

The H I Content of Compact Groups of Galaxies

J. B. Stevens^{A,B,E}, R. L. Webster^{A,F}, D. G. Barnes^A, D. J. Pisano^C, and M. J. Drinkwater^D

^A School of Physics, University of Melbourne, Melbourne VIC 3010, Australia

^B Affiliated with Australia Telescope National Facility, Epping NSW 2121, Australia

^C NSF MPS Distinguished Research Fellow, Australia Telescope National Facility, Epping NSW 2121, Australia

^D Department of Physics, University of Queensland, Brisbane QLD 4072, Australia

^E E-mail: jstevens@physics.unimelb.edu.au

^F E-mail: rwebster@physics.unimelb.edu.au

Received 2003 July 4, accepted 2004 June 1

Abstract: The H I content of Hickson Compact Groups in the southern hemisphere is measured using data from the H I Parkes All-Sky Survey (HIPASS), and dedicated observations using the narrow band filter on the Multibeam instrument on the Parkes telescope. The expected H I mass of these groups was estimated using the luminosity, diameter, and morphological types of the member galaxies, calibrated from published data. Taking careful account of non-detection limits, the results show that the compact group population that has been detected by these observations has an H I content similar to that of galaxies in the reference field sample. The upper limits for the undetected groups lie within the normal range; improvement of these limits will require a large increase in sensitivity.

Keywords: galaxies: evolution — galaxies: interactions — methods: observational

1 Introduction

Neutral hydrogen gas (H I) on the edges of spiral galaxies is only very loosely gravitationally bound. Therefore, when spiral galaxies interact, the outer H I can be easily disturbed (Hibbard & Van Gorkom 1996). During a merger, the H I may be heated and ionised, or it may cool into molecular (H₂) clouds within the remnant (Hibbard & Van Gorkom 1996). It may be tidally driven into the centre of the galaxy and converted into molecular gas or could be turned into stars when tidal streams collide and become compressed. The gas may also escape the interacting system and disperse, becoming too diffuse to detect or perhaps ionised. This paper re-examines the evidence for H I deficiency in compact groups of galaxies using a new survey, the H I Parkes All Sky Survey.

Compact groups of galaxies are excellent places to look for interactions and mergers between (spiral) galaxies. A compact group is defined as having several galaxies with similar redshifts within a small area of the sky, which are also isolated from surrounding galaxies. This ensures that the cores of clusters are excluded from the definition.

A compact group generally has a low velocity dispersion, $\sim 250 \text{ km s}^{-1}$ (Hickson 1997), and a crossing time much smaller than the Hubble time, meaning interactions and mergers are very likely. Although the fraction of spirals in the field is higher ($\sim 80\%$), the fraction of compact group galaxies which are spirals is still $\sim 50\%$ (Hickson 1997).

Although interactions are thought to be common in compact groups, the star-formation rate (SFR) and the amount of molecular gas in the galaxies within the compact groups are only weakly enhanced in comparison to field galaxies (Leon, Combes, & Menon 1998). Leon et al. (1998) concluded that only compact groups with a very small mean separation of $< 30 \text{ kpc}$ had a strong H₂ enhancement. Verdes-Montenegro et al. (1998) also showed that while 20% of compact groups had an apparent deficiency of CO emission, the rest had CO and FIR properties similar to isolated galaxies. Both Leon et al. (1998) and Verdes-Montenegro et al. (1998) concluded that tidal interactions in compact groups were very important in the groups' evolution.

The aim of this paper is to determine the extent to which the H I in compact groups has been affected by the environment. This has been investigated previously by Williams & Rood (1987, hereafter WR87), Huchtmeier (1997, hereafter H97), and Verdes-Montenegro et al. (2001), all of whom observed subsamples of the Hickson Compact Groups (HCGs; Hickson 1982, hereafter H82). Their results indicated that the average amount of H I in the groups was between 40% and 50% of the mass expected by summing the mean H I mass in comparable field galaxies.

WR87 compared the H I content of the compact groups with that of a reference sample of loose groups (RLG), and with a sample of 204 spiral galaxies in the RC2 (de Vaucouleurs et al. 1976). In the south, WR87 observed HCGs

3, 14, 16, 19, 22, 23, 25, 26, 30, 31, 40, 62, 64, 67, 88, 89, and 97. The comparison loose groups were selected to have the same number of galaxies as the compact groups, and to have the same joint distribution of galaxy luminosity and Hubble morphological type. The expected masses therefore were a function of the luminosity of the galaxies. Twenty-seven of the 34 detected compact groups had less H I than the comparable loose group.

A similar procedure was used by Huchtmeier (1997), using an explicit relation between the integrated blue luminosities of the groups and the expected H I masses of the groups. The groups observed in the south were HCGs 4, 14, 19, 21, 22, 24, 25, 26, 30, 40, 42, 48, 62, 67, 87, 91, and 97. In comparison to four nearby groups, a reference sample of 146 galaxies (Huchtmeier & Richter 1988), and to the Virgo cluster spirals, several compact groups appeared to have a very low H I mass-to-blue luminosity ratio.

These authors did not consider how the H I mass of the compact groups related to the distribution of H I masses in the comparison samples. Instead, only the average values of the sample galaxy masses were used. However, any two galaxies with identical optical measurements (diameter, morphology, luminosity) will not necessarily have the same H I mass. There may be a dependence of the H I mass on these observables, but the scatter in this correlation is significant.

Verdes-Montenegro et al. (2001) compiled the data from WR87 and H97, and added VLA imaging of 16 HCGs (2, 16, 18, 23, 26, 31, 33, 40, 44, 49, 54, 79, 88, 92, 95, 96). Using all 72 groups from the combined sample, Verdes-Montenegro et al. (2001) calculated the mean H I ‘deficiency’ to be $\text{DEF}_{\text{H I}} = 0.40 \pm 0.07$, where $\text{DEF}_{\text{H I}} \equiv \log[M_{\text{H I, predicted}}] - \log[M_{\text{H I, observed}}]$. A comparison sample from Haynes & Giovanelli (1984) was used. This comparison sample is also used in this paper and is described in Section 4. Calculated values of $\text{DEF}_{\text{H I}}$ suggested that compact groups contained only 40% of the H I mass of comparable field galaxies. Verdes-Montenegro et al. (2001) also deduced that an evolutionary sequence for HCGs could be followed, beginning with H I being associated with individual galaxies, through to the gas enveloping the entire group, to a group having no detectable H I.

This paper compares the H I content of the galaxies in compact groups with estimated values of H I mass for field galaxies. The distribution of H I mass in the reference field galaxies will be specifically included in the comparison. Thus it is possible to identify groups with masses outside the expected range. The upper limits of the non-detected groups are also taken into account using the statistical methods of Buckley & James (1979).

This paper first deals with the selection of the group sample (Section 2), and then with the H I observations, reduction and the derived parameters (Section 3). The mass estimation methods are described in Section 4, as well as the estimated H I content of the groups. Section 5 compares the relative contents of the

compact groups and the reference sample. For this paper, $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2 Sample Selection

Compact groups were selected from the optical catalogues of H82 and Prandoni, Iovino, & MacGillivray (1994, hereafter P94). The sample included all compact groups which had a declination $\delta < +2^\circ$, and a velocity $cz < 12700 \text{ km s}^{-1}$. A total of 62 groups satisfied these criteria, and are the primary sample of this paper.

HIPASS covers the entire southern sky up to a declination of $\delta < +2^\circ$, over the H I velocity range $-1280 < cz < 12700 \text{ km s}^{-1}$ (Barnes et al. 2001). Using the 13-beam multibeam instrument on the Parkes radio telescope, the survey was completed in 2000 March. The velocity resolution of the survey is 18.0 km s^{-1} , and the RMS noise per channel is typically 13 mJy.

The Hickson compact groups (H82) — for which a complete catalogue of properties of the constituent galaxies is available (Hickson 1993) — were selected so that a direct comparison with previous results could be made. The Southern compact groups (SCGs; P94) were selected by an automatic routine, and therefore have the potential to be a more complete and less biased sample (P94). However, there has not been a published follow-up survey of the member galaxies, and thus the SCGs do not have complete redshift information.

H82 selected groups from the red Palomar Observatory Sky Survey (POSS) plates by eye, with each of the groups satisfying specific criteria. Firstly, the groups needed to be compact and isolated. Specifically, there needed to be at least four galaxies within three magnitudes of the brightest galaxy, all within a circle of angular radius θ_G , the group radius. The isolation criteria specified that there should be no galaxies within three magnitudes of the brightest galaxy within $3\theta_G$ of the group centre. The magnitude concordance was used to select physically associated groups without any redshift information, while the isolation criteria ensured that the cores of clusters, and associations in clusters were not identified as compact groups.

P94 selected groups in the same way as H82, except that the search was on COSMOS scans of plates taken with the United Kingdom Schmidt Telescope (UKST), and the search algorithm was machine-implemented.

The MIRIAD task MBSPECT was used to inspect the HIPASS spectrum for each group, which was made by averaging over a region of 5×5 pixels, centered on the optically determined position of the group. Each HIPASS pixel is 4 arcmin on a side. This $20' \times 20'$ averaging area ensures that all H I flux coming from the groups is accounted for, as the largest group in the sample has an optical diameter of $10.8'$.

The spectra were examined by eye for emission lines, and groups with evidence of emission were selected for reobservation. The HCG spectra were searched at each group’s known optical velocity (Hickson 1993). The SCGs

however, do not have comprehensive velocity information for all their member galaxies, and thus the entire velocity range of HIPASS was examined at each SCG position on the sky.

In total, 19 HCGs and 21 SCGs showed possible H_I emission. Twelve HCGs and ten SCGs were clear non-detections. The primary sample for this paper is shown in Table 1.

Table 1. Basic observational data for the compact group sample and pointed H_I observations using the MX mode. Descriptions of the quantities are given in the text

Object	RA (J2000)	Dec (J2000)	Integration time [min]	RMS noise [mJy]	Peak flux [mJy]	Mass limit [10 ⁹ M _⊙]
HCG 003	00:34:27	−07:35:32	14	6.6		3.50
HCG 004	00:34:15	−21:26:47	14	5.7	44.1	3.69
HCG 006	00:39:10	−08:23:43				7.43
HCG 007	00:39:23	+00:52:40	14	8.4	44.5	1.14
HCG 014	01:59:47	−07:01:42	14	6.9		1.64
HCG 016	02:09:33	−10:09:47	14	6.7	117.3	0.86
HCG 019	02:45:45	−12:24:42	28	7.3	24.6	1.02
HCG 021	02:45:17	−17:37:09	14	10.4	21.1	2.94
HCG 022	03:03:31	−15:40:32	14	7.3	67.6	0.36
HCG 023	03:07:06	−09:35:08	14	7.6	36.1	1.34
HCG 024	03:20:18	−10:51:53				4.78
HCG 025	03:20:43	−01:03:07	14	7.4	34.9	2.22
HCG 026	03:21:54	−13:38:45	14	6.4	30.3	5.04
HCG 028	04:27:19	−10:19:00				7.27
HCG 030	04:36:28	−02:49:57				1.18
HCG 031	05:01:36	−04:15:24	14	6.6	170.6	0.50
HCG 040	09:38:54	−04:51:07				2.45
HCG 042	10:00:21	−19:38:57				0.93
HCG 043	10:11:13	−00:01:54				5.52
HCG 048	10:37:45	−27:04:50	28	7.4	19.2	0.41
HCG 048/1	10:37:45	−27:04:50		7.4	18.9	
HCG 048/2	10:37:45	−27:04:50		7.4	19.2	
HCG 062	12:53:08	−09:13:27				1.06
HCG 063	13:02:08	−32:46:04	14	7.0	30.0	5.35
HCG 064	13:25:43	−03:51:28				6.55
HCG 067	13:49:03	−07:12:20	14	7.4	23.4	3.19
HCG 086	18:47:05	−30:49:33	14	5.6		2.08
HCG 087	20:48:11	−19:50:35				4.43
HCG 088	20:52:22	−05:45:28	14	12.0	52.2	2.14
HCG 089	21:20:10	−03:54:31	14	14.0	71.8	4.59
HCG 090	22:02:05	−31:58:00				0.40
HCG 091	22:09:10	−27:47:45	14	20.2		3.05
HCG 097	23:47:22	−02:19:34				2.59
HCG 098	23:54:12	+00:22:24				3.68
SCG 06	01:49:10	−27:06:42	14	6.3		
SCG 07	00:37:43	−33:41:24	28	4.8	33.0	4.62
SCG 14	01:56:28	−20:05:51	14	6.8		
SCG 15	00:32:41	−25:36:47	14	7.1	40.4	0.50
SCG 17	00:51:14	−32:25:59	14	8.7		
SCG 18	01:18:32	−17:02:24	14	6.4		
SCG 20	00:14:56	−24:05:25	70	7.5	39.0	3.38
SCG 24	00:19:14	−26:42:49	28	5.4		
SCG 25	02:53:11	−09:29:12	14	6.8		
SCG 27	03:03:07	−22:12:10	14	7.3		
SCG 28	00:47:32	−21:29:16	14	5.8	37.6	2.41
SCG 33	03:04:51	−12:04:17	28	6.2		
SCG 34	01:19:27	−17:25:16	14	6.5		
SCG 35	03:37:06	−07:41:39	14	6.1		
SCG 39	00:37:09	−34:57:49	28	5.1		
SCG 43	01:15:57	−29:46:00	14	6.7	34.7	5.29
SCG 49	02:55:08	−21:35:43	14	6.1		
SCG 51	00:33:36	−27:47:00	14	6.5	132.2	0.14
SCG 54	00:05:58	−36:06:54	14	6.4		
SCG 55	00:58:58	−28:17:38	14	5.3		
SCG 57	01:49:57	−27:48:32	14	10.7		

3 H I Data

3.1 Follow-Up Observations

To confirm the H I properties of these groups, and to more accurately measure their masses, the 40 compact groups detected in HIPASS were reobserved with the Parkes Telescope Multibeam instrument during 2001 November, 2002 April, and 2003 January. These observations were carried out using the narrowband mode of the correlator (MX mode, 8 MHz bandwidth, 1024 channels), giving a velocity resolution of 2.0 km s^{-1} . Each group was observed for a total on-source integration time of 14 min, using the MX observing mode.

The MX observing mode places the central seven beams of the Multibeam on-source, in turn, for 2 min each. The 12 min that each beam spends off-source is used to measure the bandpass for subtraction from the on-source observation. The final spectrum is made by combining the seven separate bandpass-subtracted spectra.

The FWHM of each of the beams is $\sim 14'$, meaning each of the compact groups is entirely covered by the beam, when the beam is placed at the group centre. The largest group is HCG 21 with a diameter of $10.8'$, while the smallest groups are SCGs 54 & 57, with a diameter of $0.8'$. The median group diameter is $2.5'$. The diameter of

the groups is defined by the smallest circle which contains all the galaxies. The telescope was pointed at the centres of these circles.

3.2 Data Reduction

Initial reduction of the data used LIVEDATA (Barnes 1998), which subtracts the bandpass estimates from the on-source observations, and GRIDZILLA which then combines the seven MX observations into one. The reduced spectra are shown in Figure 1 for the detected compact groups. The velocity shown in the spectra in Figure 1 is optical velocity, in the heliocentric frame.

The compact group spectra were analysed and parameterised using the MIRIAD task MMSPECT, an extended version of MBSPECT. This task examines the combined spectra from the MX observations and calculates an estimate of the continuum baseline level using a Gaussian smoothing routine, which can then be subtracted. The baseline fit is shown in Figure 1 as a solid line. The width of the fitting Gaussian was always >40 channels to ensure that small-scale fluctuations do not affect the baseline calculation. The sources of the baseline instability after bandpass subtraction in LIVEDATA are solar interference which changes rapidly, and pointings near a strong

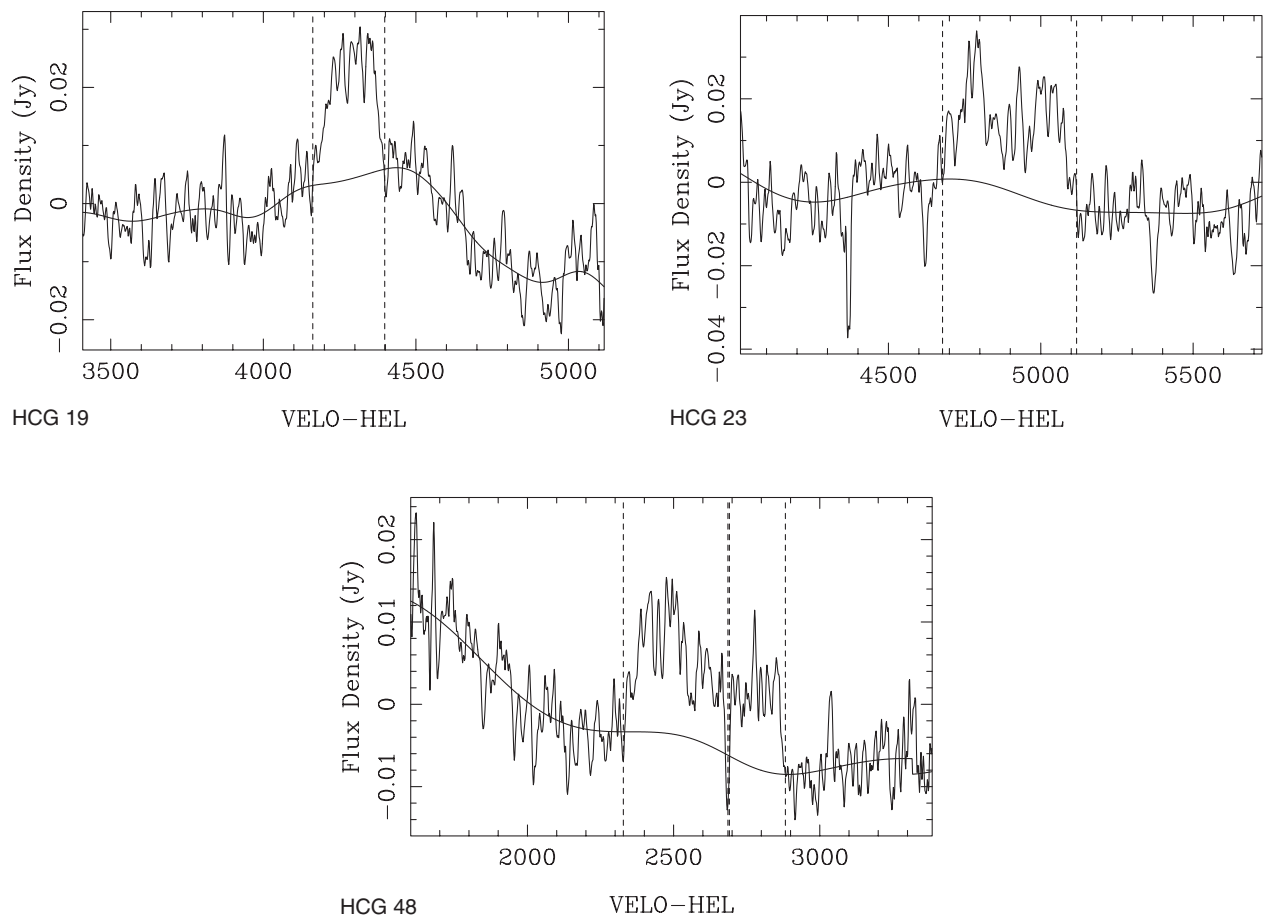


Figure 1 Some example narrowband spectra of the detected compact groups. The solid line in each spectrum is the estimated baseline, and the vertical dashed lines surround the emission region. The emission from HCG 48 has been split into two domains, separated by the area where the profile crosses the baseline, to provide a reliable estimate of the flux. More spectra are available as Accessory Material.

continuum source which creates standing waves between the dish surface and the receiver.

The channels within the profile window (marked in Figure 1 as dashed lines) were excluded from the baseline fit, and the baseline has been interpolated across this window. The profile window was determined so that the sides of the window cut the baseline at the edges of the profile, and within the window the profile does not cross the baseline.

It is assumed that the H I in each group is concentrated at the centre of the observation, and that it is optically thin. This makes the H I masses presented here lower limits. The H I mass of a galaxy is given by Equation (1):

$$M_{\text{H I}}[M_{\odot}] = 2.36 \times 10^5 \left(\frac{v^2}{H_0} \right) S_{\text{int}} \quad (1)$$

where v is the recessional velocity of the galaxy and S_{int} is the integrated flux from the galaxy. If a correction for the location of the H I in the beam were to be applied, the maximum correction needed can be calculated if all the H I is considered to be at the edge of the group diameter, and thus farthest away from the centre of the beam as possible. Since the largest group diameter is $10.8'$, and the FWHM of the observing beams is $14.3'$, the maximum possible flux correction needed would be 48% upwards.

Seven groups have had multiple observations combined to produce the final spectrum. These observations were combined using the MIRIAD task IMCOMB. Two observations were combined for HCG 19, HCG 48, SCG 7, SCG 24, SCG 33, and SCG 39, and five observations were combined for SCG 20.

The RMS noise level of the spectrum has been calculated from the non-profile channels. This calculation is performed before the baseline fit has been subtracted, so it includes the effect of the baseline ripple. It is much harder to find a signal superimposed on a varying baseline than it is on a flat one, and this measure quantifies this difference. The RMS noise level is given in Table 1.

3.3 H I Observational Parameters

Table 1 gives the basic H I observational data for the compact group sample. ‘Integration time’ gives the total narrowband integration time on the group. ‘RMS noise’ gives the RMS noise level (per smoothed channel) of the narrowband observation. If data in this column is not listed, the group has not been detected in the HIPASS data, and was not reobserved. ‘Peak flux’ gives the observed peak flux of the H I emission from the group. If data in this column is not listed, the group has not been detected by the narrowband observations. ‘Mass limit’ gives the observable H I mass limit per HIPASS spectral channel, obtained using the limit given in Section 3.3.2, and the known velocity of the group (Hickson 1993). For the SCGs, this value is calculated using the observed velocity. If an SCG has not been detected, this column is left blank.

Table 2 gives the H I profile data for each of the groups. ‘Integrated flux’ gives the integrated flux of the detected profile. ‘Mean velocity, heliocentric’ gives the observed

intensity-weighted mean velocity of the profile, computed in the following way:

$$\bar{v} = \frac{\sum I \times v}{\sum I} \quad (2)$$

where \bar{v} is the mean velocity of the profile, and I is the intensity of the emission at velocity v . ‘Mean velocity, GSR’ shows this velocity after correction to the Galactic Standard of Rest (GSR) frame, as used by Braun & Burton (1999):

$$\bar{v}_{\text{GSR}} = \bar{v} + 9 \cos(l) \cos(b) + 12 \sin(l) \cos(b) - 7 \sin(b) + 220 \sin(l) \cos(b) \quad (3)$$

where l and b are the galactic latitude and longitude of the group respectively, and \bar{v} is in the heliocentric frame. The uncertainties of the ‘Mean velocity’ columns are the same, but only the latter is shown with errors because the GSR velocities will be used from now on. The ‘Width’ columns give the maximum 50% and 20% widths of the emission profile. Maximum widths are determined from the two points at the appropriate fraction of the peak flux farthest away from each other in the emission profile. ‘Opt. group, vel. disp.’ gives the optical radial velocity dispersion of the group, which is the RMS of the galaxy velocities with respect to the velocity centroid (Hickson 1993).

A comparison of the H I fluxes measured by this study, WR87, and H97 is shown in Figure 2. Only the Hickson compact groups listed in this paper are compared in Figure 2. Good agreement is seen with the fluxes of WR87, whereas for H97 there is no correlation. Between H97 and WR87, there are four groups which agree well — HCGs 19, 22, 26, and 67. Three of these groups (HCG 19, 26, 67) have a reasonable agreement to the narrowband observations, within the WR87 uncertainties. WR87 used the NRAO 91-m and the Arecibo 305-m telescopes for their observations, while H97 used the Effelsberg 100-m telescope. There is no clear reason why the agreement to the observations of WR87 is better than to those from H97. The upper limits found by HIPASS are generally higher than those determined by WR87 and H97, as expected.

3.3.1 Errors in Measured Parameters

The uncertainties for the 50% profile width and the 20% width are given by the semi-empirical relation found by Schneider et al. (1986):

$$\Delta v(f) = 1.5\sqrt{2} |v_{80} - v_{20}| \left(\frac{0.25}{f(1-f)} \right) (S/N)^{-1} \quad (4)$$

where $\Delta v(f)$ is the uncertainty in the profile width at $(100f)\%$ of the peak flux (the 50% width has $f = 0.50$), v_{80} is the 80% width of the profile, v_{20} is the 20% width, and (S/N) is the signal-to-noise ratio of the spectrum. The latter can be estimated by the following equation:

$$(S/N) \equiv \frac{S_{\text{int}}}{\Delta S_{\text{int}}} = \frac{S_{\text{int}}}{\sigma v_{50}} \sqrt{\frac{v_{50}}{13.2}} \quad (5)$$

where S_{int} is the integrated flux of the profile, σ is the RMS noise level, v_{50} is the 50% width of the profile,

Table 2. Measured H I parameters for the compact groups detected by the pointed H I observations. Descriptions of the quantities in each column are given in the text

Object	Integrated flux [Jy km s ⁻¹]	Mean velocity, heliocentric [km s ⁻¹]	Mean velocity, GSR [km s ⁻¹]	50% Width [km s ⁻¹]	20% Width [km s ⁻¹]	Opt. group, vel. disp. [km s ⁻¹]
HCG 003			7883			251.2
HCG 004	3.2 ± 0.1	8058	8092 ± 2	83 ± 3	111 ± 5	338.9
HCG 006			11480			251.2
HCG 007	7.6 ± 0.6	4404	4506 ± 12	222 ± 24	297 ± 38	89.1
HCG 014			5393			331.1
HCG 016	26.5 ± 0.5	3888	3900 ± 5	228 ± 11	369 ± 17	123.0
HCG 019	3.4 ± 0.3	4286	4252 ± 9	167 ± 19	213 ± 29	
HCG 021	2.5 ± 0.6	7262	7228 ± 6	189 ± 11	201 ± 18	112.2
HCG 022	8.0 ± 0.2	2574	2533 ± 1	137 ± 3	160 ± 5	43.7
HCG 023	7.7 ± 1.0	4901	4875 ± 25	335 ± 50	429 ± 79	169.8
HCG 024			9209			199.5
HCG 025	7.0 ± 0.7	6285	6273 ± 9	283 ± 18	323 ± 28	61.7
HCG 026	9.0 ± 1.2	9506	9458 ± 9	440 ± 18	471 ± 28	199.5
HCG 028			11357			85.1
HCG 030			4577			72.4
HCG 031	23.9 ± 0.3	4057	3968 ± 2	125 ± 5	222 ± 7	85.1
HCG 040			6599			147.9
HCG 042			4069			213.8
HCG 043			9893			223.9
HCG 048	5.1 ± 1.7		2707 ± 24	454 ± 77	514 ± 116	302.0
HCG 048/1	3.4 ± 0.8	2497	2571 ± 22	290 ± 45	337 ± 70	
HCG 048/2	1.7 ± 0.4	2787	2861 ± 5	164 ± 11	177 ± 16	
HCG 062			4327			288.4
HCG 063	5.0 ± 0.6	9312	9479 ± 11	260 ± 21	301 ± 33	131.8
HCG 064			10785			213.8
HCG 067	3.5 ± 0.4	7731	7865 ± 23	186 ± 47	285 ± 73	208.9
HCG 086			6073			269.2
HCG 087			8868			120.2
HCG 088	5.9 ± 0.7	6015	6161 ± 4	168 ± 8	188 ± 12	26.9
HCG 089	8.7 ± 0.6	8872	9023 ± 8	154 ± 16	211 ± 24	55.0
HCG 090			2672			100.0
HCG 091			7353			182.0
HCG 097			6785			371.5
HCG 098			8083			120.2
SCG 06						
SCG 07	13.4 ± 1.5	8988	9054 ± 26	619 ± 51	724 ± 80	
SCG 14						
SCG 15	7.5 ± 0.7	2934	2982 ± 4	270 ± 9	293 ± 13	
SCG 17						
SCG 18						
SCG 20	9.4 ± 1.1	7680	7740 ± 19	358 ± 38	438 ± 60	
SCG 24						
SCG 25						
SCG 27						
SCG 28	3.9 ± 0.2	6479	6540 ± 8	139 ± 16	211 ± 24	
SCG 33						
SCG 34						
SCG 35						
SCG 39						
SCG 43	1.8 ± 0.1	9654	9691 ± 3	60 ± 6	100 ± 10	
SCG 49						
SCG 51	21.6 ± 0.2	1539	1581 ± 4	110 ± 7	341 ± 11	
SCG 54						
SCG 55						
SCG 57						

and the reduced velocity resolution of the narrowband observations is 13.2 km s⁻¹. The inverse of Equation (5) can be used as the uncertainty estimate of the integrated flux.

It is necessary to change Equation (4) to take the 50% width, instead of the 80% width. Thus v_{80} is replaced by v_{50} , and the factor of 3.0 is twice the original value because $|v_{80} - v_{20}|$ is twice $|v_{50} - v_{20}|$, assuming the profile sides

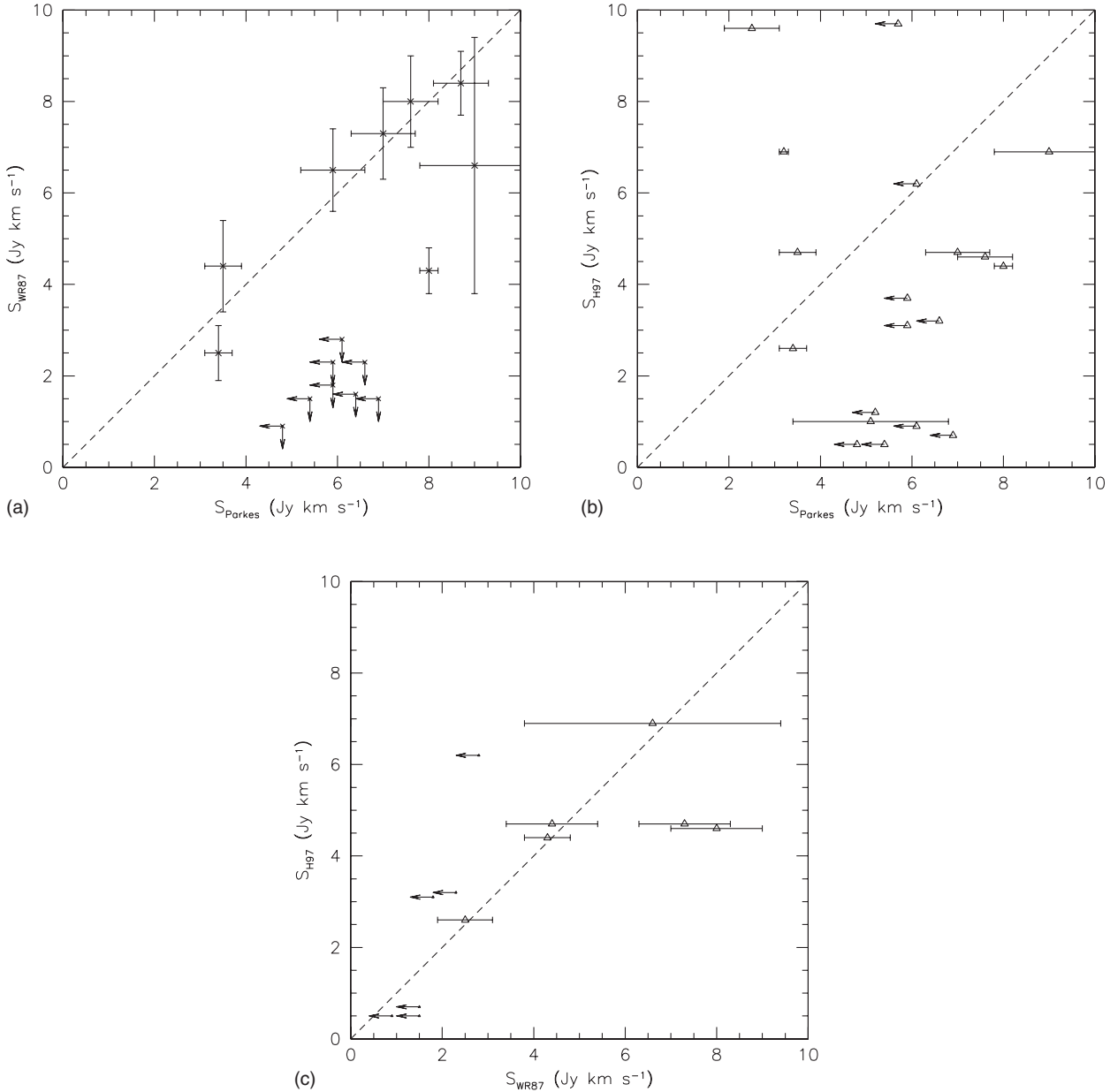


Figure 2 The fluxes of the compact groups studied in this sample compared to the measurements of (a) WR87 and (b) H97. (c) The fluxes measured by WR87 and H97 compared to each other. The line of equality is shown as a dashed line. No flux errors or upper limits were listed in H97. Upper flux limits are indicated by arrows. Three WR87 detections are not shown on these plots — HCG 16, HCG 23, and HCG 31 — which lie above the axis limits.

are straight. Equation (4) therefore becomes:

$$\Delta v(f) = 3.0\sqrt{2}|v_{50} - v_{20}|\left(\frac{0.25}{f(1-f)}\right)(S/N)^{-1} \quad (6)$$

The error in the mean velocity can be estimated in a similar way (Schneider et al. 1986), and is given by:

$$\Delta \bar{v} = \frac{3.0}{\sqrt{2}}|v_{50} - v_{20}|(S/N)^{-1} \quad (7)$$

where $\Delta \bar{v}$ is the uncertainty on the intensity-weighted mean velocity.

3.3.2 HIPASS Detection Limits

The detection of an object depends not only on its flux, but also on the number of channels over which the flux is spread. The detection experiment of the group sample was done using HIPASS data, thus to find the upper limits for the non-detected groups knowledge of the HIPASS detection limit is needed. This can be deduced from the integrated flux S_{int} and the 50% velocity width v_{50} of the 4315 galaxies in the HIPASS catalogue (Meyer et al. 2004). For a detection, the signal-to-noise ratio over the width of the galaxy profile needs to be greater than a specified level, which is taken as the detection limit. The mean single-channel RMS noise level of HIPASS is $\sigma = 13$ mJy.

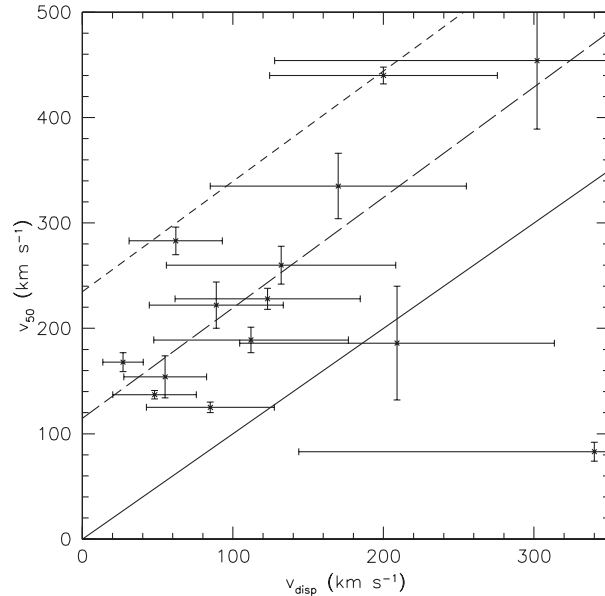


Figure 3 The 50% H I velocity width v_{50} is plotted against the optical velocity dispersion of the group for the detected compact groups. The long-dashed line is the least squares fit to the data, the short-dashed line is the fit offset to become an upper limit, and the solid line is $v_{50} = v_{\text{disp}}$.

In the HIPASS catalogue, most galaxies are detected with $\text{SNR} > 5\sigma$, and there is a sharp dropoff in detections at lower signal-to-noise. The approximate HIPASS detection limit is thus:

$$S_{\text{int}} \gtrsim 5 \frac{\sigma}{\sqrt{n}} v_{50} \quad (8)$$

where n is the number of uncorrelated channels defining the profile. An estimate of n is $v_{50}/18.0$ since the velocity resolution of HIPASS is 18.0 km s^{-1} . Here any channel-to-channel correlations are ignored, and normal statistics is assumed.

The upper mass limit for the non-detected galaxies is then:

$$M_{\text{H I}} \lesssim (2.36 \times 10^5 D^2) (5\sigma \sqrt{18.0}) \sqrt{v_{50}} M_{\odot} \quad (9)$$

where D is the distance to the group [Mpc], calculated using the median of the corrected optical velocities of the group members.

To estimate the upper H I mass limits for the non-detected groups requires that an estimate of the total velocity width of the group be made. If the inclinations of each galaxy and their optical rotation speeds were known for all the groups, and the galaxies which were likely to have H I could be identified, then estimating v_{50} would be straightforward. However, the only related quantity available for the HCGs is the optical velocity dispersion.

A plot of v_{50} versus the optical radial velocity dispersion v_{disp} for each of the detected compact groups is shown in Figure 3. It shows that v_{disp} is in practice, a useful lower limit on the H I velocity width. This is expected if the group is unresolved spectrally in H I. To get the upper mass limit

requires a upper limit on v_{50} . This limit is obtained by calculating the least squares fit to the data in Figure 3 and shifting it upwards to make it an upper bound. The group HCG 4 (lower right corner) has been left out of the regression calculation as it clearly does not fit with the general trend. It is conjectured that in HCG 4, only the central face-on spiral HCG004a is detected in H I, as the other two members of the group are ellipticals.

The equation for the line which represents a reasonable upper limit on the width of the H I profile, given the optical velocity dispersion is:

$$v_{50} = 1.05 v_{\text{disp}} + 230 \text{ km s}^{-1} \quad (10)$$

The upper mass limits for the non-detected HCGs have been calculated using Equation (10), and are shown in the second column of Table 5.

4 Predicted H I Masses

The null hypothesis is that the H I detected in a compact group is the sum of the H I in the individual member galaxies, if they were field galaxies with the same optical properties. A test of the null hypothesis requires a reliable method of calculating a galaxy's likely H I mass, depending on its type and size. In this Section, three mass estimation methods are described.

Two of the methods — the mass–diameter relation and the mass–luminosity relation — are described by Haynes & Giovanelli (1984, hereafter HG84), who examine a sample of isolated field galaxies to investigate whether correlations exist between each galaxy's H I mass, and its optical properties. The correlations with the optical diameter of a galaxy and with the blue luminosity are also examined.

The isolated galaxy sample of HG84 was extracted from the Catalogue of Isolated Galaxies (CIG) by Karachentseva (1973). This catalogue contained 1052 galaxies with a magnitude limit of $+15.7$ in the POSS red prints. The galaxies in the catalogue were determined to be isolated on the following basis. If a galaxy of diameter d has a neighbour galaxy of diameter d_1 , with the constraint $\frac{1}{4}d \leq d_1 \leq 4d$, then the first galaxy is isolated from the second if it lies more than $20d_1$ away from it. The CIG contains galaxies which are isolated from all other galaxies according to this metric. The HG84 sample contains only those galaxies from the CIG which are also included in the Uppsala General Catalog of Galaxies (UGC; Nilson 1973), and lie at a declination which can be observed by the Arecibo 305-m telescope ($-1^\circ < \delta < +38^\circ$). This leaves a sample of 324 isolated galaxies.

The third mass estimation method is that used by WR87 and H97, and is a simplified version of the mass–luminosity relation of HG84.

The application of each of these methods to this sample is discussed in turn.

4.1 Optical Diameter as an Indicator of H I Mass

HG84 report a strong correlation between the optical major diameter of a galaxy and the galaxy's H I mass,

Table 3. Coefficients of the mass–diameter relation (MDR) from Haynes & Giovanelli (1984) and Chamaraux, Balkowski, & Fontanelli (1986)

Type t	Morphology	c_1	c_2	SEE	Source
1	S0	5.68	1.00	0.25	C86
2	S0a	6.21	1.00	0.10	C86
3, 4	Sa, Sab	7.17	0.82	0.31	HG84
5	Sb	7.29	0.83	0.25	HG84
6	Sbc	7.27	0.85	0.17	HG84
7	Sc	6.91	0.95	0.18	HG84
8, 9	Scd, Sd, Irr, Sm, Sdm, dwarf sp	7.00	0.94	0.17	HG84
10	Pec	7.75	0.66	0.19	HG84

which takes the form:

$$\log(h^2 \langle M_{\text{H I}} \rangle) = c_1(t) + c_2(t) \log(h D_I)^2 \quad (11)$$

where $\langle M_{\text{H I}} \rangle$ is the expected H I mass [M_\odot], c_1 and c_2 are scalar parameters which depend on the galaxy's numerical morphological type t , D_I is the length of the major axis of the galaxy [kpc], and h is the Hubble factor such that $H_0 = 100h \text{ km s}^{-1} \text{ Mpc}^{-1}$. The numerical morphological type ranges from 0 to 10, with later type galaxies having a higher numerical type (Sandage & Tammann 1981). That is, ellipticals have a type $t = 0$, and peculiar galaxies are $t = 10$. Table 3 lists the numerical types associated with galaxy morphology. This system differs from that of the definition given in the RC2 (de Vaucouleurs 1976), which has $t = -5$ for ellipticals, $t = -2$ for S0, $t = 0$ for S0a, and increases to $t = 10$ for irregular galaxies. The system of Sandage & Tammann (1981) is used here because it matches with the system used by HG84.

Equation (11) reduces to:

$$\langle M_{\text{H I}} \rangle = \frac{1}{h^2} 10^{c_1(t)} (h D_I)^{2c_2(t)} \quad (12)$$

Equation (12) states that a galaxy has a well-defined surface density of H I dependent on its morphological type, which is expressed in the constant $10^{c_1(t)}$. However, there is an additional dependence on the diameter of the galaxy, and this is expressed by the power $c_2(t)$.

HG84 determined the values of $c_1(t)$ and $c_2(t)$ for the various morphological types from the isolated sample of 324 galaxies after observations with the Arecibo 305-m telescope and the Green Bank 91-m telescope. The rate of non-detections for the morphological types 0, 1, & 2 was very high in HG84, with only 14 of 30 galaxies of these types detected. The coefficients were calculated based only on the detected galaxies, so they are biased towards higher mass galaxies. To remove this bias, the results of Chamaraux et al. (1986, hereafter C86) are used here for galaxies of type 1 and 2. For elliptical galaxies (numerical type 0), the H I mass is assumed to be zero, as only 6% of RC3 ellipticals have been detected in HIPASS (Sadler et al. 2002). The low H I detection rate of ellipticals makes it impossible to reliably estimate the range of H I masses that these galaxies would have.

Table 4. Coefficients of the mass–luminosity relation (MLR) from Haynes & Giovanelli (1984)

Type t	c_3	c_4	SEE	Source
1, 2	2.93	0.63	0.36	HG84
3, 4	2.87	0.66	0.38	HG84
5	2.99	0.66	0.36	HG84
6	2.77	0.69	0.29	HG84
7	3.03	0.66	0.26	HG84
8, 9	1.40	0.84	0.30	HG84
10	4.99	0.45	0.25	HG84

The c_1 and c_2 values are shown in Table 3, along with the standard error of the estimate. The standard error of the estimate (SEE) is a measure of how far away a data point is from its predicted value, and is defined as:

$$\text{SEE} \equiv \sqrt{\frac{\sum (y - y')^2}{N}} \quad (13)$$

where y is the known value of some data point, y' is the value predicted by some function which fits the data point, and N is the number of data points that the function is fitted to.

C86 take into account the upper limits of detection, using the method of Chamaraux (1987). This method uses the mean H I surface density, similar to HG84, and thus the method for determining $\langle M_{\text{H I}} \rangle$ is the same. The values found by C86 are lower than those found by HG84 for the same morphological types. There was no attempt by C86 to determine the residual dependence on the galaxy's diameter, so $c_2 = 1$ for these types.

4.2 Blue Luminosity as an Indicator of H I Mass

The distance-independent relationship between the H I mass of a galaxy and its blue luminosity L_B can be expressed as:

$$\log(h^2 \langle M_{\text{H I}} \rangle) = c_3(t) + c_4(t) \log(h^2 L_B) \quad (14)$$

which, like the relation between diameter and mass, can be written as:

$$\langle M_{\text{H I}} \rangle = \frac{1}{h^2} 10^{c_3(t)} (h^2 L_B)^{c_4(t)} \quad (15)$$

Table 4 gives the values for the constants in Equation (15). We can see that, in comparison to the MDR, the errors involved with the MLR are slightly higher; on average there is a greater dispersion in the relationship between the predicted and measured values of H I mass when using Equation (15). The MLR was not investigated by C86, thus all values in Table 4 are from HG84. As HG84 again did not take into account the non-detection rates for galaxies with early type morphology, the values quoted for types $t = 1$ & 2 are most likely overestimated.

4.3 The Simplified Mass–Luminosity Relation

A simplified version of the MLR was used by both WR87 and H97. For galaxies with type $t = 0, 1$, & 2, H97 quotes

the values $\langle M_{\text{H I}} \rangle / L_{\text{B}} = 0.03$, and $\langle M_{\text{H I}} \rangle / L_{\text{B}} = 0.65$ for types $t > 2$. Since a dependence on luminosity was not considered, $c_4 = 1$, and thus Equation (15) becomes $\langle M_{\text{H I}} \rangle / L_{\text{B}} = 10^{c_3(t)}$. Thus $c_3 = -1.52$ for galaxies with types $t = 1$ & 2 (early types), and $c_3 = -0.19$ for later types. This relation will be referred to as the simplified mass–luminosity relation (SMLR).

4.4 Compact Group H I Content

For each of the HCGs, Hickson (1993) has classified the members by galaxy type. The expected H I content of each group was calculated by summing the expected $M_{\text{H I}}$ for each member's morphological type. This calculation is repeated here for each of the three methods described above.

Table 5 shows the H I masses calculated with each of the methods described above, as well as the observed H I mass. 'H I mass' gives the H I mass of the compact group. If the group has not been detected in these observations, an upper mass limit is given. This limit is obtained using the HIPASS detection limit, which is detailed in Section 3.3.2. It would be inappropriate to use the limits from the follow-up survey, because the *detection experiment* was done using the HIPASS data. The 'Expected mass' columns give the expected masses of the groups calculated using the MDR, MLR, and SMLR. The errors on these estimates are calculated from the quoted standard errors listed for each of the methods.

4.4.1 H I Content of the Compact Groups

The plot of expected mass versus observed mass for the HCGs is given in Figure 4 for each of the estimation methods described above. We no longer include the SCG detections, because without full knowledge of the optical properties of the sample as a whole, including them would only add noise to the analysis. Each plot in Figure 4 has an area bounded by long-dashed lines. These lines represent the 3σ errors for each of the estimation techniques. For a galaxy whose expected H I mass was estimated using the MDR technique for example, there is a range of observable masses for that galaxy which could be considered consistent with that estimate, because the spread of the observed galaxy masses in HG84 is broad. The long-dashed lines in Figure 4 show the maximum and minimum observable H I masses consistent with each expected mass, averaged over all morphological types. For (a) & (b) the mass uncertainties are obtained from HG84, and for (c) from H97.

The traditional way of assessing the relative H I content of a sample of galaxies to the field is to use the deficiency parameter, which is defined in the following way:

$$\text{DEF}_{\text{H I}} \equiv \log[M_{\text{H I, pred}}] - \log[M_{\text{H I, obs}}] \quad (16)$$

where $M_{\text{H I, pred}}$ is the predicted average H I mass of a field galaxy with the same optical properties as the sample galaxy, and $M_{\text{H I, obs}}$ is the observed H I mass of the galaxy. If $\text{DEF}_{\text{H I}} \leq 0$ then the galaxy would not be considered

deficient, while a $\text{DEF}_{\text{H I}} > 0$ would indicate a H I deficiency. The $M_{\text{H I, pred}}$ has an uncertainty associated with it due to the natural spread of masses seen in the field, so the deficiency parameter will also have an uncertainty.

Verdes-Montenegro et al. (2001) calculated the average deficiency of their sample of 50 HCGs to be $\langle \text{DEF}_{\text{H I}} \rangle = 0.40 \pm 0.07$. They considered a group to have an anomalous H I content when the amount predicted varied from the amount observed by twice the mean error on the predicted mass, otherwise the H I content is considered to be normal. In this way, only 23 of these 50 groups could be considered deficient, and these 23 groups have an average deficiency of 0.73. Of the other 27 groups, three have anomalously high H I masses and the other 24 are in the range considered normal by Verdes-Montenegro et al. (2001). Thus an average measure of deficiency does not give an indication of the properties of the sample as a whole. A further 14 groups were not detected by Verdes-Montenegro et al. (2001) and could not be included in the analysis.

To properly measure the correlation between the expected mass and the observed mass, a technique that takes into account the upper limits of the data is required. The Buckley–James method (Buckley & James 1979; Isobe, Fiegelson, & Nelson 1986) uses the Kaplan–Meier estimator and the EM (expectation, maximization) algorithm to find a linear regression fit to a data set which has upper limits. The EM method determines a fit using estimates of the values of the upper limits (censored points) and changes these estimates iteratively to maximise the likelihood estimators for the unknown parameters. To estimate the values of the censored points, knowledge of the distribution around the regression lines is needed. The Kaplan–Meier estimator uses the known values for the uncensored data to estimate this distribution, and hence give values to the censored points. Simulations by Buckley & James (1979) showed that this estimator performed well even if 50% of the points were censored unevenly along the distribution, finding the true value of the slope with a small number ($n = 20$) of points.

The fits for each of the methods are shown in Table 6, and the coefficients are for the relation in Equation (17):

$$\log \langle M_{\text{H I}} \rangle = a \log M_{\text{H I}} + b \quad (17)$$

The estimate of the scatter σ is given by the relation:

$$\sigma = \sqrt{\sum_{i=1}^n \frac{(y_i - \langle y_i \rangle)^2}{n - 2}} \quad (18)$$

where y_i is the actual value of point i , $\langle y_i \rangle$ is the expected value of that point given by the fit, and n is the total number of points contributing to the fit.

When only the detected groups are considered, the least-squares fits change only slightly, as shown in Table 7 (see Equation (17)).

Table 5. Predicted H I masses for each compact group using the phenomenological relationships of HG84 and H97. Details of each of the mass estimates are described in the text

Object	H I mass [$10^9 M_{\odot}$]	Expected mass (MDR) [$10^9 M_{\odot}$]	Expected mass (MLR) [$10^9 M_{\odot}$]	Expected mass (SMLR) [$10^9 M_{\odot}$]
HCG 003	≤ 18.3	5.7 ± 0.8	10.7 ± 2.0	8.4 ± 3.9
HCG 004	10.1 ± 0.3	14.5 ± 2.5	23.3 ± 6.8	48.3 ± 24.1
HCG 006	≤ 38.9	7.9 ± 1.7	7.9 ± 1.7	22.6 ± 9.1
HCG 007	7.4 ± 0.6	13.3 ± 2.1	22.3 ± 4.0	34.2 ± 11.8
HCG 014	≤ 9.3	8.3 ± 1.4	7.3 ± 2.1	9.8 ± 4.4
HCG 016	19.4 ± 0.4	6.3 ± 1.3	8.7 ± 2.1	36.8 ± 13.8
HCG 019	3.0 ± 0.3	5.9 ± 0.8	3.8 ± 0.7	5.1 ± 1.7
HCG 021	6.3 ± 1.4	27.1 ± 5.3	14.3 ± 3.2	67.0 ± 22.9
HCG 022	2.5 ± 0.1	0.1 ± 0.0	0.4 ± 0.1	0.1 ± 0.0
HCG 023	8.8 ± 1.1	6.8 ± 1.5	4.1 ± 0.8	6.9 ± 2.7
HCG 024	≤ 23.6	1.9 ± 0.5	5.1 ± 1.5	19.9 ± 9.8
HCG 025	13.3 ± 1.4	9.6 ± 1.5	7.3 ± 1.4	20.8 ± 7.3
HCG 026	38.8 ± 5.3	9.6 ± 1.3	7.1 ± 1.1	11.1 ± 3.5
HCG 028	≤ 30.6	10.6 ± 2.6	7.2 ± 2.2	9.7 ± 4.7
HCG 030	≤ 4.9	2.1 ± 0.3	4.6 ± 0.9	3.3 ± 1.2
HCG 031	18.1 ± 0.2	2.1 ± 0.3	7.9 ± 2.0	15.9 ± 6.8
HCG 040	≤ 11.3	7.9 ± 1.5	7.0 ± 2.0	8.2 ± 3.8
HCG 042	≤ 4.7	0.4 ± 0.1	1.7 ± 0.5	0.5 ± 0.2
HCG 043	≤ 28.0	19.4 ± 2.5	22.3 ± 4.5	42.2 ± 13.5
HCG 048	1.8 ± 0.4	0.6 ± 0.1	1.0 ± 0.2	1.1 ± 0.6
HCG 048/1	1.1 ± 0.2			
HCG 048/2	0.7 ± 0.1			
HCG 062	≤ 5.7	1.7 ± 0.4	3.4 ± 0.9	1.3 ± 0.5
HCG 063	21.6 ± 2.6	17.7 ± 2.4	14.2 ± 2.8	48.1 ± 18.1
HCG 064	≤ 32.9	24.7 ± 3.1	14.0 ± 2.6	44.0 ± 17.1
HCG 067	10.4 ± 1.2	26.2 ± 5.7	11.0 ± 2.6	17.2 ± 6.2
HCG 086	≤ 11.1	0.4 ± 0.1	1.4 ± 0.4	0.3 ± 0.1
HCG 087	≤ 19.7	8.8 ± 1.5	7.9 ± 1.9	7.5 ± 3.2
HCG 088	10.8 ± 1.0	25.0 ± 3.0	18.2 ± 3.8	47.1 ± 13.2
HCG 089	34.1 ± 2.2	13.4 ± 1.4	10.7 ± 1.5	25.1 ± 7.1
HCG 090	≤ 1.7	7.7 ± 1.7	6.4 ± 1.7	30.4 ± 10.9
HCG 091	≤ 14.7	33.2 ± 4.1	17.9 ± 3.4	60.3 ± 23.3
HCG 097	≤ 15.2	8.2 ± 1.1	7.0 ± 1.4	12.7 ± 5.8
HCG 098	≤ 16.4	2.0 ± 0.5	5.0 ± 1.4	2.5 ± 1.1
SCG 06				
SCG 07	52.9 ± 6.1	19.7 ± 7.1	5.4 ± 8.2	4.8 ± 1.8
SCG 14				
SCG 15	3.2 ± 0.3	0.7 ± 0.2	2.1 ± 0.2	0.3 ± 0.1
SCG 17				
SCG 18				
SCG 20	27.1 ± 3.1	5.6 ± 2.9	2.1 ± 3.3	12.5 ± 6.2
SCG 24				
SCG 25				
SCG 27				
SCG 28	8.0 ± 0.4	6.7 ± 1.9	4.4 ± 2.5	3.7 ± 1.1
SCG 33				
SCG 34				
SCG 35				
SCG 39				
SCG 43	8.1 ± 0.3	7.8 ± 3.4	5.3 ± 3.8	15.5 ± 7.1
SCG 49				
SCG 51	2.6 ± 0.1			
SCG 54				
SCG 55				

4.4.2 Flux Contamination

Consideration must also be given to the possibility that the compact group observations have measured flux in the area surrounding the compact groups due to bright galaxies not

listed in the Hickson catalogue, or within the groups due to dwarf galaxies too faint to be considered members. The former case could occur if the angular extent of the group was less than a third of the beamsize. In this case, the

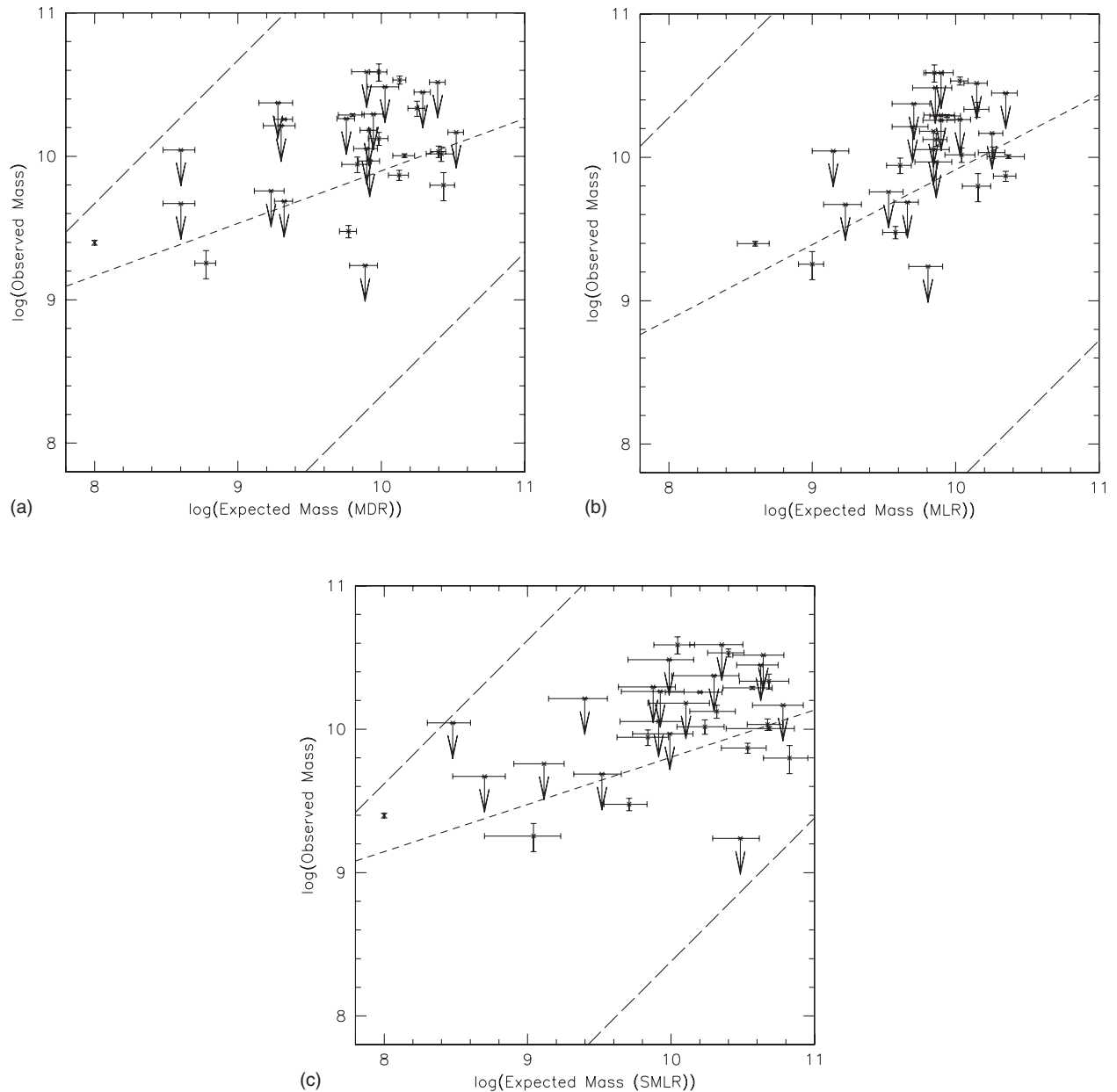


Figure 4 The observed mass is plotted against the expected mass for the (a) MDR, (b) MLR, and (c) SMLR. The long-dashed line shows the 3σ errors on each of the estimation techniques, as found by HG84 and H97, and are described further in the text. The short-dashed line is the Buckley–James fit to the data, including the non-detected groups. Upper limits for non-detections are indicated.

Table 6. The Buckley–James regression parameters for the observed compact group sample

	MDR	MLR	SMLR
a	0.36 ± 0.13	0.52 ± 0.17	0.33 ± 0.11
b	6.24	4.69	6.51
σ	0.32	0.31	0.31

Table 7. The least-squares regression parameters for the detected groups in the compact group sample

	MDR	MLR	SMLR
a	0.35 ± 0.13	0.49 ± 0.17	0.33 ± 0.11
b	6.55 ± 1.26	5.12 ± 1.68	6.65 ± 1.12
σ	0.32	0.31	0.31

H82 isolation criteria would allow bright interlopers to add H I flux to the observation. The latter case could occur at any time, as only galaxies within three magnitudes of the brightest galaxy in the group are counted as members. As dwarf galaxies can potentially have significant flux, if

they are not correctly accounted for, the content estimation may not be reliable.

The effect of any bright interlopers can be estimated by looking at the ratio of the group angular size to the beam size, and whether this ratio correlates with the ratio of

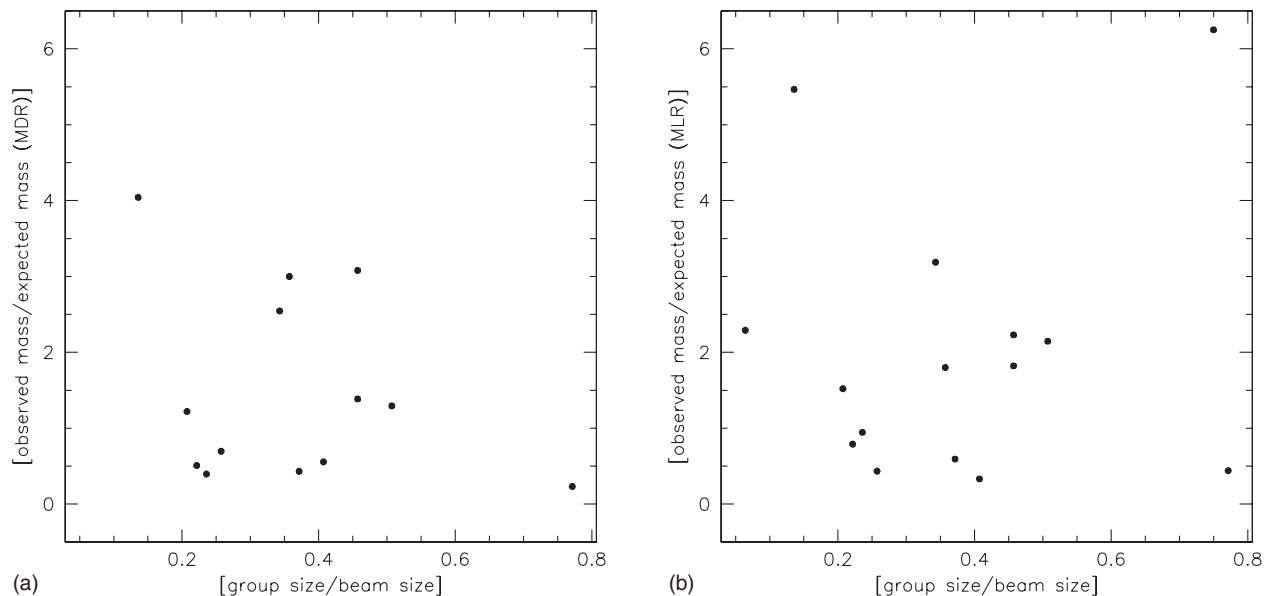


Figure 5 The observed-to-expected mass ratio is plotted against the fraction of the beam each compact group fills for (a) expected masses from the MDR and (b) expected masses from the MLR. This measures the effect interlopers outside the compact groups have on the observed masses.

the observed flux to the estimated flux. If, at beam-filling ratios below one-third, a larger observed mass is observed with no increase in expected mass, then bright interlopers may be the cause. The plots of beam-filling ratio versus the mass ratio are shown in Figure 5 for the MDR and the MLR.

The MDR plot of Figure 5 does not include an outlying group, HCG 26, to keep a reasonable scale for this plot. HCG 26 also appears anomalous in the MLR plot (topmost-right point). This group has many irregular galaxies and a diffuse dominant Scd which makes it difficult to estimate the diameters, and to a lesser extent the luminosities, of the group members. This anomalous group has a large angular extent however, and would not be affected by galaxies outside the group.

For beam-filling ratios of between 0 and 0.3, there is a chance galaxies outside the group may interfere. However, from Figure 5, there does not appear to be a significant difference in the mass ratios for the smaller groups when compared to the larger ones. There appear to be two anomalous groups using the MDR estimate and one with the MLR estimate. The compact group HCG 31 is only anomalous using the MDR mass estimate, and this may be because two galaxies in this group are aligned along the line-of-sight, making the estimation of their diameters difficult, while the sum of their luminosities is not affected. HCG 22 is an outlier using both mass estimators, and has about five times as much mass as expected. This group has several faint galaxies within the group radius which are not designated as group members due to the luminosity criteria. Also a large elliptical galaxy, and faint galaxies lie well outside the group boundaries but within the observing

beam. The H_i profile of HCG 22 is quite narrow though, so the velocity dispersion of the group must be low. It is not clear what is causing the underestimation of the mass of the group in this case.

The other four groups which are small enough to be affected by interlopers have the same observed-to-expected mass ratio as the larger groups, suggesting that interloper interference is not especially important. This may be because the beam would lessen the effect the contributions from outlying galaxies could have on the mass, or it may be that interlopers are not present. A search of the HIPASS Catalogue (Meyer et al. 2004) reveals that only two of these groups are catalogued from HIPASS, and both only have a single source within the observing beam.

An estimate of how many low-luminosity galaxies are present in the groups can be made using the compact group luminosity function, such as that found by Zepf, de Carvalho, & Ribero (1997). Zepf et al. (1997) surveyed 17 HCGs to look for galaxies which were not included in the H82 group membership, and derived a luminosity function from the results. The luminosity function had a Schechter form, with parameters $M_* = -19.5 + 5 \log h$ and $\alpha = -1.0$, where M_* is the ‘knee’ of the Schechter function and α is the slope of the function at the faint end.

The amount of mass a number of faint galaxies could contribute to the group can be estimated in the following way. To begin, the brightest galaxy which would not be classified as a group member must be three magnitudes fainter than the brightest galaxy in the group, or in terms of luminosity, 15 times fainter. The faint galaxy’s

contribution to the mass can be computed in the following way:

$$\begin{aligned}
 R &\equiv \frac{\langle M_{\text{H I, bright}} \rangle}{\langle M_{\text{H I, faint}} \rangle} \\
 &= \frac{(1/h^2)10^{c_3(t_1)}(h^2 \times 15 \times L_B)^{c_4(t_1)}}{(1/h^2)10^{c_3(t_2)}(h^2 \times L_B)^{c_4(t_2)}} \\
 &= 10^{c_3(t_1)-c_3(t_2)} \frac{(h^2 \times 15 \times L_B)^{c_4(t_1)}}{(h^2 \times L_B)^{c_4(t_2)}} \quad (19)
 \end{aligned}$$

Here R is the ratio of the predicted H I mass of the brightest galaxy in the group to the predicted H I mass of the brightest non-member galaxy, L_B is the blue luminosity of the brightest non-member galaxy, t_1 is the morphological type of the brightest member galaxy, and t_2 is the morphological type of the brightest non-member galaxy.

If we assume average values for c_3 and c_4 , Equation (19) reduces to $R = (15)^{\bar{c}_4}$, where \bar{c}_4 is an average value of c_4 and is independent of morphology. From Table 4, $\bar{c}_4 = 0.67$, thus $R \sim (15)^{0.67} = 6.1$. This means that the brightest non-member galaxy would be approximately six times less massive than the brightest member galaxy. If the brightest galaxy was not the only member galaxy contributing to the H I flux, the number of faint galaxies required to cause a significant increase in group mass would be large.

The compact group luminosity function of Zepf et al. (1997) suggests that the number of faint galaxies outnumber the bright galaxies by a factor of between 10 and 100. This provides enough faint galaxies to account for any large H I excesses. However, Zepf et al. (1997) also found that the faint galaxies have a wider spatial distribution than the bright galaxies. That is, the extent of the group as measured by faint galaxies is larger than the group as defined by H82 by a factor of several.

If it is now assumed that each group, as defined by the extent of the faint non-member galaxies, fills the HIPASS beam entirely, then groups with small angular sizes should again have larger H I excesses relative to larger groups. This effect is therefore confused with the interloper effect for groups smaller than one-third of the beam size, but would be the dominant effect in groups larger than this. Since there is no evidence that small groups have a greater excess than large groups (Figure 5), faint galaxies do not appear to have a significant effect on the observed mass for the compact group population as a whole. Again, the effect of the beam weighting the contributions from the H82 group more strongly than the surrounding area may be preventing faint galaxies from contributing.

5 Discussion

The purpose of these observations has been to test the null hypothesis that a compact group's H I mass can be calculated from the sum of the individual galaxies' masses, and that these galaxies have H I contents matching that of field galaxies with the same observed optical properties.

The Buckley–James fit, shown in Table 6 and in Figure 4, clearly does not agree with the null hypothesis for

any of the estimation methods. The slope of the line would suggest that groups with a high expected mass tend to be deficient in H I, while those with a low expected mass are not. However, this result is mostly due to the lack of sensitivity of HIPASS to low mass galaxies.

The least-squares fit parameters to the detected groups are, within errors, the same as the Buckley–James parameters. This is not surprising, because the Buckley–James fits are calculated using the distribution of the known points to estimate the upper limits.

These results show that the groups detected by these observations have H I contents similar to the reference field sample of HG84, primarily because all the detected groups lie within the area populated by the HG84 reference sample of field galaxies in the expected–observed mass plots. The slope of the line fitted to the detected groups does not match the expected slope, but this may be because the slope is being constrained only at the high expected–high observed mass end ($\log M_{\text{H I}} > 9.5$, $\log \langle M_{\text{H I}} \rangle > 9.5$).

Although the fits from each of the methods lie well within the sample region, because the majority of the groups are undetected by this survey, no definite conclusion about the H I content of the HCGs can be made. It seems that the limits of HIPASS are not low enough to properly assess the H I content of the lower mass compact groups.

It is interesting to determine what the detection limit needs to be to make a more definitive statement about the compact groups as a whole. As was mentioned before, the mass detection limit needs to be lowered to probe the low expected–low observed mass region. To move all the upper limits down to below the equality line would require an increase in sensitivity of 1.5 orders of magnitude, but to move half of them below the equality line only requires an increase in sensitivity of one order of magnitude.

If the undetected groups were detected with this sensitivity, then a much tighter constraint could be placed on the slope of the fit, and thus on the H I content of the groups. If they remained undetected, there would still be enough of a constraint at the low expected–low observed mass end to make an inference about the fraction of the population which could be considered to have normal H I content.

Practically, the extra sensitivity would require a detection limit of 0.5σ from HIPASS, which corresponds to 6.5 mJy. If this was the 5σ level, then the spectral RMS of the observation would need to be 1.3 mJy. To get this noise level would require a 100-fold increase in the integration time, if the same instrumental setup as HIPASS was used.

In the narrowband observations, a spectral RMS of ~ 7 mJy was obtained for most groups after 14 min of integration time. To get a RMS of 1.3 mJy would therefore take approximately 30 times longer, or about 6–7 h per source.

The narrowband sample can also be extended by incorporating the results of H97 to provide more confirmed detections. Of the 17 groups not detected here, H97 detects ten, but only the five detections with the highest signal-to-noise ratios are used here. The properties of these HCGs

Table 8. The H I properties of the HCGs not detected by HIPASS, but detected by H97

Name	RMS [mJy]	S_{int} [Jy km s ⁻¹]	v_{LSR} [km s ⁻¹]	v_{50} [km s ⁻¹]	v_{20} [km s ⁻¹]	$M_{\text{H I}}$ [10 ⁹ M _⊙]
HCG 003	2.3	6.2 ± 1.9	7986 ± 47	403 ± 93	1057 ± 146	18.9 ± 5.9
HCG 014	1.1	3.2 ± 0.4	5476 ± 2	1217 ± 3	1231 ± 5	4.6 ± 0.6
HCG 042	4.6	3.7 ± 1.3	3994 ± 18	201 ± 35	307 ± 55	2.8 ± 1.0
HCG 043	2.0	3.1 ± 0.7	9973 ± 8	352 ± 17	426 ± 26	14.9 ± 3.4
HCG 091	2.8	9.7 ± 1.9	7242 ± 9	483 ± 18	634 ± 29	24.4 ± 4.8

are listed in Table 8. When the H97 data is included, the detected groups all still lie within the HG84 sample region. In fact, the H97 detections lie only slightly below the estimated upper limits from HIPASS (see Figure 2). The resultant slope is slightly steeper though, with $a = 0.37$ for the MDR, $a = 0.55$ for the MLR, and $a = 0.35$ for the SMLR. Only one H97 detection (HCG 42) lies in the low expected–low observed mass region however.

Finally, this result can be contrasted to that found by Solanes et al. (2001) who studied the H I deficiency in spiral galaxies found in clusters. They found that within 1 Abell radius (1 R_A) of the cluster centre, spiral galaxies show strong deficiency, while in the outer parts of the cluster, the H I contents were similar to those found in the field. As the galactic densities of compact groups can approach that of the cores of rich clusters (H82), it is interesting then that compact groups do not show the same deficiency.

Solanes et al. (2001) concluded that hydrodynamical effects caused by the interaction of the H I with the hot intracluster medium in the centre of the clusters was the cause of the H I deficiency. They also found H I stripping was a relatively recent event (a few Gyr ago) for the deficient galaxies, as the optical properties of the galaxies have not been affected.

One possibility then for the normal H I content of the compact group sample investigated in this paper is the lack of a hot intragroup medium to strip the H I. This agrees with X-ray observations of the HCGs from Ponman, Bournier, & Ebeling (1996), who concluded that only 20% of groups had diffuse X-ray emission with luminosity greater than 10⁴² ergs. This luminosity is an order of magnitude weaker than the weakest cluster emission reported by Solanes et al. (2001).

6 Conclusions

- This study has shown that southern compact groups in the HCG catalogue appear have a H I content consistent with the hypothesis that they are comprised of galaxies similar to those found in the field.
- In comparing the mass estimation methods of HG84 and H97, it was found that they all give similar results, although the scatter in the relationships prevents them from predicting the mass of an isolated galaxy accurately.

- The observations of the detected compact groups all have measured H I contents within the predicted range for their expected content. However, the majority of groups were not detected in this study, and so the H I content of the compact group population as a whole is not yet known.
- The survey has not been significantly affected by biasing due to interlopers lying far outside the group radius; nor from faint galaxies not listed as part of the compact groups, but contributing to the H I mass.
- The H I in the compact groups may have evolved through interactions, and star-formation, but it is clear that the fraction of H I consumed is not large. It is possible that H I depletion is correlated with morphological evolution towards earlier types, which would reduce the expected H I mass of the group.
- A survey of these groups with much higher sensitivity, leading to more detections or more stringent upper limits, will give a clearer picture on their H I content.

Acknowledgments

The authors would like to thank the HIPASS team for the data used for our detection experiment, and the staff at the Parkes radio telescope for assistance/troubleshooting during our observing runs. Thanks also to Donald Payne, Martin Zwaan, and Stacey Mader for observing assistance. Thanks to Martin Meyer for the use of MMSPECT, and for providing us with the HIPASS catalogue. J.S. thanks Melbourne University for financial assistance through a Melbourne University Postgraduate Scholarship. The Parkes telescope is part of the Australia Telescope, which is funded by the Commonwealth of Australia for operation as a National Facility managed by CSIRO. This work has made use of the MIRIAD software package. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

Accessory Materials

A complete display of the profile windows (see Figure 1) is available from the authors or, until 2009, from Publications of the Astronomical Society of Australia.

References

- Barnes, D. G. 1998, in ASP Conf. Ser., Vol. 145, *Astronomical Data Analysis Software and Systems VII*, eds. R. Albrecht, R. N. Hook, & H. A. Bushouse (San Francisco, CA: ASP), p. 32
- Barnes, D. G., Staveley-Smith, L., de Blok, W. J. G., Oosterloo, T., Stewart, I. M., Wright, A. E., Banks, G. D., Bhathal, R. et al. 2001, *MNRAS*, 322, 486
- Braun, R., & Burton, W. B. 1999, *A&A*, 341, 437
- Buckley, J., & James, I. 1979, *Biometrika*, 66, 429
- Chamaraux, P. 1987, *A&A*, 177, 326
- Chamaraux, P., Balkowski, C., & Fontanelli, P. 1986, *A&A*, 165, 15
- de Vaucouleurs, G., de Vaucouleurs, A., & Corwin, H. G. 1976, in *Second Reference Catalogue of Bright Galaxies* (Austin, TX: University of Texas Press)
- Haynes, M. P., & Giovanelli, R. 1984, *AJ*, 89, 758
- Hibbard, J. E., & Van Gorkom, J. H. 1996, *AJ*, 111, 655
- Hickson, P. 1982, *ApJ*, 255, 382
- Hickson, P. 1993, in *Atlas of Compact Groups of Galaxies* (New York, NY: Gordon & Breach)
- Hickson, P. 1997, *ARA&A*, 35, 357
- Huchtmeier, W. K. 1997, *A&A*, 325, 473
- Huchtmeier, W. K., & Richter, O. G. 1988, *A&A*, 203, 237
- Isobe, T., Feigelson, E. D., & Nelson, P. I. 1986, *ApJ*, 306, 490
- Karachentseva, V. E. 1973, *AISAO*, 8, 3
- Leon, S., Combes, F., & Menon, T. K. 1998, *A&A*, 330, 37
- Meyer, M., Zwaan, M., Webster, R. L., Staveley-Smith, L., Ryan-Weber, E., Drinkwater, M. J., Barnes, D. G., Howlett, M. et al. 2004, *MNRAS*, 350, 1195
- Nilson, P. 1973, in *Uppsala General Catalogue of Galaxies* (Uppsala: Astronomiska Observatorium)
- Ponman, T., Bourner, P. D. J., & Ebeling, H. 1996, *Roentgenstrahlung from the Universe*, 357
- Prandoni, I., Iovino, A., & MacGillivray, H. T. 1994, *AJ*, 107, 1235
- Sadler, E. M., Oosterloo, T., & Morganti, R. 2002, in *The Dynamics, Structure, & History of Galaxies* (San Francisco, SF: ASP), 215
- Sandage, A., & Tammann, G. A. 1981, in *Revised Shapley-Ames Catalog of Bright Galaxies* (Washington, DC: Carnegie Institute)
- Schneider, S. E., Helou, G., Salpeter, E. E., & Terzian, Y. 1986, *AJ*, 92, 742
- Solanes, J. M., Manrique, A., Garcia-Gomez, C., Gonzalez-Casado, G., Giovanelli, R., & Haynes, M. P. 2001, *ApJ*, 548, 97
- Verdes-Montenegro, L., Yun, M. S., Perea, J., del Olmo, A., & Ho, P. T. P. 1998, *ApJ*, 497, 89
- Verdes-Montenegro, L., Yun, M. S., Williams, B. A., Huchtmeier, W. K., del Olmo, A., & Perea, J. 2001, *A&A*, 377, 812
- Williams, B. A., & Rood, H. J. 1987, *ApJS*, 63, 265
- Zepf, S. E., de Carvalho, R. R., & Ribeiro, A. L. B. 1997, *ApJL*, 488, 11