

Right handed sneutrino dark matter in inverse and linear seesaw scenarios

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Abstract. We consider supersymmetric models in which the right-handed sneutrino is a viable WIMP dark matter candidate. These are either simple extensions of the Minimal Supersymmetric Standard Model or models with the addition of an extra U(1) group. All of them can explain small neutrino masses, through either the Inverse or the Linear Seesaw mechanism. We investigate the properties of the dark matter candidate naturally arising in these scenarios. We check for phenomenological bounds, such as correct relic abundance, consistency with direct detection cross section limits and laboratory constraints. Especially, we comment on limitations of the model space due to lepton flavour violating charged lepton decays.

1. Introduction

We analyze two supersymmetric particle physics models, which contain a Right-Handed (RH) sneutrino $\tilde{\nu}_{LSP}$ as lightest supersymmetric particle (LSP) and thus as a good WIMP (weakly interactive massive particle) cold Dark Matter (CDM) candidate.

The two models we consider are:

Model I - A simple extension of the MSSM, which accounts for the *inverse seesaw* model [1], where the particle content of the SM is extended by three pairs of singlets, called ν^c and $\hat{S}$, which form “heavy” pseudo-Dirac pairs, or for an alternative seesaw scheme that can also be realized at low-scale: the *linear seesaw*, which has been suggested as arising from a particular $SO(10)$ unified model [2, 3, 4].

The total superpotential contains then the additional terms $M_R \hat{\nu} \hat{S}$, (Lepton Number - LN conserving) and $\frac{1}{2} \mu_S \hat{S} \hat{S}$, (LN violating):

$$W = Y_u \hat{u}^c \hat{q} \hat{H}_u - Y_d \hat{d}^c \hat{q} \hat{H}_d - Y_e \hat{e}^c \hat{l} \hat{H}_d + \mu \hat{H}_u \hat{H}_d + Y_\nu \hat{\nu} \hat{l} \hat{H}_u + M_R \hat{\nu}^c \hat{S} + \frac{1}{2} \mu_S \hat{S} \hat{S} \quad (1)$$

where Y_e , Y_d and Y_u are the usual MSSM Yukawa couplings for the charged leptons and the quarks. In addition there is the neutrino Yukawa coupling Y_ν .

In the linear seesaw scheme, the superpotential is

$$W = Y_u \hat{u}^c \hat{q} \hat{H}_u - Y_d \hat{d}^c \hat{q} \hat{H}_d - Y_e \hat{e}^c \hat{l} \hat{H}_d + \mu \hat{H}_u \hat{H}_d + Y_\nu \hat{\nu} \hat{l} \hat{H}_u + M_R \hat{\nu}^c \hat{S} + Y_{SL} \hat{S} \hat{l} \hat{H}_u \quad (2)$$

and Lepton Number is violated by Y_{SL} .



Model II - A left-right symmetric extension of the SM based on a variant of the basic SUSY $SO(10)$ model advocated in [2] in which an extended intermediate $U(1)_R \times U(1)_{B-L}$ stage follows a short $SU(2)_R \times U(1)_{B-L}$ phase, which can emerge, e.g., in a class of $SO(10)$ GUTs broken along the “minimal” left-right symmetric chain

$$SO(10) \rightarrow SU(3)_c \times SU(2)_L \times SU(2)_R \times U(1)_{B-L} \rightarrow SU(3)_c \times SU(2)_L \times U(1)_R \times U(1)_{B-L},$$

in which the low energy scale seesaw schemes are actually naturally realized.

The R -parity conserving superpotential is

$$W = Y_u \hat{u}^c \hat{q} \hat{H}_u - Y_d \hat{d}^c \hat{q} \hat{H}_d + Y_\nu \hat{\nu}^c \hat{l} \hat{H}_u - Y_e \hat{e}^c \hat{l} \hat{H}_d + Y_s \hat{\nu}^c \hat{\chi}_R \hat{S} + \mu \hat{H}_u \hat{H}_d - \mu_R \hat{\chi}_R \hat{\chi}_R + \frac{1}{2} \mu_S \hat{S} \hat{S} \quad (3)$$

with the same notation as before. The main introduction here is the term featuring the Yukawa coupling Y_s , which mixes the ν^c fields with the S fields giving rise to an inverse seesaw mechanism for neutrino masses.

2. Phenomenological constraints

2.1. Neutrino masses

We fit neutrino masses according to neutrino oscillation data, i.e. with the atmospheric (atm) and solar ($\odot$) scales as in the case of normal hierarchy and the lepton mixing angles according to the most recent measurements of reactor antineutrino disappearance - Double CHOOZ, Daya Bay and RENO [5, 6, 7, 8]

2.2. Lepton flavor violation

Putting all flavor in Y_ν is actually the usual option, thus keeping Y_L , M_R and μ_S diagonal. However, this framework gives strong constraints on the sneutrino LSP coming from LFV processes notably $\mu \rightarrow e\gamma$ and three leptons decays like $\mu \rightarrow eee$.

Experiments set the bounds $\text{BR}(\mu \rightarrow e\gamma) \leq 2.4 \times 10^{-12}$ (90% C.L.) [9, 10] and $\text{BR}(\mu \rightarrow eee) \leq 1.0 \times 10^{-12}$ (90% C.L.) [11]. Indeed, it has been shown [12] that in models with an extended particle content, $l_i \rightarrow 3l_j$ can be more constraining than $l_i \rightarrow l_j\gamma$.

To really relax the LFV constraints, the best option is to put off diagonal terms in μ_S (in the case of the inverse seesaw) or in Y_{SL} (linear seesaw) and to keep M_R and Y_ν diagonal

2.3. Invisible Z width constraints

Another experimental bound on MSSM neutrinos lighter than $m_Z/2$, comes from their contribution to the invisible Z -width, measured in searches at accelerators. Since in our models we have introduced also RH neutrinos ν^c , the mixing ($\sin\theta$) between the states has to be taken into account. The invisible Z -width including the mixing reads:

$$\Delta\Gamma_Z = \sin^4\theta \frac{\Gamma_\nu}{2} \left[1 - \left(\frac{2m_\nu}{m_Z} \right)^2 \right]^{3/2} \theta(m_Z - 2m_\nu). \quad (4)$$

A further contribution should in principle be considered, that is the left - right mixing of the $\tilde{\nu}_{LSP}$, which could also participate in the Z decay. However, it becomes important only when the mass of the $\tilde{\nu}_{LSP}$ is lighter than $m_Z/2$.

3. Sneutrino as Dark Matter candidate

The models considered in this paper are by construction very similar to the MSSM. In the ordinary MSSM the LSP is the neutralino in extended regions of the parameter space, which anyway provides an interesting candidate of dark matter. However our interest here is to study another dark matter candidate, which naturally arises in this simple extensions of the MSSM, which is the lightest of the supersymmetric partners of the RH neutrinos, the $\tilde{\nu}_{LSP}$. It is an interesting candidate too, since it has electroweak scale interactions, and because of the mixing between the LH and RH components, does not suffer from the problems derived from a too strong coupling to the Z boson like the LH sneutrino of the MSSM.

The existence of the RH sneutrino as LSP does not exclude the possibility of having neutralino as LSP, though; the mass of the LSP strongly depends on the new parameters added to the theory. There will then be just some regions of the parameter space where the RH sneutrino is the LSP.

3.1. Model I with inverse seesaw

The $\tilde{\nu}_{LSP}$ has a direct coupling to the Higgs, as can be seen from the term $Y_\nu \nu l H_u$ of the superpotential Eq. 1. This coupling depends on the Yukawa coupling Y_ν and on the trilinear term T_{Y_ν} and provides tree level interactions with the SM particles at the electroweak scale, thus making the $\tilde{\nu}_{LSP}$ a possible WIMP candidate. The strongest constraint on $\tilde{\nu}_{LSP}$ to be a viable CDM candidate, comes from the CMB measurements, that is from its relic abundance (RA).

In order to perform the analysis of the sneutrino dark matter phenomenology, we have created a suitable model file for this model with SARAH [13]. Then, we have used SPheno [14] to compute the low energy spectrum through a precise mass calculation using two-loop RGEs and one-loop corrections to all masses. The calculation of the relic density of the LSP is then done with MicrOmegas [15] version 2.4.5 based on the CalcHep output of SARAH. To perform the scans we used a Mathematica package (SSP) [16].

The main annihilation channels are:

- $\tilde{\nu}_{LSP}\tilde{\nu}_{LSP} \rightarrow W^+W^-, Z^0Z^0, f\bar{f}$ via s-channel Higgs (h^0, H^0, A^0) exchange;
- $\tilde{\nu}_{LSP}\tilde{\nu}_{LSP} \rightarrow W^+W^-, Z^0Z^0$ via t- and u- channel $\tilde{l}$ exchange, and through a quartic coupling;
- $\tilde{\nu}_{LSP}\tilde{\nu}_{LSP} \rightarrow H^0H^0$ through a quartic coupling which results to be very efficient;

In Fig.1, where the $\tilde{\nu}_{LSP}$ is mainly LH, it behaves like a MSSM sneutrino. When its mass approaches half the mass of the Higgs, the "Higgs pole" is clearly visible, which is due to the opening of the annihilation channel through the Higgs boson. This channel results to be highly efficient, indeed the RA goes down to very small values of the order of 10^{-5} . Those points refer the sneutrino could not account for the totality of the CDM relic abundance, but it would still be a valid subdominant dark matter candidate. The analogue Z pole is instead suppressed, because of the dependence of the RA on Y_ν ($\propto Y_\nu^4$) and because of helicity. Another efficient channel is the quartic coupling to the Higgses, which appears when the mass of the $\tilde{\nu}_{LSP}$ is larger than the mass of the Higgs.

3.1.1. Direct Detection Direct detection of the sneutrinos would consist in detecting a signal coming from their elastic scattering with nuclei inside a detector. The direct detection rate of the DM particle depends on its mass and on the scattering cross section. The interaction, which occurs in a non relativistic limit, provided that the velocity of dark matter particles in the galactic halo is small, consists of basically two diagrams contributing at tree level, which are the t-channel exchange of a neutral Higgs or of the Z boson. This last diagram is however suppressed because of the mostly singlet nature of the $\tilde{\nu}_{LSP}$.

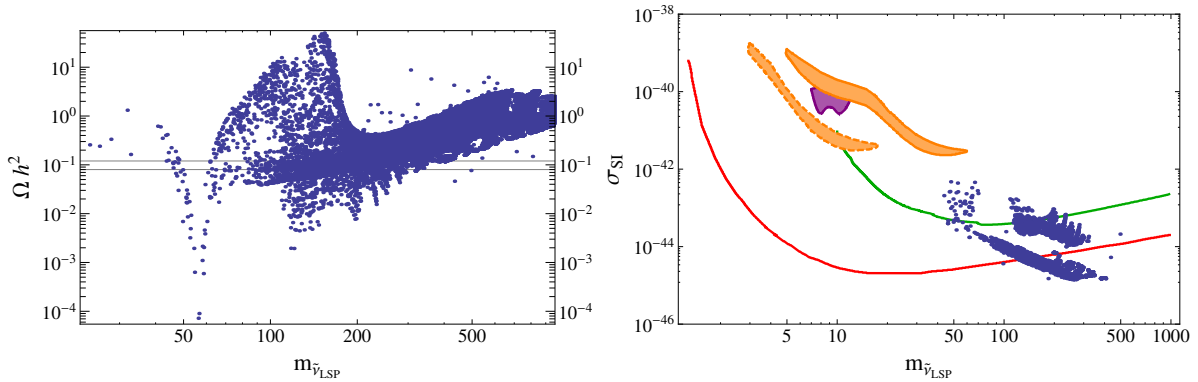


Figure 1. Fig. a (left): General scan for Model I, showing the RA versus the $\tilde{\nu}_{LSP}$ mass. Fig. b (right): Direct detection cross section for model I, and current limits from XENON100 [17] (red line), CDMS [18] (green line), DAMA (with and without channeling, orange regions) [19], and Cogent [20] (purple region).

Lepton number violation in the sneutrino mass matrix leads to a mass splitting between the real and the imaginary part of the lightest sneutrino, and the scattering via Z boson exchange occurs inelastically, through a transition from the real to the imaginary or viceversa. If the mass splitting is greater than some keV, scattering is strongly suppressed at direct detection experiments.

In Fig.1b we compare the direct detection cross section for the RH sneutrino LSP with current direct detection experiments. The major bound nowadays comes from the XENON100 experiment [17], whose best sensitivity is around 10^{-44}cm^2 for a dark matter candidate of 50 GeV. We depict the direct detection cross section versus the LSP sneutrino mass (blue points). The points are those which survive the relic abundance constraint. In the same plot, the current limits from XENON100 [17] (red line), CDMS [18] (green line), DAMA (with and without channeling, orange regions) [19], and Cogent [20] (purple region) are shown. The sneutrinos show a SI cross section $\sigma_{SI} \lesssim 10^{-42}\text{cm}^2$, and for masses $m_{\tilde{\nu}_{LSP}} \gtrsim 100$ GeV they are compatible with current limits by XENON100. However, XENON1T, whose sensibility should improve up to 10^{-46}cm^2 , will test those cross sections.

3.2. Model I with linear seesaw

The phenomenology of sneutrino dark matter in the MSSM with linear seesaw is basically the same as in the previous case. Indeed, it is not possible to distinguish between the kind of seesaw mechanism through the dark matter phenomenology.

3.3. Model II

In this subsection we discuss the DM phenomenology in the minimal supersymmetric $U(1)_R \times U(1)_{B-L}$ extension of the standard model. The presence of the extra gauge boson Z' leads to important differences with respect to the previous models.

We use a version of the model where the tadpole equations for the vevs are solved analytically for $(\mu, B_\mu, \mu_R, B_{\mu_R})$ and we perform a scan over the parameters $m_0, M_{1/2}, \tan\beta_R$. The results are shown in Figg. 2.

In both plots the main feature clearly visible is the Z' pole. Indeed, the annihilation of the $\tilde{\nu}_{LSP}$ LSPs into SM particles via the Z' becomes efficient when the mass of the $\tilde{\nu}_{LSP}$ is close

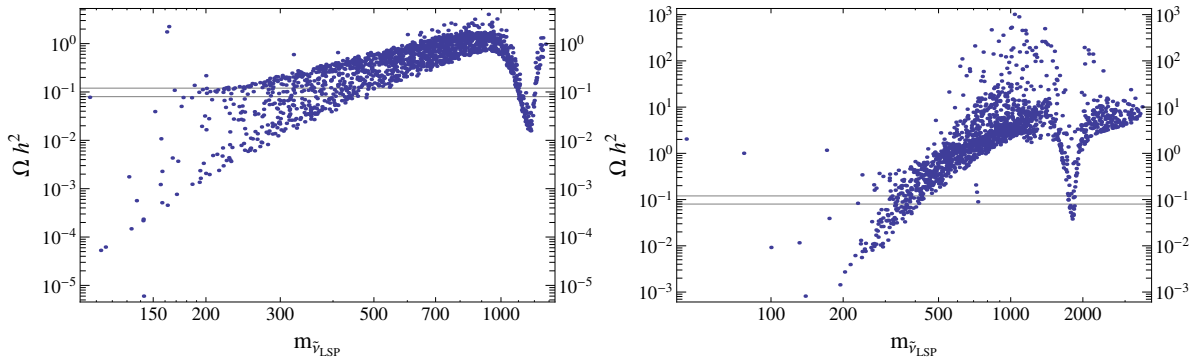


Figure 2. Scans for model II, made over the parameters $m_0 = [0, 6000]$ GeV, $M_{1/2} = [3000, 8000]$ GeV, $\tan\beta_R = [1.0, 1.3]$. The other parameters are set to the values $\tan\beta=10$, $A_0 = -4500$ GeV, $Y_S = \text{diag}(0.3)$; v_R has been chosen different in the two plots, $v_R = 5.5$ TeV (left) and $v_R = 10$ TeV (right).

to half the mass of the Z' . The mass of the Z' can be calculated analytically [21] and mainly depends on the value of v_R . The ATLAS searches for a Z' set a lower limit on its mass which is 1.8 TeV, and this translates into a lower limit on $v_R \gtrsim 5$ TeV. In Fig. 2 we chose two different values of v_R : $v_R = 5.5$ TeV and $v_R = 10$ TeV, thus leading to two values of Z' and so to a shift in the position of the pole in the relic density plot. This shows that choosing a higher value of v_R we can get heavy $\tilde{\nu}_{LSP}$ DM with the correct RA.

The quartic coupling with two Higgses (h^0 , h_{BLR}^0 and A^0 , depending on if they are kinematically allowed, according to the $\tilde{\nu}_{LSP}$ mass) is one of the most efficient annihilation channels, as before. For lower masses the annihilation via the MSSM Higgs is the most efficient, as can be noticed by the lowering of the relic density going to lower masses, especially in the first plot, where on the left end side we are approaching the Higgs pole.

3.3.1. Direct Detection As to the direct detection cross section analysis, a bound on the v_R , and then on the mass of the Z' can be set by the DD experiments, especially by Xenon100. This bound turns to be stronger than the one set by the collider experiments like LHC, as it is shown in Fig.3. The two clouds of points there refer to $v_R = 5.5$ TeV (black) and $v_R = 10$ TeV (blue).

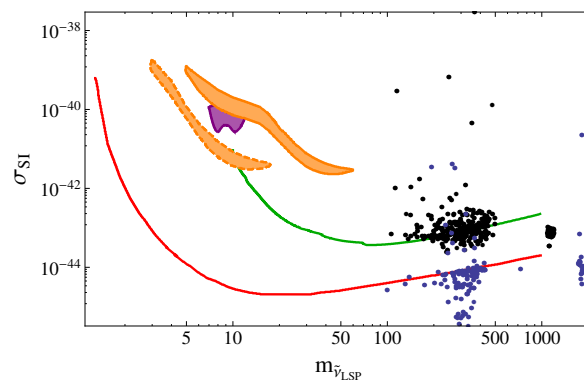


Figure 3. Direct Detection cross section limits for model II. The two clouds of points refer to $v_R = 5.5$ TeV (black) and $v_R = 10$ TeV (blue). Experimental limits as in Fig. 1b.

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References

- [1] Mohapatra R and Valle J 1986 *Phys.Rev.* **D34** 1642
- [2] Malinsky M, Romao J and Valle J 2005 *Phys.Rev.Lett.* **95** 161801 (*Preprint hep-ph/0506296*)
- [3] Akhmedov E K, Lindner M, Schnapka E and Valle J 1996 *Phys.Lett.* **B368** 270–280 (*Preprint hep-ph/9507275*)
- [4] Akhmedov E K, Lindner M, Schnapka E and Valle J 1996 *Phys.Rev.* **D53** 2752–2780 (*Preprint hep-ph/9509255*)
- [5] Abe Y *et al.* (DOUBLE-CHOOZ Collaboration) 2012 *Phys.Rev.Lett.* **108** 131801 (*Preprint 1112.6353*)
- [6] An F *et al.* (DAYA-BAY Collaboration) 2012 *Phys.Rev.Lett.* **108** 171803 (*Preprint 1203.1669*)
- [7] Ahn J *et al.* (RENO collaboration) 2012 *Phys.Rev.Lett.* **108** 191802 (*Preprint 1204.0626*)
- [8] Forero D, Tortola M and Valle J 2012 *Phys.Rev.* **D86** 073012 (*Preprint 1205.4018*)
- [9] Nishimura K, Dey B, Aston D, Leith D, Ratcliff B *et al.* 2013 *Nucl.Instrum.Meth.* **A701** 115–126
- [10] Adam J *et al.* (MEG collaboration) 2011 *Phys.Rev.Lett.* **107** 171801 (*Preprint 1107.5547*)
- [11] Nakamura K *et al.* (Particle Data Group) 2010 *J.Phys.* **G37** 075021
- [12] Hirsch M, Staub F and Vicente A 2012 *Phys.Rev.* **D85** 113013 (*Preprint 1202.1825*)
- [13] Staub F 2008 (*Preprint 0806.0538*)
- [14] Porod W 2003 *Comput.Phys.Commun.* **153** 275–315 (*Preprint hep-ph/0301101*)
- [15] Belanger G, Boudjema F, Pukhov A and Semenov A 2007 *Comput.Phys.Commun.* **176** 367–382 (*Preprint hep-ph/0607059*)
- [16] Staub F, Ohl T, Porod W and Speckner C 2012 *Comput.Phys.Commun.* **183** 2165–2206 (*Preprint 1109.5147*)
- [17] Aprile E *et al.* (XENON100 Collaboration) 2012 (*Preprint 1207.3458*)
- [18] Ahmed Z *et al.* (CDMS-II Collaboration) 2010 *Science* **327** 1619–1621 (*Preprint 0912.3592*)
- [19] Bernabei R *et al.* (DAMA Collaboration, LIBRA Collaboration) 2010 *Eur.Phys.J.* **C67** 39–49 (*Preprint 1002.1028*)
- [20] Aalseth C *et al.* (CoGeNT collaboration) 2011 *Phys.Rev.Lett.* **106** 131301 (*Preprint 1002.4703*)
- [21] Hirsch M, Porod W, Reichert L and Staub F 2012 *Phys.Rev.* **D86** 093018 (*Preprint 1206.3516*)