

High entropy and evidence for a merger in the low surface brightness cluster Abell 2399

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Abstract

We present results of the X-ray analyses of the nearby cluster of galaxies Abell 2399 ($z = 0.058$) using the XMM–Newton and Suzaku satellites. This cluster is classified as a low surface brightness (LSB) cluster. It has a bimodal structure in X-rays, and the X-ray-emission peaks are significantly offset from the peaks in gas temperature. By de-projecting the annular spectra, the temperature and electron density in the central $r < 0.1 r_{200}$ region are measured to be 3.6 keV and $1.4 \times 10^{-3} \text{ cm}^{-3}$, respectively. This yields a very high gas entropy, $\sim 300 \text{ keV cm}^2$, in the central region, comparable to the values observed in other LSB clusters. The scaled entropy of Abell 2399 is the highest among the REXCESS cluster sample. The spatial distribution of the member galaxies exhibits multiple substructures, the locations of which are significantly different from those in the X-ray image. We discovered a clear discontinuity in the X-ray brightness and temperature distributions in the western gas clump, which provides evidence of the presence of a cold front. Therefore, our results strongly suggest that A2399 has experienced a merger and that the high central entropy originates from the merger activity.

Key words: galaxies: clusters: individual (Abell 2399) — galaxies: clusters: intracluster medium — X-rays: galaxies: clusters

1 Introduction

A low surface brightness cluster of galaxies (hereafter an LSB cluster) is a type of cluster that has an extremely low X-ray surface brightness and a highly irregular morphology. Several LSB clusters have been identified in the

REXCESS cluster sample, which make up 5%–10% of the cluster sample detected by the ROSAT All-Sky Survey (Böhringer et al. 2013, 2017). One requires deep observations to investigate the detailed gas properties of such clusters because of their extremely low brightness (typically $\sim 10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ arcmin}^{-2}$ in the 0.1–2.4 keV band) and

diffuse gas distribution without a prominent core. Hence their evolution remains a puzzle. Since upcoming cluster surveys—by, e.g., eROSITA on the SRG satellite—are expected to detect a large number of clusters including faint ones (Merloni et al. 2012), it is important to understand the properties of LSB clusters, which may not follow the mass-scaling relations of bright clusters.

The Suzaku satellite (Mitsuda et al. 2007) has been used to analyze a few LSB clusters in detail. A76 is an archetypical LSB cluster; it has the lowest surface brightness among the ROSAT clusters studied by Neumann and Arnaud (1999) and an irregular morphology. Ota et al. (2013) reported that A76 has an exceptionally high gas entropy in the central region, $K \equiv kTn_e^{-2/3} \sim 400 \text{ keV cm}^2$, and a low electron density, of the order of 10^{-4} – 10^{-3} cm^{-3} , for the observed mean gas temperature of 3 keV. This result is not explained either by gravitational heating or by preheating. A similar trend has been found in other LSB clusters, such as A548W (Nakazawa et al. 2016) and A1631 (Babazaki et al. 2018). Babazaki et al. (2018) noted that the central entropy of A1631 is one of the highest ($\gtrsim 400 \text{ keV cm}^2$) among known nearby clusters, while the X-ray luminosity is lower by a factor of 3 than that expected from the luminosity–temperature relation for relaxed clusters. In addition, they found that the spatial distribution of the gas is significantly different from that of the member galaxies and they proposed a post-merger scenario to explain the production of high-entropy gas in this irregular cluster.

To clarify the nature of LSB clusters, we need to increase the number of clusters in the sample. In this paper, we conduct an X-ray imaging and spectroscopic study of the LSB cluster Abell 2399 at $z = 0.058$ using the XMM–Newton and Suzaku satellites. The good spatial resolution of XMM–Newton enables us to derive the temperature distribution of the hot gas associated with substructures in the cluster. In addition, we have searched for discontinuities in the X-ray surface brightness as an indicator of the merger activity. Suzaku’s low-background level allows us to measure the gas-entropy profile out to the virial radius. We have also compared the spatial distributions of the galaxies and the hot gas in this system because rich spectroscopic data in optical are available for the member galaxies (Moretti et al. 2017). In previous studies, A2399 (the alternative name is RXC J2157–0747) has been analyzed within the REXCESS sample (Böhringer et al. 2007); the profiles of gas density and entropy were determined up to $\sim r_{500}$ by Croston et al. (2008) and Pratt et al. (2010), and the morphological investigation was also studied by Böhringer et al. (2010). They noted that RXC J2157–0747 has a higher entropy and is one of outliers in the substructure characterization, such as a power ratio and a centroid shift parameter among the

REXCESS sample, even though they did not pursue the origin of the specific properties.

This paper is organized as follows: In section 2 we present the X-ray observations of A2399 with Suzaku and XMM–Newton and the data reduction. Sections 3 and 4 describe the results from our imaging and spectral analyses. In section 5, we discuss the X-ray and optical properties and propose a scenario to interpret the data. Where necessary, we assume the standard cosmological model, with a matter density $\Omega_M = 0.27$, a cosmological constant $\Omega_\Lambda = 0.73$, and a Hubble constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$. At the cluster redshift $z = 0.058$, $1'$ corresponds to 68 kpc. In this paper, we have employed HEASoft v6.21 and XSPEC version 12.9 to perform χ^2 fitting of the X-ray spectra, and we have used the metal-abundance table of Anders and Grevesse (1989). Unless otherwise stated, the error ranges show the 90% confidence levels from the central values.

2 Observation and data reduction

2.1 Suzaku

In 2014 November, four-pointed Suzaku observations of A2399 (the center, east, west, and north) were performed. Table 1 summarizes the basic information about the observations. Four X-ray sensitive CCD cameras—known as the XIS instruments—are installed on Suzaku. Three of the CCD cameras are front-illuminated (XIS-0, -2, -3), and the other is back-illuminated (XIS-1) (Koyama et al. 2007). During the observation periods, the XIS-0, -1, and -3 CCD cameras were operated in normal mode with space charge injection enabled (Uchiyama et al. 2009).

We used standard Suzaku pipeline processing to create event files with CALDB [v2016-06-07 for XIS and v2011-06-30 for the X-ray telescopes (XRTs: Serlemitsos et al. 2007)]. We used the following criteria to screen the data: Earth elevation angle $> 5^\circ$, day-Earth elevation angle $> 20^\circ$, and time removal during the South Atlantic Anomaly. As an additional correction, we removed flickering pixels.¹ We also excluded point sources detected by XMM–Newton (see subsection 2.2) from this Suzaku analysis.

We determined the X-ray background from the XIS spectra in the northern field, $\sim 1.3 \text{ Mpc}$ away from the center (figure 1a). We subtracted the detector background and evaluated other background components assuming the same models as used in, e.g., Mitsuishi et al. (2012) and Babazaki et al. (2018), which consists of the cosmic X-ray background (CXB) and the Galactic X-ray background (GXB). Table 2 summarizes the best-fitting parameters, and they are consistent with typical values for the CXB and GXB components (e.g., Kushino et al. 2002; Yoshino et al. 2009).

¹ (<https://heasarc.gsfc.nasa.gov/docs/suzaku/analysis/xisnxbnew.html>).

Table 1. Logs of observations of A2399 with Suzaku and XMM–Newton.

Target	Obs ID	Date	RA (J2000.0) [$^{\circ}$]	Dec (J2000.0) [$^{\circ}$]	Net exposure [ks]
Suzaku/XIS					
A2399 West	809020010	2014-11-11	329.1831	−7.8643	23.1
A2399 Center	809022010	2014-11-12	329.2851	−7.5651	26.0
A2399 North	809021010	2014-11-12	329.3708	−7.7975	18.6
A2399 East	809023010	2014-11-15	329.5573	−7.7299	20.9
XMM–Newton/EPIC					
RXC J2157–0747 (A2399)	0654440101	2010-06-07	329.3573	−7.7987	54.9* (MOS1), 59.5* (MOS2), 26.2 (PN)

*Net exposure time after removal of periods of high background flaring.

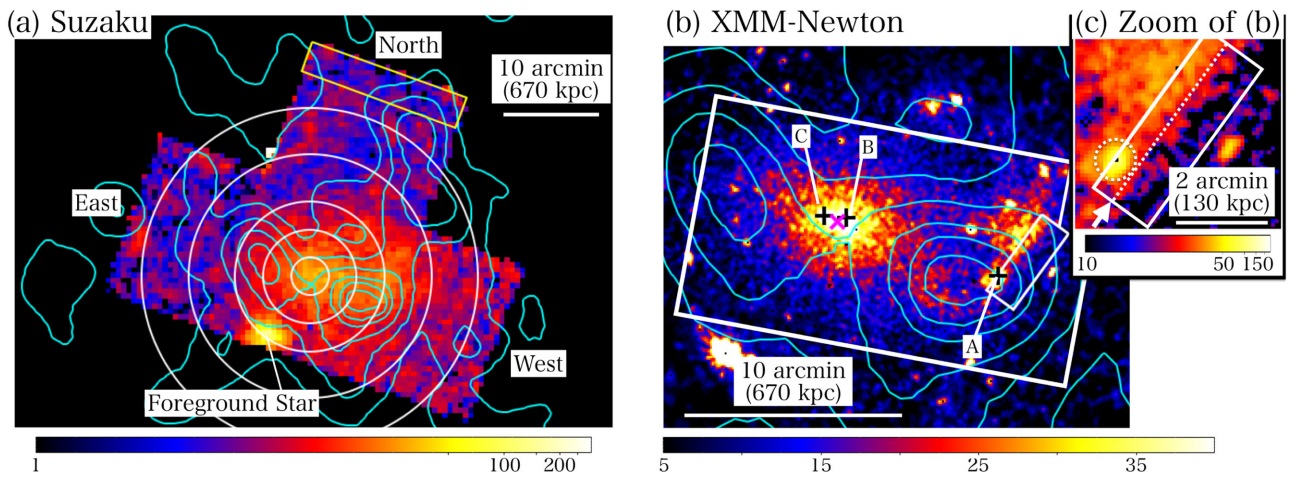


Fig. 1. Suzaku and XMM–Newton images in the 0.5–5 keV band. Solid white rings and rectangles and cyan contours indicate the regions used in the imaging and spectral analysis and the galaxy number density distribution, respectively. (a) Suzaku XIS-1 mosaic X-ray image of A2399 in units of $\text{cts} (20 \text{ ks})^{-1} (1024 \text{ pixels})^{-1}$. We subtracted the NXB component. We corrected for exposure and vignetting effect and smoothed the image using a Gaussian function with $\sigma = 25'$. We used the rectangular region outlined in yellow to estimate the X-ray background emission. A bright foreground star, SDSS J215751.40–075348.1, is shown and excluded in the spectral analysis. (b) XMM–Newton EPIC composite images in units of $\text{cts s}^{-1} \text{ degree}^{-2}$. The QBP background has been subtracted. The positions of the first, second, and third brightest galaxies are denoted by A, B, and C with black cross marks. The magenta cross shows the X-ray peak, which is derived from the XMM–Newton image after excluding point sources. (c) A close-up view of (b) to emphasize an extracted area (a white rectangle) for the cold front study. The arrow and the dotted line show the plane of the discontinuity in the surface brightness. The dotted white circle corresponds to a point source that was excluded in the analysis. (Color online)

Thus, we incorporate this background model in the spectral analysis.

2.2 XMM–Newton

We retrieved three data sets containing A2399 (RXC J2157–0747) from the XMM–Newton Science Archive. We used only one of these data sets, OBSID 0654440101, in the analysis because the other two are heavily contaminated by soft-proton flares.

We performed data reduction for the set of three X-ray CCD cameras onboard XMM–Newton—the European Photon Imaging Camera (EPIC)—in the standard manner, using the ESAS (Extended Source Analysis Software) package (Snowden et al. 2008) in SAS version 16.0.0. We removed high background periods, for which the rates were beyond the 2σ range of the rate distribution. Table 1 shows the observation identifications and the net exposure times after filtering. We used an ESAS routine to detect point sources in each detector and removed them from our analysis.

3 X-ray images

Figures 1a and 1b show Suzaku/XIS-1 and XMM-Newton/EPIC images of A2399 in the 0.5–5.0 keV band, respectively. In the XIS image, diffuse cluster emission is detected out to the virial radius $r_{200} = 1.2$ Mpc or $17'$. Here the virial radius is estimated from the global temperature $\langle kT \rangle = 2.6$ keV and the T - r_{200} relation (Arnaud et al. 2005) which is obtained from clusters in the temperature range of 2–9 keV.

In the EPIC image, the cluster shows a bimodal structure; the main gas clump in the east and a smaller (“the sub”) gas clump in the west each have a projected distance of 500 kpc. The X-ray surface brightness of the main clump has a peak (329.3790, -7.7982) after removal of point-like sources, and it appears to be extended in the southwest–northeast direction. On the other hand, the sub clump has a rather elongated shape, extending in the southeast–northwest direction (see also Böhringer et al. 2007).

The bright X-ray point-like sources located around the center of the main/sub clumps correspond to the three brightest cluster galaxies (BCGs; Smith et al. 2004; Cava et al. 2009; Varela et al. 2009; Moretti et al. 2017) around redshift of A2399, i.e., WINGS J215701.72–075022.0, WINGS J215729.43–074744.1, and WINGS J215733.47–074738.8, and they are denoted by A, B, and C, respectively, in figure 1b.

4 X-ray spectra

To investigate the spatial distributions of the physical quantities that characterize the cluster gas, we performed X-ray spectral analyses using the XMM-Newton and Suzaku data. We define two kinds of spectral-integration regions in subsection 4.1, and we present the results of our analysis in subsections 4.2 and 4.3.

4.1 Definitions of spectral regions

As figure 1 indicates, we extracted spectra from the following two kinds of regions:

- (i) Bridge regions along the gas clumps (12 boxes; $1.5 \times 10' \times 12 \simeq 100 \text{ kpc} \times 700 \text{ kpc} \times 12$).
- (ii) Annular regions centered on the X-ray peak of the main clump after removal of point sources (section 3). Considering the PSF size of the Suzaku XRT, five radial bins are chosen as follows: $0' < r < 2'$, $2' < r < 5'$, $5' < r < 8'$, $8' < r < 13'$, and $13' < r < 18'$, which correspond to $0 < r < 140 \text{ kpc}$, $140 < r < 340 \text{ kpc}$, $340 < r < 540 \text{ kpc}$, $540 < r < 880 \text{ kpc}$, and $880 < r < 1200 \text{ kpc}$. Note that the first three rings are extracted from the “A2399 Center” field and thus cover a position angle 0° – 360° , while the

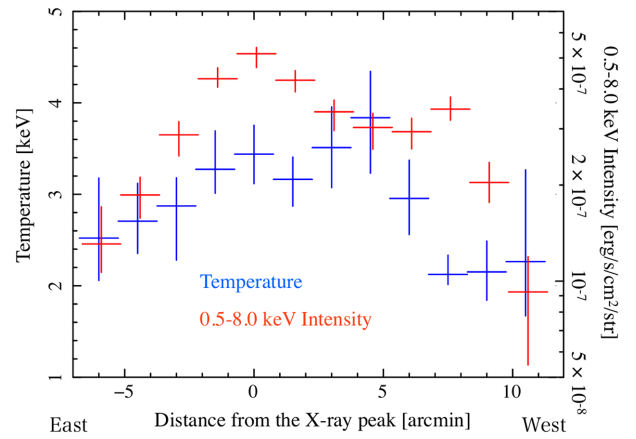


Fig. 2. Temperature (blue) and intensity (red) profiles across the main and sub gas clumps. The horizontal axis shows the distance from the X-ray peak (negative is east, and positive is west; see figure 1b). (Color online)

outer two rings in the east, north, and west directions are extracted from the “A2399 East,” “A2399 North,” and “A2399 West” fields of view, respectively, and thus each region has a sector angle of about 90° .

For (i), we used only the XMM-Newton data, to take advantage of its higher angular resolution and to avoid contamination from the point sources (subsection 4.2). For (ii), we used the Suzaku data to capitalize on its stable and low background characteristics, which enable us to study low surface brightness signals out to the virial radius for the first time (subsection 4.3).

4.2 Bridge regions along the gas clumps

To investigate the temperature and intensity profiles along the axis connecting the main and sub gas clumps, we extracted XMM-Newton spectra from the 12 boxes defined in region (i) of the previous subsection (see also figure 1b). We fitted the 0.5–11.0 keV spectra to the sum of the X-ray background, the cluster, and a soft proton background model, in the same manner as in Babazaki et al. (2018). For the X-ray background emission, we took into account the local hot bubble (LHB), the Galactic halo (GH), and the CXB components. The parameters of the X-ray background components are limited to the error ranges given in table 2. The power-law function representing the soft-proton contamination was included with the diagonal response files distributed by the XMM-Newton team. The photon indices for the MOS1 and MOS2 spectra are linked to each other, as indicated in the guidelines given in the ESAS manual.² The Galactic column density N_H is fixed at $3.0 \times 10^{20} \text{ cm}^{-2}$ based on the H I maps (Kalberla et al.

² (<https://heasarc.gsfc.nasa.gov/docs/xmm/esas/cookbook/xmm-esas.html>).

Table 2. Best-fitting parameters for the X-ray background emission.

kT_{LHB} (keV)	$Norm_{\text{LHB}}^*$	kT_{GH} (keV)	$Norm_{\text{GH}}^*$	Γ_{CXB}	$Norm_{\text{CXB}}^\dagger$	$\chi^2/\text{d.o.f.}$
0.1 (fixed)	$11.3^{+4.3}_{-7.8}$	$0.30^{+0.50}_{-0.12}$	$5.4^{+1.9}_{-3.5}$	1.4 (fixed)	$10.4^{+1.7}_{-1.5}$	72/62

*Normalization of the APEC model divided by a solid angle Ω (sr). $Norm = (10^{-14}/\Omega) \int n_e n_H dV / \{4[(1+z)D_A]^2\}$ in units of $\text{cm}^{-5} \text{sr}^{-1}$, where n_e , n_H , D , and V are the electron and hydrogen number densities (cm^{-3}), the angular diameter distance (cm) and the emission volume (cm^3), respectively. Here, “LHB” denotes the local hot bubble, and “GH” the Galactic halo, respectively.

† The unit is photon $\text{keV}^{-1} \text{cm}^{-2} \text{s}^{-1} \text{sr}^{-1}$ at 1 keV.

Table 3. Best-fitting parameters for the bridge region between the main and sub gas clumps.

Distance* (')	kT (keV)	$Norm^\dagger$	$\chi^2/\text{d.o.f.}$
−6.0	$2.14^{+0.71}_{-0.42}$	89^{+25}_{-24}	300/323
−4.5	$2.52^{+0.66}_{-0.46}$	146^{+45}_{-27}	363/358
−3.0	$2.71^{+0.42}_{-0.35}$	200^{+27}_{-31}	437/440
−1.5	$2.87^{+0.31}_{-0.59}$	299^{+30}_{-42}	574/541
0	$3.27^{+0.42}_{-0.26}$	421^{+35}_{-25}	564/606
+1.5	$3.44^{+0.32}_{-0.32}$	492^{+23}_{-47}	673/710
+3.0	$3.16^{+0.24}_{-0.29}$	424^{+30}_{-34}	732/677
+4.5	$3.51^{+0.45}_{-0.44}$	321^{+29}_{-41}	529/550
+6.0	$3.84^{+0.51}_{-0.61}$	276^{+30}_{-40}	581/542
+7.5	$2.95^{+0.42}_{-0.40}$	302^{+31}_{-35}	437/462
+9.0	$2.12^{+0.21}_{-0.11}$	419^{+38}_{-32}	524/491
+10.5	$2.15^{+0.34}_{-0.31}$	245^{+38}_{-33}	489/447

*The distance from the peak position in the main gas clump (minus: east, plus: west).

† See footnote * in table 2.

2005). In the model for the hot cluster gas, the redshift is fixed at 0.0579, and the metal abundance is fixed at 0.22 solar, as obtained from the spectral analysis of the entire region.

Figure 2 shows the temperature and intensity profiles, and table 3 summarizes the data. The intensity profile shows two clear peaks, at 0' and 7.5', corresponding to the positions of the main and sub gas clumps. In the bridge region between the two gas clumps, the intracluster medium (ICM) shows a relatively flat temperature profile, with $kT \sim 3$ keV. The temperature decreases to ~ 2 keV outside the bridge, and a steep gradient is seen, particularly around the sub gas clump. There are no prominent temperature peaks at the positions of the X-ray-intensity peaks, indicating a significant difference between the temperature and gas-density profiles.

4.3 Annular regions

To derive the three-dimensional structure of the temperature, density, and entropy of the hot gas, we de-projected the data, assuming a spherically symmetric gas distribution (figure 1a). We derived the distributions separately for each of the three directions (east, north, and west) to study the spatial variations. In each direction, we fitted the five annular spectra in the 0.7–8 keV band simultaneously and the de-projection operation by using the “project” model in XSPEC. We assumed the APEC model for cluster emission attenuated by Galactic absorption for each annular bin. We assumed the metal abundance to be constant, independent of radius, and fixed at 0.22 solar, as is also the case in subsection 4.2. Because the temperature was not well constrained in the outermost annular regions (13'–18') in the east and north directions, we linked the temperature of that region to the next-innermost annular region (8'–13'). The fit was reasonable, giving $\chi^2/\text{degrees of freedom (d.o.f.)} = 408/382$, $484/414$, and $498/438$ for the east/north/west directions.

Table 4 lists the derived temperature values, and figure 3 a plots their profiles. The temperature is about 3.6 keV at the center and drops by a factor of 2 in the outskirts. In the ring $0.30 < r/r_{200} < 0.48$, the temperature in the west, $4.1^{+1.7}_{-1.1}$ keV, is statistically consistent with the XMM result (figure 2), and it is marginally higher than in the other two directions.

We calculated the electron density n_e from the normalization of the APEC model. The central density is $1.4 \times 10^{-3} \text{cm}^{-3}$, which is extremely low in comparison with other nearby clusters. Here the density is averaged within $0.1 r_{200}$ because we used the radial bin-size of 2', considering the spatial resolution of Suzaku, and assumed that the gas is uniform in each spherical shell. At a very small radius, however, the gas density can be higher; in fact, the previous XMM–Newton measurement yielded $3.5 \times 10^{-3} \text{cm}^{-3}$ at $r = 4.7 \text{kpc}$ or $0.004 r_{200}$ (Croston et al. 2008). For A2399 West, a density excess is seen in the ring $0.48 < r/r_{200} < 0.76$, which can be attributed to the western sub gas clump

Table 4. Results of the de-projection analysis for A2399 East, North, and West.

Region	A2399 East			A2399 North			A2399 West		
	kT (keV)	n_{e0} (10^{-4} cm^{-3})	K (keV cm^2)	kT (keV)	n_{e0} (10^{-4} cm^{-3})	K (keV cm^2)	kT (keV)	n_{e0} (10^{-4} cm^{-3})	K (keV cm^2)
0'–2'	$3.58^{+0.68}_{-0.50}$	13.73 ± 0.53	289 ± 49	$3.60^{+0.69}_{-0.51}$	13.70 ± 0.53	292 ± 49	$3.68^{+0.74}_{-0.55}$	13.73 ± 0.54	302 ± 53
2'–5'	$4.84^{+1.09}_{-0.88}$	5.00 ± 0.19	769 ± 157	$4.67^{+1.18}_{-0.80}$	5.02 ± 0.20	739 ± 158	$4.28^{+0.87}_{-0.65}$	5.29 ± 0.19	655 ± 117
5'–8'	$2.62^{+0.31}_{-0.21}$	4.71 ± 0.12	434 ± 44	$2.78^{+0.33}_{-0.30}$	4.52 ± 0.14	472 ± 54	$4.11^{+1.67}_{-1.07}$	2.91 ± 0.20	937 ± 315
8'–13'	$1.59^{+0.50}_{-0.25}$	1.04 ± 0.19	719 ± 190	$1.63^{+0.43}_{-0.14}$	1.15 ± 0.19	688 ± 142	$2.16^{+0.25}_{-0.15}$	2.88 ± 0.09	495 ± 48
13'–18'	\uparrow^*	0.84 ± 0.13	828 ± 212	\uparrow^*	1.25 ± 0.09	651 ± 118	$1.24^{+0.09}_{-0.11}$	1.28 ± 0.09	487 ± 46

*The upward-pointing arrow means that the parameter is linked to that of the inner region.

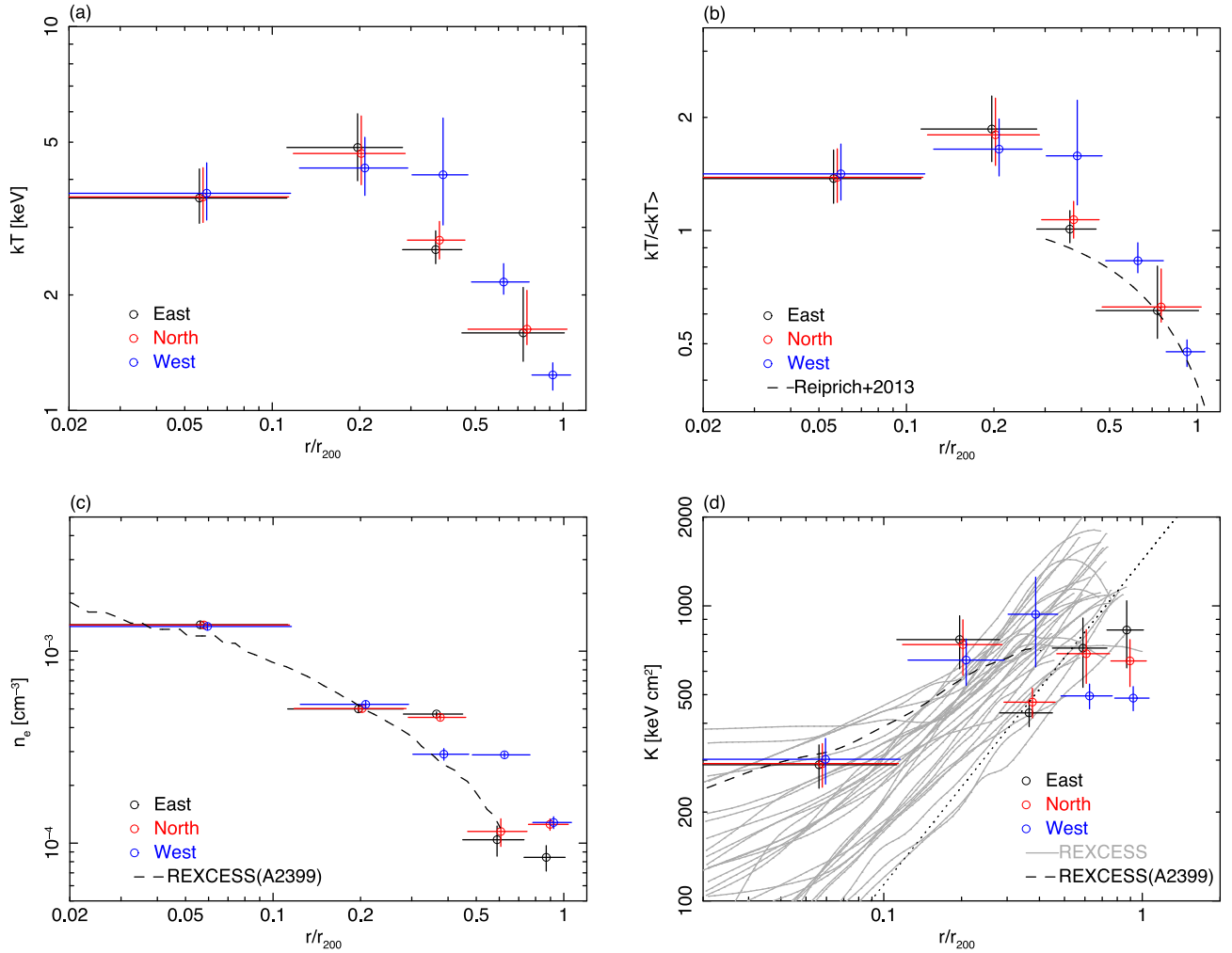


Fig. 3. (a) Temperature, (b) normalized temperature, (c) gas density, and (d) entropy profiles of A2399 derived from de-projection of the Suzaku spectra. The black, red, and blue symbols show the profiles in the east, north, and west directions, respectively. The markers for the north (west) direction are shifted horizontally by 3(6)% for clarity. In panel (b), the temperature is normalized relative to the mean value (kT) = 2.6 keV. The dashed line indicates the best-fitting temperature profile obtained from the Suzaku observations of relaxed clusters (Reiprich et al. 2013). In panel (c), the gas density derived by Croston et al. (2008) is indicated by the dashed line. In panel (d), the entropy profiles of the REXCESS sample (Pratt et al. 2010) are also shown for comparison and the data for A2399 is denoted by the dashed line. The dotted line indicates the baseline entropy profile (Voit et al. 2005) calculated for the mean temperature of A2399. The radius is normalized by r_{200} . (Color online)

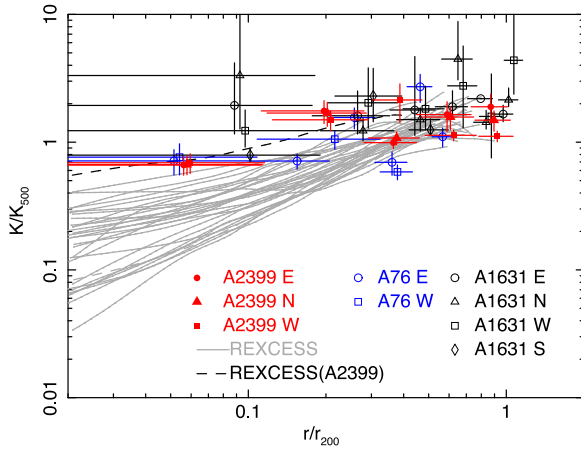


Fig. 4. Scaled entropy profile of A2399 in the east (red filled circle), north (red filled triangle), and west (red filled square) directions, as derived from the Suzaku XIS analysis. For comparison, the profiles of A76 in the east (blue open circle) and west (blue open square) and A1631 in the east (black open circle), north (black open triangle), west (black open square), and south (black open diamond) obtained from previous Suzaku observations are shown (Ota et al. 2013; Babazaki et al. 2018). The profiles of the REXCESS sample are shown by the gray lines, and A2399 (RXC J2157–0747) is plotted by the dashed black line. (Color online)

(figure 1b). Because of the presence of this high-density clump, the de-projection analysis leads to a relatively lower density bin in the ring $0.30 < r/r_{200} < 0.48$ in comparison with the other directions, even though the surface brightness is apparently the same.

We evaluated the entropy $K = kTn_e^{-2/3}$ from the temperature and gas density within r_{200} . The entropy is very high; the azimuthally-averaged value is $\sim 290 \pm 30 \text{ keV cm}^2$ in the central $r < 0.1 r_{200}$ region, and flat outside $0.1 r_{200}$. The entropy for $r/r_{200} > 0.48$ in the west is marginally lower than in the other directions, which again coincides with the location of the sub clump. In comparison with the previous XMM–Newton study of the REXCESS sample, the Suzaku result for $r < 0.36 r_{200}$ agrees with their result within the measurement uncertainties. For $r \sim 0.36 r_{200}$, there are significant azimuthal variations due to the clumpy structure and the difference in the entropy values can be attributed to the difference in their analysis methods because in the XMM–Newton deprojection analysis by Pratt et al. (2010), the western gas clump is excluded and the maximum radius is about a half of that used in the Suzaku analysis.

5 Discussion

In this section, we discuss the thermodynamic properties of A2399 based on the XMM–Newton and Suzaku results. Then, we compare the X-ray and optical observations and discuss the dynamical state of the system.

5.1 Thermodynamic properties of the cluster gas

Previous X-ray observations have shown that the temperature profiles of relaxed clusters exhibit similarities at large radii, once they are scaled by the mean temperature and the virial radius (e.g., Vikhlinin et al. 2005; Pratt et al. 2007; Reiprich et al. 2013; Walker et al. 2013; Ichikawa et al. 2013; Ghirardini et al. 2018). Figure 3b compares the normalized temperature profile of A2399 and the average of relaxed clusters observed with Suzaku (Reiprich et al. 2013). For $r > 0.3 r_{200}$, A2399 shows a decline similar to that of a relaxed cluster, while we found no significant cool-core structure for $r < 0.3 r_{200}$.

Figure 4 compares the A2399 scaled-entropy profile, as derived from our Suzaku analysis in subsection 4.3, with those of the REXCESS sample derived from the XMM–Newton observations (Pratt et al. 2010). Here the Suzaku data of A2399 are scaled by the value of K_{500} estimated from M_{500} ($\sim 1.3 h_{70}^{-1} \times 10^{14} M_{\odot}$) obtained by using the M_{500} – Y_X scaling relation in table 1 and equation (3) (Pratt et al. 2010). In the central region, A2399 has the highest scaled entropy, $K/K_{500} \sim 0.7$, in comparison with other REXCESS clusters as suggested by the previous study, and the present Suzaku result agrees well with that from XMM–Newton. We confirmed the difference in M_{500} is only $\sim 20\%$ corresponding to $\sim 10\%$ in K_{500} even though we adopt the M_{500} – T scaling relation (Arnaud et al. 2005) to obtain M_{500} . At larger radii ($r \gtrsim 0.3 r_{200}$) up to around the virial radius, even in the east and north directions, the normalized entropy values seem to be systematically lower than those of relaxed clusters (Walker et al. 2012) indicating that a large-scale event gives rise to the discrepancy in the whole cluster system. The scaled-entropy profiles of other two LSB clusters, A76 and A1631, obtained from previous Suzaku studies (Ota et al. 2013; Babazaki et al. 2018) are also shown in figure 4. Note that these two clusters are not a part of the REXCESS sample. The central scaled-entropy values of A2399 and A76 agree well with each other, while they tend to be lower than those of A1631. This mainly reflects a difference in electron density; A1631 has the lowest density of a few 10^{-4} cm^{-3} among known nearby clusters. The uncertainty in the scaling factor may also affect the difference.

From numerical simulations that include gravitational heating effect, Voit, Kay, and Bryan (2005) derived a “baseline” entropy model for the cluster gas; $K = 550 \text{ keV cm}^2 (kT/1 \text{ keV}) (r/r_{200})^{1.1}$. Figure 3d compares the observed A2399 profile with the baseline model calculated for the averaged temperature of $kT = 2.6 \text{ keV}$. The azimuthally-averaged central entropy, $294 \pm 29 \text{ keV cm}^2$, is significantly higher than that expected from gravitational heating alone ($\sim 100 \text{ keV cm}^2$ at $0.1 r_{200}$).

Non-gravitational heating processes can also contribute to the entropy excess, and numerical simulations suggest that active galactic nuclei (AGN) feedback plays a major role (e.g., Dubois et al. 2011). Although no signs of cavities or jets have been reported in A2399, we can assess the impact of AGN feedback as discussed in previous studies (Ota et al. 2013; Babazaki et al. 2018). Using the *K*-band luminosity of the brightest galaxy 2MASX J21572939–0747443, $\log(L_K/L_{K,\odot}) = 11.5$ (Skrutskie et al. 2006), and the empirical relationship given by Wang et al. 2010, we estimate the excess entropy in A2399 to be of the order of 1 keV cm^2 . Therefore, we conclude that the observed high entropy also cannot be explained by AGN feedback. Another likely origin of the excess entropy is disturbance by a cluster merger, which we will discuss in the following subsections.

5.2 Comparison between X-ray and optical properties

The galaxy distribution provides a clue as to the dynamical state of a cluster. For instance, if a cluster merger occurs, substructures can be identified in both spatial and velocity distributions of member galaxies (Geller & Beers 1982). Such substructures are detected in the adaptively smoothed galaxy maps of many merging systems (e.g., Owers et al. 2017).

The Omega WINGS survey (Moretti et al. 2017) has identified 234 spectroscopic galaxies as member galaxies of A2399. The smoothed galaxy map (see contours in figure 1) shows several substructures, with the peak number density in the western and eastern galaxy clumps being ~ 90 and 70 Mpc^{-2} , respectively. Thus, the spatial offset between the gas and galaxy clumps is $> 100 \text{ kpc}$. We have measured the central redshift and velocity dispersion of the galaxies that constitute the main and sub galaxy clumps. There is no significant difference in central redshift between the galaxy clumps, while the observed velocity dispersion is ~ 800 and $\sim 600 \text{ km s}^{-1}$ in the main and sub galaxy clumps, respectively. The total mass is estimated to be ~ 9 and $\sim 3 \times 10^{14} M_\odot$, respectively (Finn et al. 2005).

5.3 X-ray surface brightness discontinuity on the rim of the sub gas clump

To constrain the dynamical state of the system, we searched for discontinuities in the surface brightness distribution. As seen from figures 1c and 5, there is a clear surface-brightness discontinuity near the rim of the western sub gas clump. Therefore we modeled the surface brightness by integrating the following double power-law model for the electron-density profile along the line of sight (e.g., Owers et al.

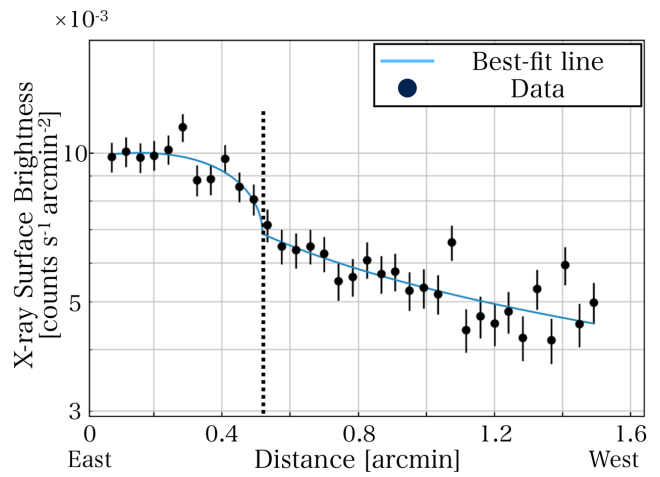


Fig. 5. Surface brightness profile around the rim of the sub gas clump as shown in the white rectangular region in figure 1c. The X-axis labeled “Distance” is defined from the east to the west direction in figure 1. The blue line indicates the best-fitting surface-brightness profile model that is calculated from the double power-law model [equation (1)]. The error bars indicate the 1σ statistical uncertainties. The dotted line corresponds to the position of the discontinuity. (Color online)

2009):

$$n_e(r) = n_0 \left(\frac{r}{r_s} \right)^{-\alpha_1} \quad (r < r_s), \quad \frac{n_0}{j_{12}} \left(\frac{r}{r_s} \right)^{-\alpha_2} \quad (r \geq r_s). \quad (1)$$

Here we assume spherical symmetry for simplicity, and n_0 , r_s , and j_{12} represent the central density, the radius at the point of discontinuity, and the density-jump parameter, respectively. The quantities α_1 and α_2 are the power-law indices for the high-density and low-density regions, respectively.

A case with no density jump, that is, the model with the jump parameter fixed at 1, does not reproduce the observed profile well. However, when the jump parameter is set free, the fit improves significantly (the null hypothesis probability is 1.2%, according to the F-test). The best-fitting parameter is $j_{12} = 1.6 \pm 0.2$.

The XMM spectral analysis demonstrated that the temperatures of the upstream and downstream regions across the plane of the discontinuity are $1.83^{+0.26}_{-0.17} \text{ keV}$ and $3.71^{+1.27}_{-1.04} \text{ keV}$, respectively. We find that pressure equilibrium holds at the boundary of the density jump to within the error bars. Thus the density discontinuity originates from a cold front rather than a shock.

5.4 A possible scenario

As is the case for A1631 (Babazaki et al. 2018), the observational characteristics in both X-ray and optical—i.e., no prominent cool-core structure, a flatter density profile, excess entropy at the center, and a spatial offset between the

peaks in the X-ray surface brightness and galaxy number density distributions—suggest that the A2399 system is also in some stage of merging.

By utilizing hydrodynamic simulations, ZuHone (2011) has studied the effect of ICM mixing as a function of a mass ratio and impact parameter of a cluster collision. The simulations reproduce a high-entropy core—where the initial core is replaced by a low-density, high-temperature and high-entropy core—and in which the temperature and gas-density profiles become flatter, relative to those of the initial profiles. Therefore, we consider that the excess entropy in A2399 can be attributed to ICM mixing during a cluster merger.

Underluminosity in X-rays at the given velocity dispersion $\sigma = 729 \pm 35 \text{ km s}^{-1}$ (Moretti et al. 2017) extracted by using the whole member galaxies is also seen, with the observed luminosity being approximately 1/3 that expected from the L – σ relation (Zhang et al. 2011) predicted from the hydrodynamic simulations (e.g., ZuHone 2011). Tovmassian (2015) argues that some merging clusters show a higher (>1.2) or smaller (<0.8) ratio of n_l/n_h , where n_l and n_h correspond to the numbers of galaxies with velocities lower and higher than the mean velocity of all member galaxies, and the observed ratio for A2399 is 0.7 ± 0.1 . In the case of A2399, the significant substructures seen in both X-ray and optical data and the discrepancy in the X-ray peak positions between the temperature and surface brightness also support the merger scenario (e.g., Markevitch et al. 2004; Harvey et al. 2015).

Furthermore, the main gas clump seen in X-rays is located between the main and sub galaxy clumps observed in the optical, and the cold front is found around the rim of the gas clump. The X-ray morphology is elongated along an axis connecting the galaxy clumps, and there is little difference in redshift between the two galaxy clumps (see subsection 5.2), indicating that the cluster–cluster collision occurred nearly in the plane of the sky.

Thus, we conclude that A2399 may be a system in a head-on collision, like the Bullet cluster (e.g., Markevitch & Vikhlinin 2007). Deep X-ray observations and weak lensing will enable the determination of further details, such as shock structures and dark matter distributions. Because this LSB cluster resides in a supercluster, the merger events may have resulted from its dense environment.

6 Summary

To study the nature of low surface brightness clusters, we have analyzed the XMM–Newton and Suzaku data of an LSB cluster, A2399 at $z = 0.058$. The X-ray image is elongated and has a bimodal structure, while the temperature

peak is significantly offset from the X-ray emission peaks. We have derived the profiles of gas temperature, density, and entropy out to the viral radius. The observed high entropy in the central region, $\sim 300 \text{ keV cm}^2$, is comparable to that of other LSB clusters like A76 and A1631, and the scaled entropy of A2399 is the highest among the entire REXCESS sample. The galaxy map also exhibits main and sub clumps; however, there is a considerable offset between the peak positions in the main galaxy clump and the main gas clump. Furthermore, we have detected the brightness and temperature discontinuities near the sub galaxy clump in the west, providing evidence for the presence of a cold front. Therefore, our results support the scenario that this cluster has experienced a merger in the plane of the sky, and the excess central entropy and complex X-ray and optical morphologies are caused by the merger. Since the X-ray properties of LSB clusters tend to deviate from the scaling relations of known nearby clusters, we expect the present study to have a significant impact on the detection techniques and interpretations of faint, diffuse clusters in future cluster surveys.

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