Dong-han Yeom^{4,6,7,e}

¹ Basic Science Research Institute, Ewha Womans University, Seoul 120-750, Republic of Korea

² Institute for the Early Universe, Ewha Womans University, Seoul 120-750, Republic of Korea

³ Department of Physics, Ewha Womans University, Seoul 120-750, Republic of Korea

⁴ Center for Quantum Spacetime, Sogang University, Seoul 121-742, Republic of Korea

⁵ Department of Physics, Sogang University, Seoul 121-742, Republic of Korea

⁶ Yukawa Institute for Theoretical Physics, Kyoto University, Kyoto 606-8502, Japan

⁷ Leung Center for Cosmology and Particle Astrophysics, National Taiwan University, Taipei 10617, Taiwan

Received: 26 January 2015 / Accepted: 6 March 2015 / Published online: 20 March 2015

© The Author(s) 2015. This article is published with open access at Springerlink.com

Abstract Cosmic bubbles are nucleated through the quantum tunneling process. After nucleation they would expand and undergo collisions with each other. In this paper, we focus in particular on collisions of two equal-sized bubbles and compute gravitational waves emitted from the collisions. First, we study the mechanism of the collisions by means of a real scalar field and its quartic potential. Then, using this model, we compute gravitational waves from the collisions in a straightforward manner. In the quadrupole approximation, *time-domain* gravitational waveforms are directly obtained by integrating the energy-momentum tensors over the volume of the wave sources, where the energy-momentum tensors are expressed in terms of the scalar field, the local geometry and the potential. We present gravitational waveforms emitted during (i) the initial-to-intermediate stage of strong collisions and (ii) the final stage of weak collisions: the former is obtained numerically, in *full General Relativity* and the latter analytically, in the flat spacetime approximation. We gain *qualitative* insights into the time-domain gravitational waveforms from bubble collisions: during (i), the waveforms show the non-linearity of the collisions, characterized by a modulating frequency and cusp-like bumps, whereas during (ii), the waveforms exhibit the linearity of the collisions, featured by smooth monochromatic oscillations.

Contents

1 Introduction	1
2 Gravity–scalar field dynamics for colliding bubbles	2
3 Gravitational waves from bubble collisions	4
4 Conclusions	10
References	11

1 Introduction

A detection of signatures of primordial gravitational waves (GWs) in the cosmic microwave background was claimed by the BICEP2 experiment in 2014 [1]. But this was shown to be likely caused by interstellar dust soon thereafter [2], and the search for true signatures of primordial GWs is still ongoing. The detection of the GW signatures, if confirmed, would gain the greatest importance, among other things, from its link to cosmic “inflation”: primordial GWs are seen as the smoking gun for the “Big Bang” expansion. According to the inflation theory, the early universe experienced an extreme burst of expansion, which lasted a tiny fraction of a second, but smoothed out irregularities—inhomogeneities, anisotropies, and the curvature of space, and made the universe appear homogeneous and isotropic [3–5].

It has been suggested that inflationary models of the early universe most likely lead to a “multiverse” [6, 7]. One such model is “eternal inflation” [8]: it proposes that many bubbles of spacetime individually nucleate and grow inside an ever-expanding background multiverse. The nucleation and growth of such bubbles can be modeled by a Coleman–de Luccia (CDL) instanton, a type of quantum transition between two classically disconnected vacua at different energies; the higher energy (false vacuum), the lower energy (true

^a e-mail: ki13130@gmail.com

^b e-mail: bhl@sogang.ac.kr

^c e-mail: warrior@sogang.ac.kr

^d e-mail: jyang@ewha.ac.kr

^e e-mail: innocent.yeom@gmail.com

vacuum) [9–11]. A scalar field initially in the false vacuum state may tunnel quantum mechanically to the true vacuum state. This nucleates bubbles of the true vacuum (new phase) inside of the false vacuum (old phase) background; through a first-order phase transition. These bubbles then expand and collide with each other. The mechanism of bubble collisions can be effectively modeled by the CDL instanton: as bubbles continue to collide repeatedly, the scalar field transitions back and forth repeatedly between the false vacuum and the true vacuum, eventually settling down in the true vacuum as the collision process is gradually terminated.

From the viewpoints of physical cosmology, bubble collisions and GWs emitted from the collisions are interesting in the following contexts: (1) Our primordial inflation would be completed by a second-order (not by a first-order) phase transition. However, there is a possibility that some weaker inflation could occur after the primordial inflation; for example, “thermal inflation” [12]. It is quite probable that the thermal inflation is completed by a first-order phase transition, and therefore bubble collisions could take place through a CDL instanton. Then there would be some signatures of bubble collisions, which would presumably be carried via GWs [13]. (2) Suppose that we live in a single large true vacuum bubble and that the boundary of our bubble would collide with another bubble that is located outside our observable universe [14]. Then there would exist some signatures of bubble collisions and these could be carried via GWs. For scenario (1), the mechanism of bubble collision–GW emission should be modeled stochastically. However, for scenario (2), the mechanism can be well approximated by a two-bubble collision model.

There were numerous studies about bubble collisions and GWs emitted from the collisions. Among others, Hawking et al. [15] and Wu [16] studied the mechanism of the collision of two bubbles using the thin-wall approximation. Johnson et al. [17] and Hwang et al. [18] investigated the collision of two bubbles in full General Relativity via numerical computations. Kosowsky et al. [19] computed the GW spectrum resulting from two-bubble collisions in first-order phase transitions in flat spacetime using numerical simulations. Caprini et al. [20] developed a model for the bubble velocity power spectrum to calculate analytically the GW spectrum generated by two-bubble collisions in first-order phase transitions in flat spacetime.

In this paper, we focus on collisions of two equal-sized bubbles and compute GWs emitted from the collisions in *time domain*. Largely, our analysis proceeds in two steps through Sects. 2 and 3. In Sect. 2, we study the mechanism of bubble collisions by means of a real scalar field and a quartic potential of this field, building the simplest possible model for a CDL instanton. The Einstein equations and a scalar field equation are derived for this system and are solved simultaneously for the full General Relativistic treatment of the

collision dynamics. Hwang et al. [18] is closely reviewed for this purpose. In Sect. 3, using the scalar field model from Sect. 2, we compute GWs from the bubble collisions in a straightforward manner. In the quadrupole approximation, time-domain gravitational waveforms are directly obtained by integrating the energy-momentum tensors over the volume of the wave sources, where the energy-momentum tensors are expressed in terms of the scalar field, the local geometry and the potential; therefore, containing all necessary information about the bubble collisions. Part of the computational results from Ref. [18] are recycled here to build the energy-momentum tensors. In parallel with the scalar field solutions in Sect. 2, which have been obtained with various false vacuum field values [18], we present gravitational waveforms emitted during

- (i) the initial-to-intermediate stage of strong collisions, and
- (ii) the final stage of weak collisions.

The former is obtained numerically, in *full General Relativity* and the latter analytically, in the flat spacetime approximation. The thin-wall and quadrupole approximations are assumed to simplify our analysis and the next-to-leading order corrections beyond these approximations are disregarded. However, the approximations serve our purpose well: we aim to provide *qualitative* illustrations of the time-domain gravitational waveforms from the bubble collisions, which will be useful for constructing the templates for observation in the future. We adopt the unit convention, $c = G = 1$ for all our computations of GWs.

2 Gravity–scalar field dynamics for colliding bubbles

The mechanism of two equal-sized colliding bubbles can be effectively modeled by means of a CDL instanton [8–10]. Basically, one can build a model for this, which consists of gravitation, a real scalar field and a potential of the field. In this section we introduce one such model from Hwang et al. [17], which is built with a quartic potential, the simplest possible one for the CDL instanton.

2.1 Dynamics of bubble collisions

A system of Einstein gravity coupled with a scalar field that is governed by a potential can be described by the following action:

$$S = \int d^4x \sqrt{-g} \left[\frac{1}{16\pi} R - \frac{1}{2} \nabla_\mu \phi \nabla^\mu \phi - V(\phi) \right], \quad (1)$$

where R denotes the Ricci scalar, ϕ the scalar field, and $V(\phi)$ the potential of the scalar field [18]. From this system the Einstein equations are derived:

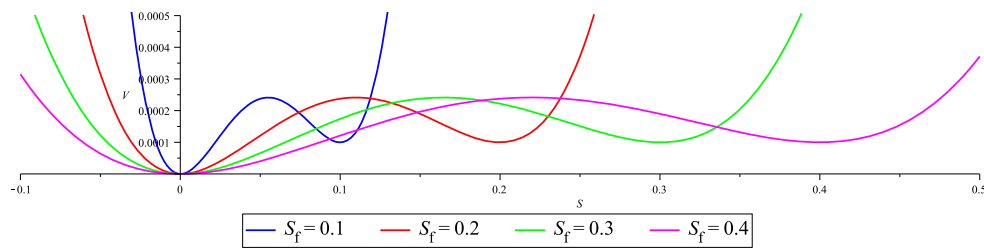


Fig. 1 The quartic potential $V(S) = \frac{3V_f}{\beta S_f^4} S^4 - \frac{2(\beta+3)V_f}{\beta S_f^3} S^3 + \frac{3(\beta+1)V_f}{\beta S_f^2} S^2$; expressed in terms of the false vacuum field $S_f = \sqrt{4\pi}\phi_f$, the vacuum energy of the false vacuum V_f and a free parameter β .

The patterns of the potential are shown for $\beta = 0.1$, $V_f = 10^{-4}$ and $S_f = \sqrt{4\pi}\phi_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$ (credit: Hwang et al. [18])

$$R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} = 8\pi T_{\mu\nu}, \quad (2)$$

where the energy-momentum tensors on the right-hand side are written as

$$T_{\mu\nu} = \phi_{;\mu}\phi_{;\nu} - \frac{1}{2}\phi_{;\rho}\phi_{;\sigma}g^{\rho\sigma}g_{\mu\nu} - V(\phi)g_{\mu\nu}. \quad (3)$$

Also, the scalar field equation for the system reads

$$\nabla^2\phi = \frac{dV}{d\phi}. \quad (4)$$

The Einstein equations (2) and the scalar field equation (4) constitute a scalar field model that effectively describes the mechanism of two colliding bubbles in curved space-time [18]. Given a potential $V(\phi)$, the scalar field solution ϕ and the geometry solution $g_{\mu\nu}$ should be obtained by solving Eqs. (2) and (4) simultaneously.¹ To this end, we prescribe an ansatz for the geometry $g_{\mu\nu}$ with the hyperbolic symmetry, using the *double-null* coordinates:

$$ds^2 = -\alpha_h^2(u, v)dudv + r_h^2(u, v)dH^2, \quad (5)$$

where $dH^2 = d\chi^2 + \sinh^2\chi d\theta^2$ with $0 \leq \chi < \infty$, $0 \leq \theta < 2\pi$ [18], and $\alpha_h(u, v)$ and $r_h(u, v)$ are to be determined by solving Eqs. (2) and (4) simultaneously in the coordinates (u, v, χ, θ) . In the flat spacetime limit, the double-null coordinates are defined as $u \equiv \tau - x$ and $v \equiv \tau + x$ with $\tau^2 \equiv t^2 - y^2 - z^2$, $t = \tau \cosh \chi$, $y = \tau \sinh \chi \sin \theta$, $z = \tau \sinh \chi \cos \theta$: in our analysis, the x -axis of Cartesian coordinates is chosen to coincide with a line adjoining the centers of the two bubbles, and the y -axis and the z -axis lie in a plane perpendicular to the x -axis.

2.2 Solving the scalar field equation

To build the simplest model of the CDL instanton for two identical colliding bubbles, we consider the potential in Sect. 2.1 to be

¹ This inevitably results in the effects of radiation reaction being included in the solutions, ϕ and $g_{\mu\nu}$ [18].

$$V(\phi) = \frac{p_4}{4!}\phi^4 + \frac{p_3}{3!}\phi^3 + \frac{p_2}{2!}\phi^2, \quad (6)$$

where p_2 , p_3 , and p_4 are constants which can be appropriately chosen to tune the shape of the potential. Bubble collisions are represented by the scalar field moving along this potential: the field initially in the false vacuum state (at higher local minimum of potential) tunnels quantum mechanically to the true vacuum state (at lower local minimum of potential), repeating the transitions back and forth between the two states, eventually settling down in the true vacuum state.

Following Ref. [18], we may rescale the scalar field, $S \equiv \sqrt{4\pi}\phi$ for computational convenience, and can specify p_2 , p_3 , and p_4 in terms of the false vacuum field $S_f = \sqrt{4\pi}\phi_f$, the vacuum energy of the false vacuum V_f and a free parameter β . The potential (6) can then be rewritten as

$$V(S) = \frac{3V_f}{\beta S_f^4} S^4 - \frac{2(\beta+3)V_f}{\beta S_f^3} S^3 + \frac{3(\beta+1)V_f}{\beta S_f^2} S^2. \quad (7)$$

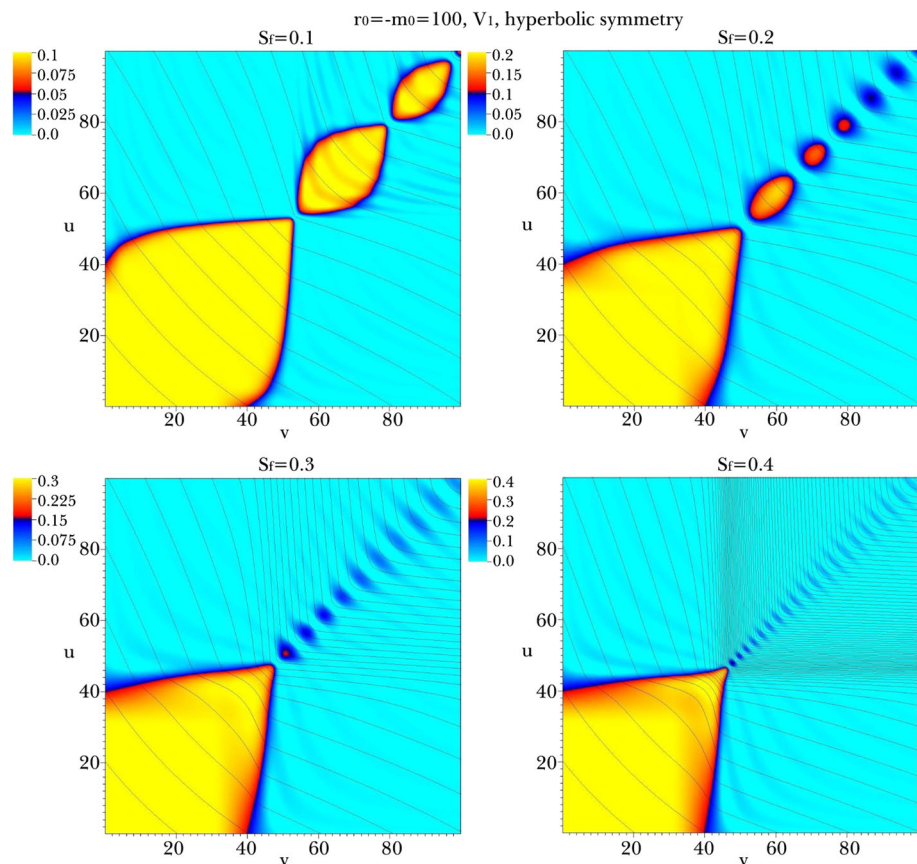
With this potential the scalar field equation (4), which is now rescaled, reads

$$\nabla^2 S = \frac{12V_f}{\beta S_f^4} S^3 - \frac{6(\beta+3)V_f}{\beta S_f^3} S^2 + \frac{6(\beta+1)V_f}{\beta S_f^2} S. \quad (8)$$

This is a non-linear wave equation whose analytical solution is not generally known: we normally approach this type of problem with numerical methods.

Now, we solve the scalar field equation (8) simultaneously with the Einstein equations (2), using the ansatz given by (5), in the coordinates (u, v, χ, θ) . However, it turns out that our scalar field solution is independent of the coordinates χ and θ and is expressed in the coordinates (u, v) only; namely, $S(u, v) = \sqrt{4\pi}\phi(u, v)$ [15, 18, 19]. With the choice of the constants, $\beta = 0.1$, $V_f = 10^{-4}$ and $S_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$ in Eq. (7), the potential takes the forms as given by Fig. 1 [18]. With this potential, our numerical solution $S(u, v)$ is obtained as presented by Fig. 2 [18]. In each case of S_f , (1)–(4), the bubble wall has a different value of tension due to a different value of S_f as shown by Fig. 1. In the top left of Fig. 2 the bubble has the lowest tension while in the bottom right it has the highest tension among the four cases

Fig. 2 The numerical solutions $S(u, v) = \sqrt{4\pi}\phi(u, v)$ obtained with the potential from Fig. 1; with various false vacuum field values, $S_f = \sqrt{4\pi}\phi_f =$ (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4 (credit: Hwang et al. [18])



of S_f . This results in the wall crossing regions in the top left being relatively wider than those in the bottom right.

3 Gravitational waves from bubble collisions

In Sect. 2 we have built a system of two equal-sized colliding bubbles in curved spacetime by means of a CDL instanton model, considering the potential given by Eq. (7) [18]. Now, we consider that the a present observer lives in the true vacuum region of one of the two bubbles and that signatures of bubble collisions which took place in the distant past are being carried to the present observer via GWs. Here we assume that the distance between the center of collision region and the observer can be arbitrarily large (within the size of our universe), and thus that the time for a collision event, which is the retarded time to the present observer, can be quite far in the distant past; namely, $t_R = t_P - r/c \ll t_P$, where t_R denotes the retarded time, t_P the present time and r the distance. Therefore, our bubble collision may be regarded as a localized event, as long as the observer is reasonably far away from the collision region. In Fig. 3 a causal relationship between two colliding bubbles and an observer is depicted using a null-cone. Here collision events that took place in the distant past are placed within the intersection of the timelike

zone (green-colored region) of a null-cone (green dashed lines) and the ‘diamond’ zone (region enclosed by black dashed lines): all the collision events as our GW sources, namely the collisions in Fig. 2 (strong collisions in the initial-to-intermediate stage) and the collisions to be discussed in Sect. 3.2 later (weak collisions in the final stage) should be considered to have taken place within this intersection.²

In the above scenario, our GWs from bubble collisions can be computed in a straightforward manner: by integrating the energy-momentum tensors combined with a Green function over the volume of the wave sources, where the energy-momentum tensors are expressed in terms of the scalar field, the local geometry and the potential by means of Eqs. (3), (5), and (8); therefore, they contain all necessary information

² In principle, the diamond zone can be extended to cover the longer evolution of bubble collision. This will result in the larger intersection area with the timelike zone of the null-cone. However, no matter how large the intersection area is, it should still be regarded as a well-localized region for our GW sources: we assume that a present observer is reasonably far away from the sources in our computations of GWs. This naturally renders our results convergent. While we consider only the retarded field for our ‘time-domain’ GWs, Kosowsky et al. [19] take both the retarded and the advanced fields for their ‘frequency-domain’ GWs. On account of this, their computation domain is unbounded, but they obtain convergent results using a method of ‘phenomenological cutoff’.

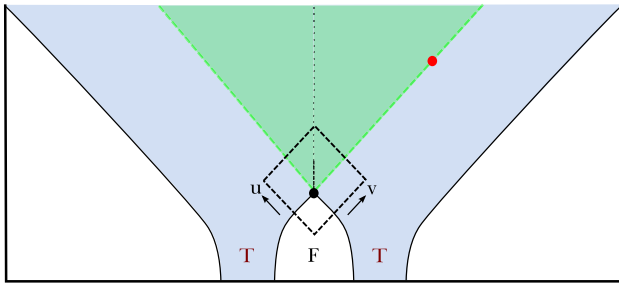


Fig. 3 A causal relationship between two colliding bubbles and an observer: gravitational waves from bubble collisions can be observed by a distant observer (*red spot*) who is in the causal future of a collision event (*black spot*). The *dotted vertical line* inside the timelike zone (*green-colored region*) of a null-cone (*green dashed lines*) represents the time-axis, along which bubble collisions took place in the distant past; within the intersection of the timelike zone of the null-cone and the ‘diamond’ zone (region enclosed by *black dashed lines*). *u* and *v* represent the double-null coordinates as defined in Sect. 2. *T* and *F* denote a true vacuum and a false vacuum, respectively

about the bubble collisions. A mathematical description of this computation is given as follows. In the transverse trace-free (TT) gauge, GWs, as derived from the perturbed Einstein equations in linearized gravity, can be expressed as

$$h_{ij}^{\text{TT}}(t, \mathbf{x}) = \bar{h}_{ij}^{\text{TT}}(t, \mathbf{x}) = \frac{4G}{c^4} \Lambda_{ij,kl}(\mathbf{n}) \times \int d^3x' \frac{1}{|\mathbf{x} - \mathbf{x}'|} T^{kl} \left(t - \frac{|\mathbf{x} - \mathbf{x}'|}{c}, \mathbf{x}' \right), \quad (9)$$

where $\bar{h}_{ij} \equiv h_{ij} - \frac{1}{2} \delta_{ij} h_k^k$ and the unit vector \mathbf{n} denotes the propagation direction of the waves, and the projection tensor for gravitational radiation,

$$\Lambda_{ij,kl}(\mathbf{n}) \equiv P_{ik} P_{jl} - \frac{1}{2} P_{ij} P_{kl} \quad (10)$$

with

$$P_{ij} = \delta_{ij} - n_i n_j. \quad (11)$$

We find that the computation would be technically quite difficult with the integral as it is in Eq. (9): the way the source point \mathbf{x}' (integration variable) is combined with the field point \mathbf{x} in the integrand would make our calculation quite intractable. However, expressing the integral in expansion, we obtain a more computationally favorable form:

$$h_{ij}^{\text{TT}}(t, \mathbf{x}) = \frac{4G}{c^4 r} \Lambda_{ij,kl}(\mathbf{n}) \left[\int d^3x' T^{kl}(t_R, \mathbf{x}') + \frac{1}{c} n^m \frac{d}{dt_R} \int d^3x' T^{kl}(t_R, \mathbf{x}') x'_m + \frac{1}{2c^2} n^m n^p \frac{d^2}{dt_R^2} \int d^3x' T^{kl}(t_R, \mathbf{x}') x'_m x'_p + \dots \right]_{t_R=t-r/c}, \quad (12)$$

where $r = |\mathbf{x}|$ and $t_R = t - r/c$ denotes the retarded time. In particular, the computation resulting from the first term alone

in the square bracket in Eq. (12) is called the “quadrupole approximation”. The next terms will provide corrections to this computation.

3.1 Computation of gravitational waves in the quadrupole approximation

The complete information about the motion of the colliding two-bubble system is encoded in the scalar field solution $S = \sqrt{4\pi\phi}$, as given by Fig. 2, and thus is carried by the energy-momentum tensors through Eq. (3): to be precise, the energy-momentum tensors are comprised of the scalar field S and the geometry $g_{\mu\nu}$, which are obtained by solving Eqs. (2) and (4) simultaneously [18]. As described by Eqs. (9) and (12), GWs from the system are computed with the energy-momentum tensors being the sources. It is believed that the two bubbles will be in highly relativistic motion when they collide [15]. In view of this, corrections due to the next-to-leading order terms in Eq. (12) should not be disregarded if one aims to compute GWs from the system accurately. However, although not perfectly accurate, the leading order term alone in Eq. (12) provides the “quadrupole approximation” of GWs:

$${}_Q h_{ij}^{\text{TT}}(t, \mathbf{x}) = \frac{4}{r} \Lambda_{ij,kl}(\mathbf{n}) I_{kl}(t_R), \quad (13)$$

where we have adopted the unit convention $c = G = 1$, and

$$I_{kl}(t_R) \equiv \int d^3x' T^{kl}(t_R, \mathbf{x}'). \quad (14)$$

Throughout this paper our computation is carried out only based on this piece. Our main purpose is to provide *qualitative* insights into patterns of GWs from the colliding two-bubble system in time domain, and the next-to-leading order corrections in Eq. (12) are disregarded in our analysis.

Following Ref. [19], we can reduce the amount of computation in a great deal. As described in Sect. 2.1, we choose the x -axis to coincide with the line adjoining the centers of the two bubbles. With the axial symmetry about the x -axis, the off-diagonal components are zero and we can put I_{kl} in the form

$$I_{kl} = D \delta_{kl} + \Delta \delta_{kx} \delta_{lx}. \quad (15)$$

Here, the first term turns out to be

$$D = \frac{1}{2} (I_{yy} + I_{zz}), \quad (16)$$

which does not contribute to gravitational radiation due to Eqs. (10) and (11). The second term is given by

$$\Delta = I_{xx} - \frac{1}{2} (I_{yy} + I_{zz}). \quad (17)$$

Therefore, I_{kl} is practically equivalent to $\Delta\delta_{kx}\delta_{lx}$:

$$I_{kl} \sim \delta_{kx}\delta_{lx} \left[I_{xx} - \frac{1}{2}(I_{yy} + I_{zz}) \right]. \quad (18)$$

Then by Eqs. (14) and (18) we may express

$$I_{kl}(t_R) = \delta_{kx}\delta_{lx} \int d^3x' \times \left[T^{xx}(t_R, \mathbf{x}') - \frac{1}{2}[T^{yy}(t_R, \mathbf{x}') + T^{zz}(t_R, \mathbf{x}')] \right]. \quad (19)$$

Now, recall from Sect. 2.1 that in the flat spacetime we define the hyperbolic coordinates τ, χ, θ by

$$t_R = \tau \cosh \chi, \quad (20)$$

$$y = \tau \sinh \chi \sin \theta, \quad (21)$$

$$z = \tau \sinh \chi \cos \theta, \quad (22)$$

so that

$$\tau^2 = t_R^2 - \rho^2, \quad (23)$$

where $0 \leq \chi < \infty$, $0 \leq \theta < 2\pi$ and $\rho \equiv \sqrt{y^2 + z^2}$. In these coordinates the flat spacetime metric takes the form

$$ds^2 = -d\tau^2 + dx^2 + \tau^2(d\chi^2 + \sinh^2 \chi d\theta^2). \quad (24)$$

In this geometry, however, the scalar field solution is independent of the coordinates χ and θ and is expressed in the coordinates (τ, x) only; namely, $\phi(\tau, x)$ [15, 19]. Taking this into account, we should rewrite the volume element for the integral (19) as

$$d^3x = dx\rho d\rho d\theta = -dx\tau d\tau d\theta, \quad (25)$$

which is defined at the instant t_R by means of Eq. (23).

From Eq. (23) we see that τ has an upper bound t_R with $\rho = 0$. This represents the exterior surface of the bubble walls, i.e.

$$\tau = t_R. \quad (26)$$

However, the interior surface is found from Eq. (23) to be

$$\tau = t_R - \frac{\eta^2}{8t_R} + \mathcal{O}\left(\frac{\eta^4}{t_R^3}\right), \quad (27)$$

given a wall thickness $\eta \sim 2\rho \ll t_R$. Then from Eqs. (19), (25), (26), and (27) we can compute $I_{kl}(t_R)$ effectively out of a volume piece \mathcal{V} :

$$\begin{aligned} I_{kl}(t_R) &= \delta_{kx}\delta_{lx} \int_{\mathcal{V}} d^3x' \left[T_{xx} - \frac{1}{2}(T_{yy} + T_{zz}) \right] \\ &= 2\pi\delta_{kx}\delta_{lx} \int_{t_R - \eta^2/(8t_R)}^{t_R} d\tau' \tau' \int_{-x_0}^{x_0} dx' \\ &\quad \times \left[T_{xx} - \frac{1}{2}(T_{yy} + T_{zz}) \right] + \mathcal{O}\left(\frac{\eta^4}{t_R^3}\right) \end{aligned}$$

$$\begin{aligned} &= \frac{\pi}{4}\delta_{kx}\delta_{lx}\eta^2 \int_{-x_0}^{x_0} dx' \left[T_{xx} - \frac{1}{2}(T_{yy} + T_{zz}) \right]_{\tau=t_R} \\ &\quad + \mathcal{O}\left(\frac{\eta^4}{t_R^3}\right), \end{aligned} \quad (28)$$

where the volume piece \mathcal{V} is defined from a thin cylindrical shell in motion, having the thickness $\eta^2/(8t_R)$, extending along the x -axis, by means of Eqs. (26) and (27): from $t_R - \eta^2/(8t_R) + \mathcal{O}(\eta^4/t_R^3) \leq \tau \leq t_R$ we find $\mathcal{V} = \Delta x |\tau \Delta \tau| \Delta \theta = \frac{\pi}{4}\eta^2 \Delta x [1 + \mathcal{O}(\eta^2/t_R^2)]$, and the limit of the integral $x_0 = \Delta x/2$ should be chosen to be sufficiently large such that collision effects be fully covered in numerical integration.³ Following Ref. [15], we estimate a bubble wall thickness η , assuming that the walls will be highly relativistic when they collide, having the Lorentz factor γ :

$$\eta \sim \frac{3}{2}\phi_f/(\xi^2\gamma) \sim \frac{3}{2}\phi_f^2/(b\epsilon^4), \quad (29)$$

where ϕ_f denotes the scalar field value at the false vacuum and ξ^4 the effective height of the potential barrier between the two minima, and the Lorentz factor $\gamma = b\epsilon^4/(\xi^2\phi_f)$ with $2b$ representing the separation of the bubbles and ϵ^4 the potential difference between the two minima (which is equivalent to V_f in our analysis in Sect. 2.2) [15].

However, as described in Sect. 2.2, our scalar field $S = \sqrt{4\pi}\phi$ is obtained by solving Eqs. (2) and (8) simultaneously, using the ansatz (5), in the coordinates (u, v, χ, θ) . Then by Eq. (3) the energy-momentum tensors should be expressed in the same coordinates. Now, due to the definitions of u and v in the flat spacetime limit, and by Eqs. (20), (21), and (22) we have

$$u = \tau - x = \sqrt{t_R^2 - (y^2 + z^2)} - x, \quad (30)$$

$$v = \tau + x = \sqrt{t_R^2 - (y^2 + z^2)} + x, \quad (31)$$

$$\chi = \tanh^{-1} \left(\sqrt{\frac{y^2 + z^2}{t_R^2}} \right). \quad (32)$$

Using these relations, we find

$$\begin{aligned} T_{xx} &= T_{uu} \left(\frac{\partial u}{\partial x} \right)^2 + 2T_{uv} \frac{\partial u}{\partial x} \frac{\partial v}{\partial x} + T_{vv} \left(\frac{\partial v}{\partial x} \right)^2 \\ &\quad + T_{\chi\chi} \left(\frac{\partial \chi}{\partial x} \right)^2 = T_{uu} - 2T_{uv} + T_{vv}, \end{aligned} \quad (33)$$

³ By Eqs. (23) and (25) the volume piece can also be viewed as $\mathcal{V} = \Delta x \rho \Delta \rho \Delta \theta = \frac{\pi}{4}\eta^2 \Delta x [1 + \mathcal{O}(\eta^2/t_R^2)]$. Then it may be stated that the volume integral in Eq. (28) will be equivalently evaluated out of this volume piece, whose shape is a long thin cylinder with the diameter (thickness) η surrounding the x -axis. This is in agreement with the statement from Ref. [15]: “The kinetic energy of the bubble walls will be concentrated in a small region around the x -axis of a wall thickness η ”.

$$\begin{aligned}
T_{yy} &= T_{uu} \left(\frac{\partial u}{\partial y} \right)^2 + 2T_{uv} \frac{\partial u}{\partial y} \frac{\partial v}{\partial y} \\
&\quad + T_{vv} \left(\frac{\partial v}{\partial y} \right)^2 + T_{\chi\chi} \left(\frac{\partial \chi}{\partial y} \right)^2 \\
&= \frac{y^2}{t_R^2 - (y^2 + z^2)} (T_{uu} + 2T_{uv} + T_{vv}) \\
&\quad + \frac{t_R^2 y^2}{[t_R^2 - (y^2 + z^2)]^2 (y^2 + z^2)} T_{\chi\chi}, \quad (34)
\end{aligned}$$

$$\begin{aligned}
T_{zz} &= T_{uu} \left(\frac{\partial u}{\partial z} \right)^2 + 2T_{uv} \frac{\partial u}{\partial z} \frac{\partial v}{\partial z} + T_{vv} \left(\frac{\partial v}{\partial z} \right)^2 \\
&\quad + T_{\chi\chi} \left(\frac{\partial \chi}{\partial z} \right)^2 \\
&= \frac{z^2}{t_R^2 - (y^2 + z^2)} (T_{uu} + 2T_{uv} + T_{vv}) \\
&\quad + \frac{t_R^2 z^2}{[t_R^2 - (y^2 + z^2)]^2 (y^2 + z^2)} T_{\chi\chi}. \quad (35)
\end{aligned}$$

Substituting Eqs. (33), (34), and (35) into Eq. (28), we obtain

$$\begin{aligned}
I_{kl}(t_R) &= \frac{\pi}{4} \delta_{kx} \delta_{lx} \eta^2 \int_{-x_0}^{x_0} dx' \\
&\quad \times \left[T_{uu} - 2T_{uv} + T_{vv} - \frac{1}{2t_R^2} T_{\chi\chi} \right]_{t_R} + \mathcal{O} \left(\frac{\eta^4}{t_R^3} \right), \quad (36)
\end{aligned}$$

where the subscript t_R outside the square bracket means that the double-null coordinates (u, v) are defined at $\tau = t_R$; namely, $u = t_R - x$ and $v = t_R + x$. In the actual computation of Eq. (36), we integrate $T_{uu}(u, v)$, $T_{uv}(u, v)$, $T_{vv}(u, v)$ and $T_{\chi\chi}(u, v)$, which are constructed out of the scalar field solution $S(u, v) = \sqrt{4\pi} \phi(u, v)$, the geometry solution $g_{uv}(u, v)$, $g_{\chi\chi}(u, v)$, $g_{\theta\theta}(u, v)$ and the potential $V(S)$ via Eq. (3). Then we need to change the variable of integration, from x to u or v . Using the relations $u = t_R - x$ and $v = t_R + x$, we can convert

$$dx = -du \quad \text{or} \quad dx = dv. \quad (37)$$

Then we may rewrite

$$\begin{aligned}
\int_{-x_0}^{x_0} dx' T_{ab}(u, v) &= \int_{t_R - x_0}^{t_R + x_0} du T_{ab}(u, 2t_R - u) \\
&= \int_{t_R - x_0}^{t_R + x_0} dv T_{ab}(2t_R - v, v) \\
&= 2 \int_0^{x_0} du T_{ab}(t_R + u, t_R - u) \\
&= 2 \int_0^{x_0} dv T_{ab}(t_R - v, t_R + v), \quad (38)
\end{aligned}$$

where T_{ab} represents any of T_{uu} , T_{uv} , T_{vv} and $T_{\chi\chi}$, and the expressions in the second line have been obtained via trans-

lations, $u \rightarrow u - t_R$ and $v \rightarrow v - t_R$. Then by Eqs. (36) and (38) $I_{kl}(t_R)$ can be expressed as

$$\begin{aligned}
I_{kl}(t_R) &= \frac{\pi}{2} \delta_{kx} \delta_{lx} \eta^2 \left[\int_0^{x_0} du T_{uu}(t_R \pm u, t_R \mp u) \right. \\
&\quad - 2 \int_0^{x_0} du T_{uv}(t_R \pm u, t_R \mp u) \\
&\quad + \int_0^{x_0} du T_{vv}(t_R \pm u, t_R \mp u) \\
&\quad \left. - \frac{1}{2t_R^2} \int_0^{x_0} du T_{\chi\chi}(t_R \pm u, t_R \mp u) \right] + \mathcal{O} \left(\frac{\eta^4}{t_R^3} \right). \quad (39)
\end{aligned}$$

If the wall thickness η can be taken sufficiently small in Eq. (39), then by Eq. (13) we can compute the bubble-collision-induced GWs in the quadrupole approximation as

$$\begin{aligned}
Qh_{ij}^{\text{TT}}(t, \mathbf{x}) &\approx \frac{2\pi}{r} \eta^2 \Lambda_{ij,kl}(\mathbf{n}) \delta_{kx} \delta_{lx} \\
&\quad \times \left[\int_0^{x_0} du T_{uu}(t_R \pm u, t_R \mp u) \right. \\
&\quad - 2 \int_0^{x_0} du T_{uv}(t_R \pm u, t_R \mp u) \\
&\quad + \int_0^{x_0} du T_{vv}(t_R \pm u, t_R \mp u) \\
&\quad \left. - \frac{1}{2t_R^2} \int_0^{x_0} du T_{\chi\chi}(t_R \pm u, t_R \mp u) \right]. \quad (40)
\end{aligned}$$

Now, without loss of generality we may choose

$$\mathbf{n} = (n_x, n_y, n_z) = (\cos \vartheta, \sin \vartheta, 0), \quad (41)$$

where ϑ denotes the angle of propagation taken from the x -axis. From this it follows that

$$\begin{aligned}
\Lambda_{ij,kl}(\mathbf{n}) \delta_{kx} \delta_{lx} &= \delta_{ix} \delta_{jx} - 2\delta_{ix} n_j \cos \vartheta \\
&\quad + \frac{1}{2} n_i n_j (1 + \cos^2 \vartheta) - \frac{1}{2} \delta_{ij} \sin^2 \vartheta, \quad (42)
\end{aligned}$$

due to Eqs. (10) and (11). Substituting this into Eq. (40), we finally find the expression

$$\begin{aligned}
Qh_{ij}^{\text{TT}}(t, \mathbf{x}) &\approx \frac{2\pi}{r} \eta^2 \left[\delta_{ix} \delta_{jx} - 2\delta_{ix} n_j \cos \vartheta \right. \\
&\quad \left. + \frac{1}{2} n_i n_j (1 + \cos^2 \vartheta) - \frac{1}{2} \delta_{ij} \sin^2 \vartheta \right] \\
&\quad \times \left[\int_0^{x_0} du T_{uu}(t_R \pm u, t_R \mp u) \right. \\
&\quad \left. - 2 \int_0^{x_0} du T_{uv}(t_R \pm u, t_R \mp u) \right]
\end{aligned}$$

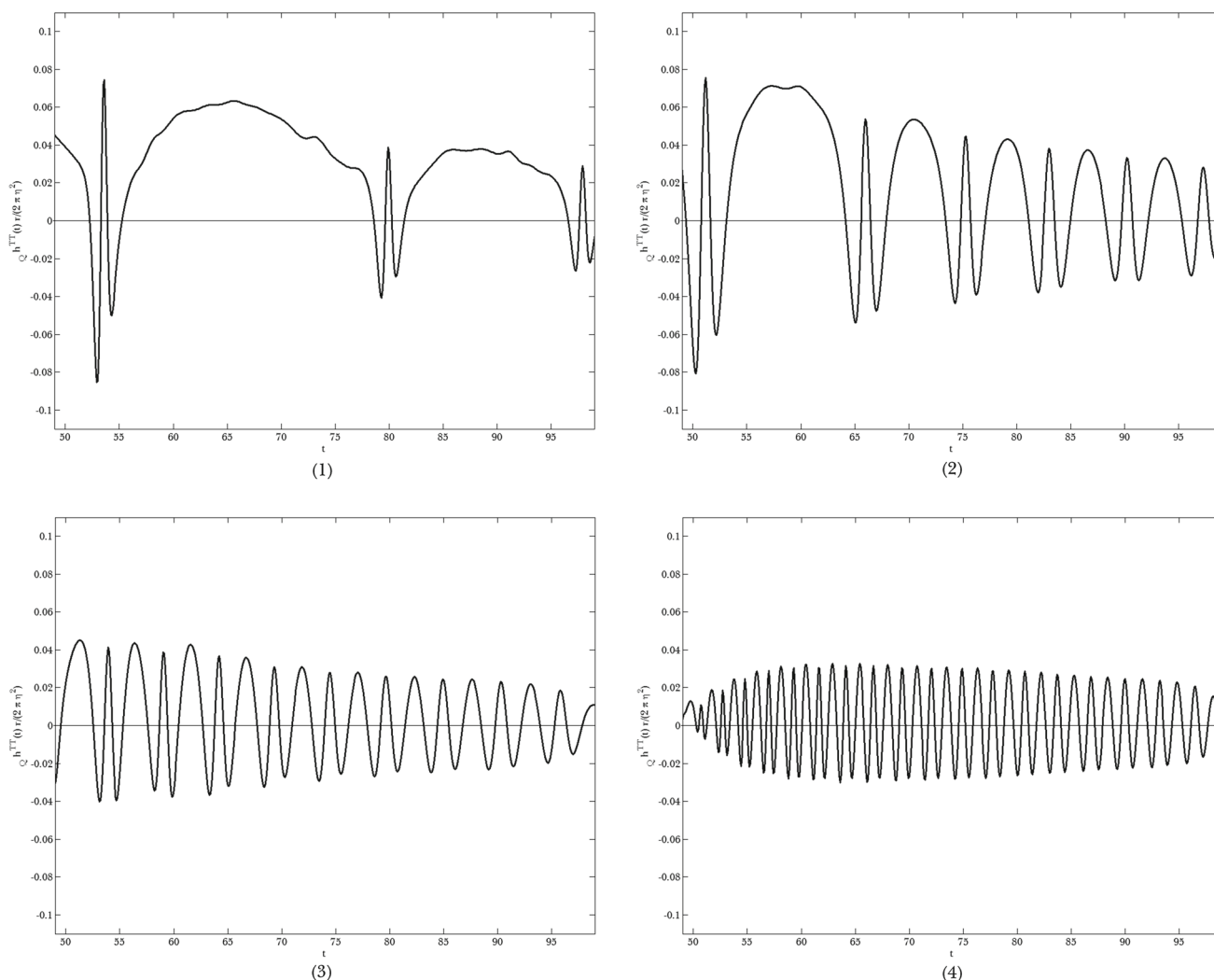


Fig. 4 The numerical plots of $Qh^{\text{TT}}(t)r/(2\pi\eta^2)$ with various false vacuum field values, $S_f = \sqrt{4\pi}\phi_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$. As $\eta^2 \sim \phi_f^4$, the amplitude of $Qh^{\text{TT}}(t)$ should scale as (1) 1, (2) 2^4 , (3) 3^4 , (4) 4^4

$$+ \int_0^{x_0} du T_{vv}(t_R \pm u, t_R \mp u) - \frac{1}{2t_R^2} \int_0^{x_0} du T_{\chi\chi}(t_R \pm u, t_R \mp u) \Big], \quad (43)$$

where $t_R = t - r = t - |\mathbf{x}|$, and the wall thickness η can be specified by means of (29); namely, in terms of the quantities for the bubble collision profiles, such as the false vacuum field ϕ_f (equivalent to $S_f/\sqrt{4\pi}$), the potential difference between the two minima ϵ^4 (equivalent to V_f) and half the separation of the bubbles b [15].

Result 1: The numerical computations of (43) are presented in Fig. 4; with various false vacuum field values, $S_f = \sqrt{4\pi}\phi_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$, in accordance with the scalar field solutions as presented by Fig. 2. Due to Eqs. (29) and (43), the amplitude of our GWs $Qh^{\text{TT}}(t)$ scales as ϕ_f^4 if the other conditions, ϵ^4 and b are kept the same. Thus, with $S_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$, the amplitude scales

as (1) 1, (2) 2^4 , (3) 3^4 , (4) 4^4 . The frequency of the waves is modulating due to the non-linearity of the collision dynamics in the all four cases of S_f , (1)–(4). However, the modulating frequency increases overall as S_f increases, which is analogous to the tendency exhibited by $S(u, v) = \sqrt{4\pi}\phi(u, v)$ as shown in Fig. 2. In Fig. 4, we present $Qh^{\text{TT}}(t)r/(2\pi\eta^2)$ instead of $Qh^{\text{TT}}(t)$, and thus all the waveforms are plotted in the same scale. One should note here that our actual numerical data of the energy-momentum tensors T_{ab} for Eq. (43) have been obtained via Eq. (3) after solving Eqs. (2) and (4) simultaneously [18]. Therefore, our T_{ab} contain the full physical information about the bubble collisions in terms of the scalar field $S = \sqrt{4\pi}\phi$, the geometry g_{ab} and the potential $V(S)$; with the radiation reaction effects included in S and g_{ab} .⁴

⁴ The way our GWs are calculated here resembles a “semi-relativistic treatment”, due to Ruffini and Sasaki [21], in the following senses: (a)

3.2 A simplified method to compute gravitational waves in the quadrupole approximation

Reference [19] presents a simplified method to compute the GWs $Qh_{ij}^{TT}(t, \mathbf{x})$ of Eq. (13) by neglecting the gravitational effects on the bubbles: namely, $g_{\mu\nu}$ in Eq. (3) is replaced by $\eta_{\mu\nu}$, assuming that the bubbles are in flat spacetime. Then Eq. (14) can be simplified as

$$\begin{aligned} I_{kl}(t_R) &= \int d^3x' T^{kl}(t_R, \mathbf{x}') \\ &= \int d^3x' \partial_k \phi(t_R, \mathbf{x}') \partial_l \phi(t_R, \mathbf{x}'), \end{aligned} \quad (44)$$

where the energy-momentum tensors from Eq. (3) have been reduced; $T_{ij} \rightarrow \partial_i \phi \partial_j \phi$ because the terms proportional to δ_{ij} in T_{ij} makes no contribution to gravitational radiation (13) due to the property of Eq. (10); namely, $\Lambda_{ij,kl} \delta_{ij} = 0$ [19].

By Eqs. (18) and (44) we have

$$\begin{aligned} I_{kl} &= \delta_{kx} \delta_{lx} \int d^3x' \left[\left(\frac{\partial \phi}{\partial x'} \right)^2 - \frac{1}{2} \left(\frac{\partial \phi}{\partial y'} \right)^2 - \frac{1}{2} \left(\frac{\partial \phi}{\partial z'} \right)^2 \right] \\ &= \delta_{kx} \delta_{lx} \int d^3x' \left[\left(\frac{\partial \phi}{\partial x'} \right)^2 - \frac{1}{2} \left(\frac{\partial \phi}{\partial \rho'} \right)^2 \right]. \end{aligned} \quad (45)$$

Using Eq. (23), we can modify

$$\left(\frac{\partial \phi}{\partial \rho} \right)^2 = \left(\frac{\partial \phi}{\partial \tau} \right)^2 \left(\frac{\partial \tau}{\partial \rho} \right)^2 = \frac{t_R^2 - \tau^2}{\tau^2} \left(\frac{\partial \phi}{\partial \tau} \right)^2. \quad (46)$$

Then in the same manner as described above by Eq. (28), the integral $I_{kl}(t_R)$ is computed out of the volume piece $\mathcal{V} = \Delta x |\tau \Delta \tau| \Delta \theta = \frac{\pi}{4} \eta^2 \Delta x [1 + \mathcal{O}(\eta^2/t_R^2)]$:

$$\begin{aligned} I_{kl}(t_R) &= \delta_{kx} \delta_{lx} \int_{\mathcal{V}} d^3x' \left[\left(\frac{\partial \phi}{\partial x'} \right)^2 - \frac{1}{2} \left(\frac{\partial \phi}{\partial \rho'} \right)^2 \right] \\ &= 2\pi \delta_{kx} \delta_{lx} \int_{t_R - \eta^2/(8t_R)}^{t_R} d\tau' \tau' \int_{-x_0}^{x_0} dx' \\ &\quad \times \left[\left(\frac{\partial \phi}{\partial x'} \right)^2 - \frac{t_R^2 - \tau'^2}{2\tau'^2} \left(\frac{\partial \phi}{\partial \tau'} \right)^2 \right] + \mathcal{O}\left(\frac{\eta^4}{t_R^3}\right) \\ &= \frac{\pi}{2} \delta_{kx} \delta_{lx} \eta^2 \int_0^{x_0} dx' \left(\frac{\partial \phi}{\partial x'} \right)^2_{\tau=t_R} + \mathcal{O}\left(\frac{\eta^4}{t_R^3}\right). \end{aligned} \quad (47)$$

If the wall thickness η can be taken sufficiently small in Eq. (47), by Eqs. (13) and (47) we can compute the bubble-collision-induced GWs in the quadrupole approximation as

the field (h_{ab}) radiates as if it were in flat spacetime, (b) the source (T_{ab}) contains the General Relativistic information about its local spacetime. In Eq. (9) we see that our GWs h_{ab} result from distant sources T_{ab} , which are composed of the scalar field and the local geometry given via Eqs. (3), (5), and (8), thus containing the full General Relativistic information, including the radiation reaction effects.

$$Qh_{ij}^{TT}(t, \mathbf{x}) \approx \frac{2\pi}{r} \Lambda_{ij,kl}(\mathbf{n}) \delta_{kx} \delta_{lx} \eta^2 \int_0^{x_0} dx' \left(\frac{\partial \phi}{\partial x'} \right)^2_{\tau=t_R}. \quad (48)$$

Substituting Eq. (42) into Eq. (48), we finally express

$$\begin{aligned} Qh_{ij}^{TT}(t, \mathbf{x}) &\approx \frac{2\pi}{r} \eta^2 \left[\delta_{ix} \delta_{jx} - 2\delta_{ix} n_j \cos \vartheta \right. \\ &\quad \left. + \frac{1}{2} n_i n_j (1 + \cos^2 \vartheta) - \frac{1}{2} \delta_{ij} \sin^2 \vartheta \right] \\ &\quad \times \int_0^{x_0} dx' \left(\frac{\partial \phi}{\partial x'} \right)^2_{\tau=t_R}, \end{aligned} \quad (49)$$

where $t_R = t - r = t - |\mathbf{x}|$, and η is specified by (29).

Result 2: Toward the end of the bubble collisions, $\tau \gg 1$, the scalar field oscillates around the true vacuum state, i.e. $|\phi| \ll 1$, being nearly *monochromatic*. Then we can approximate Eq. (8) as

$$\square \phi(\tau, x) \approx \frac{6(\beta+1)V_f}{\beta S_f^2} \phi(\tau, x), \quad (50)$$

where we have replaced the curved spacetime Laplacian ∇^2 by the flat spacetime d'Alembertian $\square \equiv -\partial^2/\partial \tau^2 - (\partial/\partial \tau)\partial/\partial \tau + \partial^2/\partial x^2$, neglecting the gravitational effects on the bubbles to simplify the problem.⁵ With the help of Refs. [22, 23], we obtain a solution for Eq. (50):

$$\phi(\tau, x) = \phi_0 \frac{J_1(\omega_t \sqrt{\tau^2 - x^2})}{\omega_t \sqrt{\tau^2 - x^2}}, \quad (51)$$

where J_n denotes the Bessel function of the first kind, $\omega_t \equiv \sqrt{6(\beta+1)V_f/(\beta S_f^2)}$ represents the ‘terminal’ frequency of the bubble collisions, and ϕ_0 is the amplitude which is determined by the initial conditions of the field. Substituting Eq. (51) into Eq. (49), the GWs emitted from the bubble collisions in the final stage can be computed. Figure 5 shows the GWs $Qh_{ij}^{TT}(t)r/(2\pi\eta^2)$ computed with various false vacuum field values, $S_f = \sqrt{4\pi}\phi_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$. Corresponding to the field values are the frequencies, $\omega_t \simeq (1) 0.8124, (2) 0.4062, (3) 0.2708, (4) 0.2031 \sim (1) 1, (2) 1/2, (3) 1/3, (4) 1/4$, as can be seen from Fig. 5. Also, the amplitude of $Qh_{ij}^{TT}(t)r/(2\pi\eta^2)$ scales as (1) 1, (2) 2, (3) 3, (4) 4 due to the factor $\omega_t^{-1} \sim S_f$, as can be seen from Fig. 5. Then the amplitude of $Qh_{ij}^{TT}(t)$ should scale as $\eta^2 \omega_t^{-1} \sim S_f^5 \sim (1) 1, (2) 2^5, (3) 3^5, (4) 4^5$.

⁵ A similar analysis is found in Ref. [17], in which the scalar field equation is solved in hyperbolic “de Sitter” spacetime in the limit, $\tau \gg H^{-1}$, where H is the Hubble parameter. The solution shows fluctuations of decreasing amplitude and increasing period (or decreasing frequency) in τ . However, in our analysis, the equation is solved in hyperbolic “flat” spacetime and our solution given by (51) has fluctuations of decreasing amplitude and fixed period (or single frequency; monochromatic) in τ .

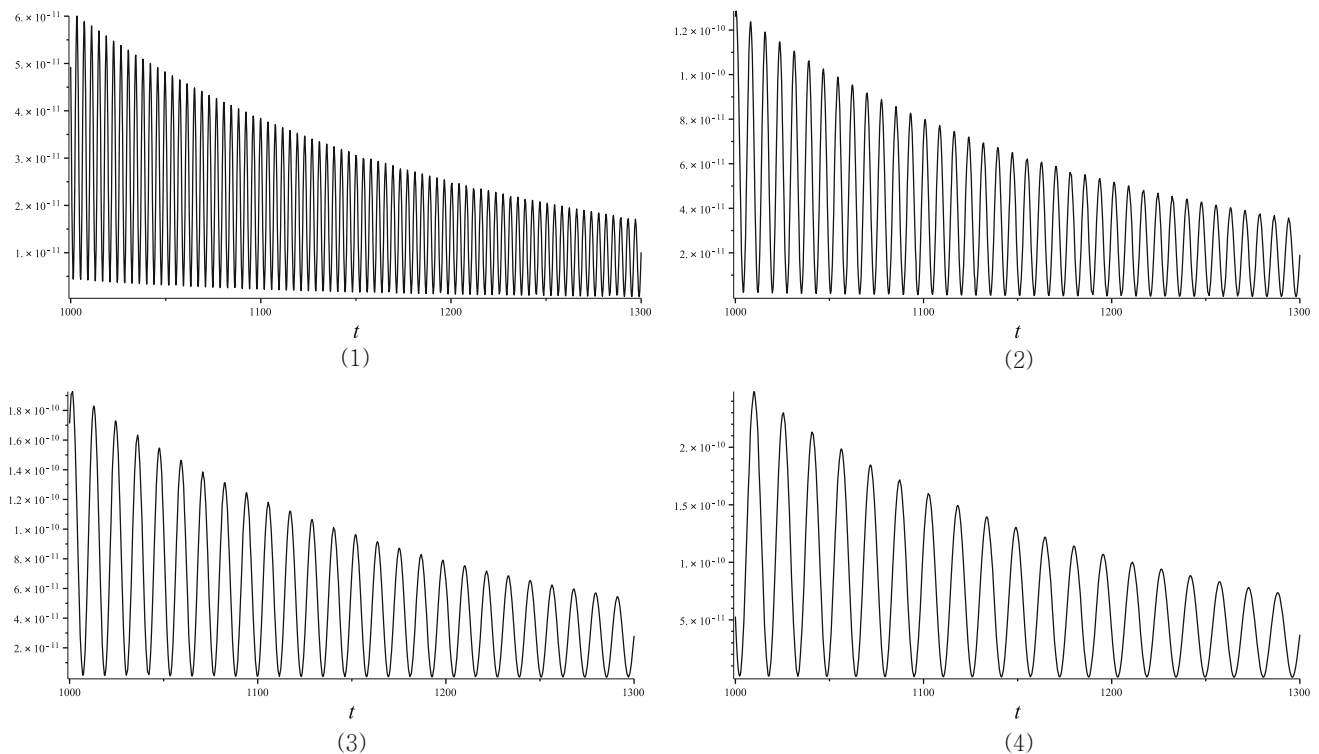


Fig. 5 Plots of $Qh^{TT}(t)r/(2\pi\eta^2)$ computed via the expression (49), with ϕ given by (51). Corresponding to the false vacuum field values, $S_f = \sqrt{4\pi}\phi_f = (1) 0.1, (2) 0.2, (3) 0.3, (4) 0.4$, are the frequencies, $\omega_t \simeq (1) 0.8124, (2) 0.4062, (3) 0.2708, (4) 0.2031 \sim (1) 1, (2) 1/2, (3)$

$1/3, (4) 1/4$. The amplitude of $Qh^{TT}(t)r/(2\pi\eta^2)$ scales as (1) 1, (2) 2, (3) 3, (4) 4 due to the factor $\omega_t^{-1} \sim S_f$. Then the amplitude of $Qh^{TT}(t)$ should scale as $\eta^2\omega_t^{-1} \sim S_f^5 \sim (1) 1, (2) 2^5, (3) 3^5, (4) 4^5$

4 Conclusions

We have computed GWs emitted from collisions of two equal-sized bubbles in time domain. The waveforms have been obtained for:

- (i) the initial-to-intermediate stage of strong collisions, and
- (ii) the final stage of weak collisions, in *full General Relativity* and in the flat spacetime approximation, using numerical and analytical methods, respectively.

During (i), the waveforms show the non-linearity of the collisions, characterized by a modulating frequency and cusp-like bumps, whereas during (ii), the waveforms exhibit the linearity of the collisions, featured by a constant frequency and smooth oscillations, as can be checked from Figs. 4 and 5, respectively. Also, depending on the false vacuum field value ϕ_f , the waveforms have different scales of frequency. During (i), the modulating frequency increases overall as the false vacuum field value ϕ_f increases, whereas during (ii), the frequency ω_t decreases as the false vacuum field value ϕ_f increases in an inversely proportional relationship, i.e. $\omega_t \sim \phi_f^{-1}$. It is interesting to note that the relationship between the false vacuum field value and the frequency during (i) changes almost inversely during (ii). In addition,

the false vacuum field value ϕ_f affects the amplitude of the waveforms. During (i), the amplitude scales as $\eta^2 \sim \phi_f^4$, whereas during (ii), the amplitude scales as $\eta^2\omega_t^{-1} \sim \phi_f^5$, where η is a bubble wall thickness.

One of the notable differences between the waveforms emitted during (i) and during (ii) is the sign, as can be seen from Figs. 4 and 5. This is due to the difference between Eqs. (43) and (49): the integral in (49) is always positive while its counterpart in (43) is not necessarily. This has to do with the composition of the integrands in the two expressions. The integrand in (43) consists of the energy-momentum tensors T_{ab} which have been obtained via Eq. (3) after solving Eqs. (2) and (4) simultaneously [18]: thus T_{ab} contain the full physical information of bubble collisions in terms of the scalar field ϕ , the geometry g_{ab} and the potential $V(\phi)$; with the radiation reaction effects included in ϕ and g_{ab} . However, as explained in the beginning of Sect. 3.2, the integrand in (49) comes only from the first term, with the second and third terms being disregarded in (3) as the gravitational effects on the bubbles are assumed to be neglected, following Ref. [19]. This, combined with the thin-wall approximation, results in the integrand in (49) being positive, which leads to the integral being also positive. But this is not the case for the integral in (43) due to the minus signs appearing in (3) and in the integrand in (43).

Throughout the paper, we used the thin-wall and quadrupole approximations to simplify our computations. These approximations served our purpose well in that we were able to gain some *qualitative* insights into the time-domain gravitational waveforms emitted from bubble collisions. However, to obtain more physically reasonable waveforms, taking into account a generic thickness and relativistic motion of bubble wall, it will be inevitable to include in our computations the next-to-leading order corrections beyond each approximation. A huge amount of computation will be involved in this task, and we leave it for follow-up studies.

Acknowledgments The authors would like to thank Dong-il Hwang for his valuable comments and assistance during an early stage of this work. The authors also would like to thank Hongsu Kim, Sang Pyo Kim, Hyung Won Lee, Gungwon Kang, and Inyong Cho for fruitful discussions and helpful comments. BHL, WL, and DY appreciate Pauchy W. Y. Hwang and Sang Pyo Kim for their hospitality at the 9th International Symposium on Cosmology and Particle Astrophysics in Taiwan, 13–17 November, 2012. DHK and WL appreciate APCTP for its hospitality during completion of this work. DHK and JY were supported by Basic Science Research Program through the National Research Foundation of Korea (NRF) funded by the Ministry of Education (2013R1A1A2008901 and 2013R1A1A2A10004883). BHL was supported by the National Research Foundation of Korea (NRF) grant funded by the Korea government (MSIP) (2014R1A2A1A01002306). WL was supported by Basic Science Research Program through the National Research Foundation of Korea (NRF) funded by the Ministry of Education (2012R1A1A2043908). DY was supported by the JSPS Grant-in-Aid for Scientific Research (A) (No. 21244033) and also supported by Leung Center for Cosmology and Particle Astrophysics (LeCosPA) of National Taiwan University (103R4000).

Open Access This article is distributed under the terms of the Creative Commons Attribution License which permits any use, distribution, and reproduction in any medium, provided the original author(s) and the source are credited.

Funded by SCOAP³ / License Version CC BY 4.0.

References

1. BICEP2 2014 Results Release, <http://bicepkeck.org/>
2. R. Flauger, J.C. Hill, D.N. Spergel, Toward an understanding of foreground emission in the BICEP2 region. *JCAP* **8**, 39 (2014)
3. A.H. Guth, Inflationary universe: a possible solution to the horizon and flatness problems. *Phys. Rev. D* **23**, 347 (1981)
4. A.D. Linde, A new inflationary universe scenario: a possible solution of the horizon, flatness, homogeneity, isotropy and primordial monopole problems. *Phys. Lett. B* **108**, 389 (1982)
5. A. Albrecht, P.J. Steinhardt, Cosmology for grand unified theories with radiatively induced symmetry breaking. *Phys. Rev. Lett.* **48**, 1220 (1982)
6. A.D. Linde, Nonsingular regenerating inflationary universe (Cambridge University preprint, print-82-0554) (1982)
7. A.H. Guth, Eternal inflation and its implications. *J. Phys. A* **40**, 6811 (2007)
8. A. Vilenkin, The birth of inflationary universes. *Phys. Rev. D* **27**, 2848 (1983)
9. S.R. Coleman, F. De Luccia, Gravitational effects on and of vacuum decay. *Phys. Rev. D* **21**, 3305 (1980)
10. S. Parke, Gravity and the decay of the false vacuum. *Phys. Lett. B* **121**, 313 (1983)
11. B.-H. Lee, W. Lee, Vacuum bubbles in a de Sitter background and black hole pair creation, *Class. Quantum Grav.* **26**, 225002 (2009). [arXiv:0809.4907](https://arxiv.org/abs/0809.4907)
12. D.H. Lyth, E.D. Stewart, Thermal inflation and the moduli problem. *Phys. Rev. D* **53**, 1784 (1996). [arXiv:hep-ph/9510204](https://arxiv.org/abs/hep-ph/9510204)
13. R. Easther, J.T. Giblin Jr, E.A. Lim, W.I. Park, E.D. Stewart, Thermal inflation and the gravitational wave background. *JCAP* **0805**, 013 (2008). [arXiv:0801.4197](https://arxiv.org/abs/0801.4197) [astro-ph]
14. S. Chang, M. Kleban, T.S. Levi, When worlds collide. *JCAP* **0804**, 034 (2008). [arXiv:0712.2261](https://arxiv.org/abs/0712.2261) [hep-th]
15. S.W. Hawking, I.G. Moss, J.M. Stewart, Bubble collisions in the very early universe. *Phys. Rev. D* **26**, 2681 (1982)
16. Z.-C. Wu, Gravitational effects in bubble collisions. *Phys. Rev. D* **28**, 1898 (1983)
17. M.C. Johnson, H.V. Peiris, L. Lehner, Determining the outcome of cosmic bubble collisions in full general relativity. *Phys. Rev. D* **85**, 083516 (2012). [arXiv:1112.4487](https://arxiv.org/abs/1112.4487)
18. D. Hwang, B.-H. Lee, W. Lee, D. Yeom, Bubble collision with gravitation, *JCAP* **1207**, 003 (2012). [arXiv:1201.6109](https://arxiv.org/abs/1201.6109) [gr-qc]
19. A. Kosowsky, M.S. Turner, R. Watkins, Gravitational radiation from colliding vacuum bubbles. *Phys. Rev. D* **45**, 4514 (1992)
20. C. Caprini, R. Durrer, G. Servant, Gravitational wave generation from bubble collisions in first-order phase transitions: An analytic approach. *Phys. Rev. D* **77**, 124015 (2008). [arXiv:0711.2593](https://arxiv.org/abs/0711.2593) [astro-ph]
21. R. Ruffini, M. Sasaki, On a semi-relativistic treatment of the gravitational radiation from a mass thrust into a black hole. *Prog. Theor. Phys.* **66**, 1627 (1981)
22. A.D. Polyanin, V.F. Zaitsev, *Handbook of Nonlinear Partial Differential Equations*, 2nd edn (CRC Press, Boca Raton, 2012)
23. A.D. Polyanin, V.F. Zaitsev, *Handbook of Exact Solutions for Ordinary Differential Equations*, 2nd edn (CRC Press, Boca Raton, 2003)